

DETECTION OF GALACTIC CENTER SOURCE G2 AT 3.8 μm DURING PERIAPSE PASSAGE

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ABSTRACT

We report new observations of the Galactic Center source G2 from the W. M. Keck Observatory. G2 is a dusty red object associated with gas that shows tidal interactions as it nears its closest approach with the Galaxy's central black hole. Our observations, conducted as G2 passed through periaapse, were designed to test the proposal that G2 is a 3 Earth mass gas cloud. Such a cloud should be tidally disrupted during periaapse passage. The data were obtained using the Keck II laser guide star adaptive optics system (LGSAO) and the facility near-infrared camera (NIRC2) through the K' [2.1 μm] and L' [3.8 μm] broadband filters. Several results emerge from these observations: (1) G2 has survived its closest approach to the black hole as a compact, unresolved source at L' , (2) G2's L' brightness measurements are consistent with those over the last decade, (3) G2's motion continues to be consistent with a Keplerian model. These results rule out G2 as a pure gas cloud and imply that G2 has a central star. This star has a luminosity of $\sim 30 L_{\odot}$ and is surrounded by a large (~ 2.6 AU) optically thick dust shell. The differences between the L' and Br- γ observations can be understood with a model in which L' and Br- γ emission arises primarily from internal and external heating, respectively. We suggest that G2 is a binary star merger product and will ultimately appear similar to the B-stars that are tightly clustered around the black hole (the so-called S-star cluster).

Key words: black hole physics – Galaxy: center – Galaxy: nucleus – infrared: stars – techniques: high angular resolution – techniques: photometric

1. INTRODUCTION

Gillessen et al. (2012) reported the detection of a very red infrared object approaching the supermassive black hole (SMBH) at the Galactic Center on an orbit with a predicted closest approach of only 3000 times the radius of the event horizon. Their detection of Brackett- γ emission led them to interpret it as a dusty, 3 Earth mass (M_{\oplus}) gas cloud. If this object (G2) is indeed a gas cloud, it would be disrupted by the tidal forces of the SMBH during its closest approach and some of it would be accreted (Burkert et al. 2012; Schartmann et al. 2012; Anninos et al. 2012). Consequently, G2 has generated tremendous interest since it can be followed through the predicted accretion event and, possibly, provide new insight into accretion physics (Shcherbakov 2014; Abarca et al. 2014; Scoville & Burkert 2013; Mościbrodzka et al. 2012; Sądowski et al. 2013; Haggard et al. 2014; Chandler & Sjouwerman 2014; Gillessen et al. 2012, 2013a, 2013b).

The identification of G2 as a pure gas cloud, however, is controversial. In the pure disrupting gas cloud scenario, G2 is required to have formed relatively recently (~ 1995 –2000), close to the moment when it first became detectable with new adaptive optics technologies and at a position that is well inside the formal apoapse position (Burkert et al. 2012). Since this scenario has the unusual requirement that G2 has been fortuitously observed during the exact decade of its entire existence, many alternative models containing a central stellar source have been proposed (e.g., Murray-Clay & Loeb 2012; Miralda-Escudé 2012; Morris et al. 2012; Scoville & Burkert 2013; Ballone et al. 2013; Zajaček et al. 2014; see also Guillochon et al. 2014). The

presence of a central star would allow G2 to survive its closest approach and would not demand a recent formation event. It would also, in most scenarios, reduce the amount of gas expected to be accreted onto the central black hole following G2's closest approach (e.g., Fragile et al. 2014).

Observationally, very little is known about G2. It has been imaged as a very red point source in the near-infrared with adaptive optics systems,⁷ where it has been detected at wavelengths of 3–5 μm ($L' - M$), but is very faint at 2 μm ($m_{K'} < 20$ mag; Phifer et al. 2013; Gillessen et al. 2012; see also Eckart et al. 2013). Spectroscopically, it is a faint emission-line object, best detected in Br- γ (Phifer et al. 2013; Gillessen et al. 2013b), which shows a slightly elongated, rather compact core and some low surface brightness emission that appears to form leading and trailing tidal tails. The interpretation of exactly how much of the low surface brightness ionized gas is associated with G2 is complicated by the highly structured gas streams, also seen in emission, that abound in the projected vicinity of Sgr A* (Gillessen et al. 2012, 2013a, 2013b; Phifer et al. 2013; Meyer et al. 2014).

With a predicted closest approach having occurred in Spring 2014, the pure gas model can now be tested. In this model, the L' -band component of G2 is interpreted as dust embedded in the gas. It should follow the spatial evolution of the Br- γ component, and thus, lose its compactness at L' . In this paper, we present new L' imaging observations of G2 at the predicted moment of its closest approach and during the following few months.

⁷ Which are necessary to isolate G2 from other sources in this crowded region of the Galaxy.

Table 1
Observations of G2 and Sgr A*

Date	N_{frames}	Data Quality		Reddening Corrected Flux Density			obs. Photometry		radial profile	
		K'/L'	(FWHM) K'/L'	(Strehl) K'/L'	Sgr A* $S_{K'}$ (mJy) Measured	Sgr A* $S_{L'}$ (mJy) Inferred	G2 + Sgr A* $S_{L'}$ (mJy) Measured	G2 $m_{L'}$ Inferred	χ^2/DOF	size limit AU
Primary Data										
2014 Mar 20	22/21	67/91	0.21/0.44	1.1 ± 0.1	1.5 ± 0.2	3.7 ± 0.4	13.88 ± 0.16	0.48	<370	
2014 May 11	9/9	67/90	0.21/0.44	1.5 ± 0.2	2.2 ± 0.2	4.7 ± 0.4	13.70 ± 0.14	1.08	<350	
2014 Jul 3	64/64 ^a	70/108	0.19/0.32	1.6 ± 0.3	2.3 ± 0.4	4.9 ± 0.5	13.67 ± 0.23	1.14	<380	
2014 Aug 4	28/28	63/92	0.20/0.42	1.8 ± 0.2	2.5 ± 0.4	4.6 ± 0.6	13.92 ± 0.23	0.65	<260	
average							13.8 ± 0.2			
Auxiliary Data										
2005 Jul 30	-/56	-/81	-/0.66				13.96 ± 0.05			
2006 May 21	-/19	-/82	-/0.56				14.05 ± 0.10			
2009 Jul 22	4/-	-/86	-/0.46				14.1 ± 0.4			
2012 Jul 20–23	1314/-	-/91	-/0.48				13.9 ± 0.3			
2014 Aug 5	127/-	57/-	0.26/-	7.0 ± 0.3						

Notes. ^a(58/58) in the case of aperture photometry on the individual deconvolved frames.

2. OBSERVATIONS

New near-infrared images of G2 were obtained on 2014 March 20, May 11, July 3, August 4 and 5; 2006 May 21; and 2005 July 30 using the LGSAO system (Wizinowich et al. 2006; van Dam et al. 2006) and NIRC2 (P.I.: K. Matthews) at the W. M. Keck Observatory as part of our long-term study of the central SMBH and its environs (Ghez et al. 1998, 2008). At the time of our observations, G2 and SgrA*, the emissive source associated with the black hole, are expected to be spatially unresolved from each other in our NIR observations. The observational setup used during these measurements enables us to disentangle the emission of G2 from that of Sgr A*, which is a highly variable source at infrared wavelengths (Genzel et al. 2003; Ghez et al. 2004; Eckart et al. 2004; Hornstein et al. 2007).

Each night, images were obtained in NIRC2's narrow field mode (10 mas pixel⁻¹), interleaving observations through the K' [2.1 μm] and L' [3.8 μm] broadband filters. Individual exposure times of 28 s (10 coadds \times 2.8 s) and 30 s (60 coadds \times 0.5 s) at K' and L' , respectively, resulted in a duty cycle time of 134 s for the two-wavelength cycle. To optimize the efficiency of operations and performance, the images on G2 were obtained in a stare-mode described in Hornstein et al. (2007) and the L' sky images were taken over a range of rotator angles according to Stolte et al. (2010). Table 1 summarizes all of the new data collected and the historic data used in this analysis.

3. ANALYSIS

All the images were analyzed with standard image reduction techniques as laid out in detail for our group's earlier work (e.g., Ghez et al. 2008; Lu et al. 2009; Yelda et al. 2014). Each image was sky-subtracted, flat-fielded, and bad-pixel and image-distortion corrected. All the individual images from a given night and at the same wavelength were then averaged to obtain a deep combined map for each night and were de-convolved frame by frame with a Lucy–Richardson algorithm (Lucy 1974) to obtain time series of flux density for both bands. To determine statistical uncertainties beyond the formal fitting errors, we additionally created three submaps from three simultaneous subsets of frames.

Since G2 and Sgr A* are spatially unresolved, we disentangled the measurement of their brightness spectrally. At K' , the source at G2's predicted location is assumed to be dominated by Sgr A*.⁸ At L' , the source is expected to be the combination of G2 and Sgr A*. While Sgr A*'s brightness is highly variable, our interleaved K' measurements of Sgr A*, and the well measured and constant K' – L' color for Sgr A* (Hornstein et al. 2007; Witzel et al. 2014) allow Sgr A*'s L' flux to be estimated and removed. The details of this approach are described below.

Photometric estimates of G2 and Sgr A* are extracted using two different approaches: point-spread-function (PSF) fitting and aperture photometry on deconvolved images. In the first method, *StarFinder*, a PSF-fitting program (Diolaiti et al. 2000), is used to identify and characterize point sources in each combined map, and in the corresponding submaps. This resulted in astrometric and photometric values for Sgr A* in K' and for the unresolved G2/Sgr A* source in L' . In the second method, aperture photometry is carried out on the individual deconvolved frames. The PSF for the deconvolution process is obtained by running *StarFinder* on individual frames. The restoring beam had an FWHM that was half the resolution at each wavelength, and the aperture diameter was 60 and 120 mas for K' and L' , respectively. The aperture photometry values are obtained at each wavelength by averaging over the resulting time-series of flux densities. While this second procedure was intended to assist with any confusion with additional sources that might be near G2 or Sgr A*, the photometry in both approaches agrees well with each other. The final values are the average of the two approaches, and the differences are treated as an additional, albeit negligible, source of uncertainty.

Photometric calibration was accomplished with nonvariable calibrators (Rafelski et al. 2007; Stolte et al. 2010) in the immediate surroundings of Sgr A*. The exact set of stars used is identical to that in Ghez et al. (2008) for PSF-fitting results and to that in Witzel et al. (2012) and G. Witzel et al. (in preparation) for the results from aperture photometry on deconvolved images. The different sets were chosen to guarantee comparability between the results here and in earlier measurements, but both establish consistent zeropoints.

⁸ G2 in all earlier measurements is fainter than Sgr A* by at least a factor of 10 ($m_{K'G2} < 20$; Phifer et al. 2013; Eckart et al. 2013; Meyer et al. 2014).

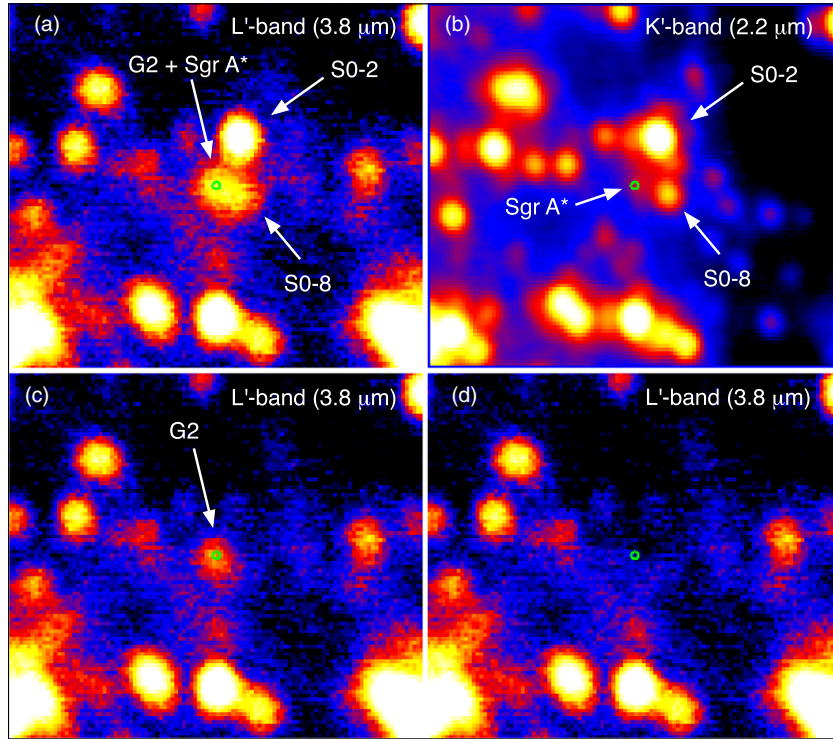


Figure 1. $1'' \times 1''$ region of L' ($3.8 \mu\text{m}$; a, c, d) and K' ($2.2 \mu\text{m}$; b) images centered on Sgr A*. These images are constructed from data obtained on 2014 March 20. The L' image (a) shows the combined flux of Sgr A* and G2, which are unresolved in these observations. The K' image (b) shows that Sgr A* is in a low emission state. Figure (c) shows the point source subtracted image (with S0-2, S0-8, and Sgr A* removed) that reveals G2. Panel (d) shows the same image after subtracting a PSF scaled to the inferred flux density of G2 at the position of G2. The clean result implies compactness of G2's spatial structure. The green circle depicts the position of Sgr A*.

To infer the flux density of Sgr A* in L' from our K' measurements, we applied the known spectral index and the extinction values in the infrared. Sgr A*'s spectral index has been shown to be constant with brightness and constant in time to within $\Delta\alpha = 0.1$ (Witzel et al. 2014), and we adopt a value of $\alpha = 0.6 \pm 0.2$, which includes the systematics of the extinction correction (Hornstein et al. 2007; Witzel et al. 2014) and which is based on the same photometric calibration applied here. For this analysis, an extinction law published by Schödel et al. (2010) and zero-points from Tokunaga (2000) were applied. We subtract the inferred L' band, reddening-corrected flux density for Sgr A* from the reddening-corrected value for the combined G2 + Sgr A* point source.

The final G2 brightness values are reported without reddening correction, for ease of comparison with earlier observed photometry. We report the photometric values for each individual observation in 2014 as well as the variance-weighted average. The uncertainties for the individual nights incorporate the statistical errors of the zero-point calibration and the submap photometry. The average value additionally includes the systematic differences between the PSF-fitting and the deconvolution method, the error of the spectral index, and the uncertainties of the extinction values.

An upper limit on G2's size for each 2014 epoch is obtained from radial profiles of G2's emission. To create the radial profile from the epoch images, the inferred L' flux of Sgr A* and the closest neighboring stars (S0-2, $m_{L'} \sim 12.7$; S0-8, $m_{L'} \sim 13.9$) are removed through PSF subtraction. The positional and brightness information for S0-2 and S0-8 are obtained from the *StarFinder* analysis discussed above. The exact location of Sgr A* is best inferred from a data set in which Sgr A* is particularly bright. Since Sgr A* was faint (< 2 mJy at K' dered-

dened) during all of the K'/L' interleaved data sets, we make use of the K' -band data obtained on 2014 August 5 when Sgr A* was bright (7 mJy, dereddened). This allows for a positional accuracy for Sgr A* of ~ 1 mas, which is dominated by the reference frame alignment uncertainty between the dates. The radial profile of G2 is compared with the PSF and the reduced χ^2 -value is calculated. Additionally, to derive an upper limit for the size of G2, the PSF is convolved with a Gaussian. The upper limit is the FWHM of the Gaussian that corresponds to a χ^2 probability of $p = 0.003$. The centroid of G2 with Sgr A* and nearby sources removed gives the position of G2, which can be determined to within 10 mas (uncertainty dominated by dust contributions; see Phifer et al. 2013).

4. RESULTS

Figure 1 shows that G2 is easily detected in our L' images. The following key results emerge from the analysis of these images.

1. *G2 has survived its closest approach to the central black hole as a compact, unresolved source at L' .* Our observations took place very near to G2's closest approach of 2014 March 16 (± 2 months) and extended well after this date (5 months). All images reveal G2 and there is no evidence for this object to be extended at L' in any 2014 epoch, as implied by the χ^2 values in Table 1. We place a 3σ upper limit on the diameter of G2's emission at L' of 32 mas, which correspond to 260 AU at a distance of ~ 8 kpc (see Table 1 and Figures 2(a) and (b)).
2. *G2's L' brightness measurements are consistent with those made over the past decade.* As the inset to Figure 2(b) displays, all our G2 L' brightness estimates from 2014 are in

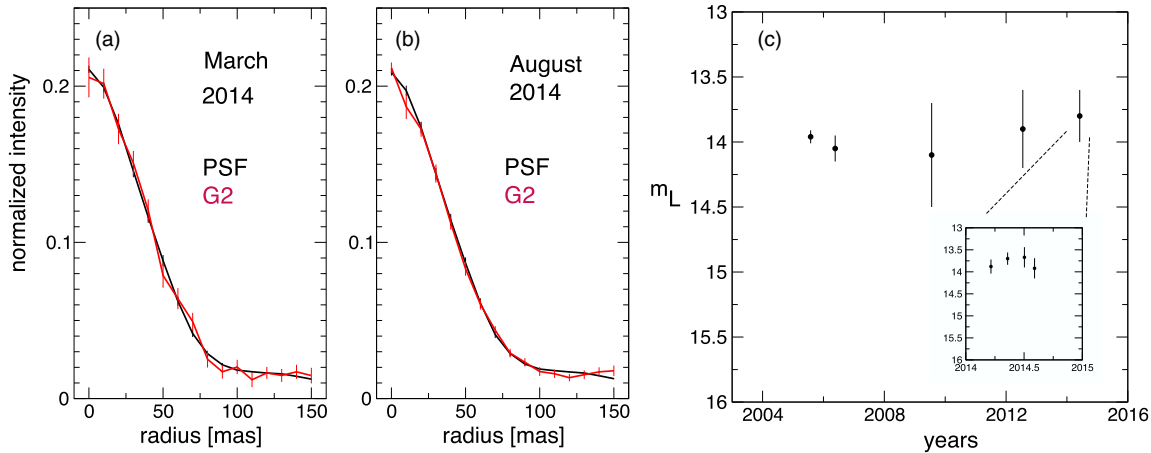


Figure 2. (a) and (b) Radial plots of G2 and the PSF, shown for the high quality 2014 March and August data sets. In the March epoch, Sgr A* was at its minimum flux of the four data sets presented in this work and, therefore, this radial plot has the smallest systematic uncertainty due to the subtraction process of Sgr A*. The error bars represent the binning statistics. G2 is fully consistent with a point source. (c) Photometry of G2 over the last 9 yr. The inset depicts the individual measurements in 2014.

agreement with one another. Furthermore, the 2014 average value of 13.8 ± 0.2 is also consistent with earlier measurements (Figure 2(b)). Table 1 summarizes the quantitative results for G2’s brightness.

3. *G2’s motion continues to be consistent with a Keplerian orbit model.* The L' position of G2 in 2014 is consistent with the predictions of our orbital model (Meyer et al. 2014). Within these predictions, the new position favors orbital solutions with shorter periods (< 500 yr) and reduces the uncertainty in the lower range of the periaapse distance a_{\min} by a factor of three ($a_{\min} = 215 \pm 30$ AU).

5. DISCUSSION AND CONCLUSIONS

Our L' measurements present a very different view of G2 from what has been seen through the Br- γ line emission measurements. The Br- γ line emission traces hot gas ($T_e \sim 10^4$ K; Gillessen et al. 2012) that appears to be *externally* heated by ionizing photons from massive stars in the vicinity of G2. In Br- γ , G2 shows clear evidence of tidal interaction with the black hole (Gillessen et al. 2012, 2013a, 2013b; Phifer et al. 2013; Pfuhl et al. 2014). This interaction and its evolution have been seen both in the increasing linewidth associated with the spatially compact “head” as well as the increasing extent of the low surface brightness tail(s) associated with G2. However, it is important to note that the Br- γ line emission measurements imply only that *some* gas associated with G2 has a size that exceeds its tidal radius.

Unlike Br- γ , the L' emission remains spatially unresolved, continues to follow a well-defined Keplerian orbit, and is constant in brightness. This allows us to rule out a pure gas cloud model as such a gas cloud model predicts that G2’s brightness and size should undergo substantial changes (e.g., Anninos et al. 2012; Pfuhl et al. 2014). Instead, G2’s L' emission appears to be coming from an optically thick dust shell surrounding an underlying star. Two lines of evidence point to this conclusion.

First, the following facts are now established.

1. The L' flux has been invariant at ~ 2.1 mJy (reddening-corrected; see Table 1) since 2005, in spite of a factor of 10 change in distance from the central black hole.
2. In 2004, G2 had an $L' - M'$ color of ~ 0.3 (reddening-corrected; Gillessen et al. 2012), which corresponds to a blackbody temperature of ~ 560 K.

3. No K' detection has been made in any epoch.
4. The L' emission is much more compact than the emission in Br- γ and thus originates in a different region.

We conclude that this is most likely explained by optically thick dust that has had constant temperature and size over the past decade and is *internally* heated. Modeling the dust emission as a blackbody results in an inferred luminosity $L_{G2} = 29 L_{\odot}$, which can easily be generated by a $2 M_{\odot}$ main-sequence star or a somewhat lower mass post- or pre-main-sequence star that is temporally very close to the main-sequence (see Figure 3(a)); higher mass stars are ruled out with the inferred luminosity.

The challenge is to explain the unusually large size inferred from the blackbody. This size is ~ 2 AU and is ~ 100 times larger than the photospheric size of $30 L_{\odot}$ stars. Several scenarios can account for such a large size, but can be excluded based on other properties. A single young star surrounded by a protoplanetary disk, as proposed by Murray-Clay & Loeb (2012), would have to be observed edge-on, which is rather unlikely. The corresponding face-on systems would not appear as such highly obscured objects. A common envelope surrounding a binary system containing a giant star would be too luminous.

As we had suggested in Phifer et al. (2013) and was followed up by Prodan et al. (2014), a binary merger is a natural model for G2. As Figure 3(b) shows, G2 has a radius in the range calculated for several putative stellar merger products (Tylanda et al. 2013; also Kamiński et al. 2010 and Zhu et al. 2013). A second line of evidence that G2 indeed is a $\sim 2 M_{\odot}$ merger product comes from tidal radius arguments. The constancy of the L' flux and the compactness of G2 lead us to conclude that the tidal radius of the source, which is proportional to its distance from the black hole, has not become smaller than the source size for most of the time. For a black hole mass of $M_{\text{BH}} = 4.3 \times 10^6 M_{\odot}$, the tidal radius is

$$r_t = 1.31 \text{ AU} \cdot \frac{R_{3\text{D}}}{215 \text{ AU}} \times \left(\frac{M_{G2}}{2 M_{\odot}} \right)^{1/3}, \quad (1)$$

where $R_{3\text{D}}$ is the 3D distance of G2 from the black hole. Figure 3(b) shows the time development of the tidal radius r_t . For a mass of $\sim 2 M_{\odot}$ the derived L' size of G2 does not show any tidal interaction with the black hole, except possibly near periaapse passage. This is consistent with G2’s L' size and

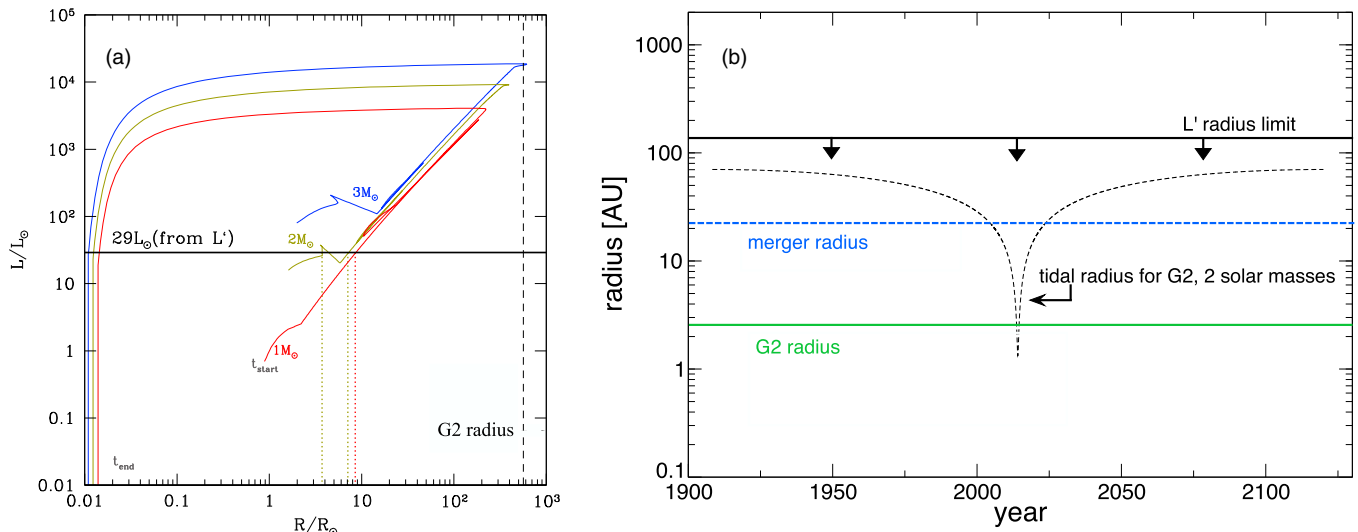


Figure 3. (a) Stellar luminosity as a function of radius. We consider three different masses, 1, 2, and $3 M_{\odot}$ and evolve them according to the stellar evolution code SSE (Hurley et al. 2000). Overplotted is the luminosity of G2 as measured from L' emission. (b) Tidal radius of G2 as a function of time assuming as mass of ~ 2 solar masses, and the radius of G2 derived for an optically thick blackbody. For comparison we show the L' radius limit derived directly from our observations and the radius of a model binary star merger.

photometry not evolving during its approach to the black hole. We note that the mass of G2, in the case in which it is indeed a merger product, can be different from the main-sequence star mass assumed here. However, the tidal radius is only weakly dependent on the mass (a factor of 10 in mass corresponds to a factor of ~ 2.2 in tidal radius).

In this picture, in which G2's L' emission originates within the tidal radius, gas, and dust beyond the tidal radius are removed by the tidal forces, creating extended optically thin tidal tails. While this extended emission is faint at L' , it is competitive with the central source at $\text{Br-}\gamma$. Thus, the $\text{Br-}\gamma$ emission is a by-product of the tidal interaction whereas the L' emission traces the properties of the merger product.

The interpretation of G2 as a merger product suggests that it will eventually look like a typical member of the S-star cluster after the extended atmosphere contracts on a Kelvin–Helmholtz time scale. The environs of the central SMBH might be particularly conducive to mergers as close binary stars in orbit around the SMBH could experience Kozai oscillations that increase the eccentricity of their orbits (Prodan et al. 2014, and references therein; Naoz et al. 2013). If this mechanism is indeed at work, it could dramatically increase the rate of stellar mergers. The plausibility of such binary interactions contributing to the central stellar population relies on the poorly known distribution and fraction of stars in binary systems at the GC (e.g., Prodan et al. 2014). Further investigations of the stellar dynamics are underway to address the question of whether it is possible that an important fraction of the S-stars has resulted from such mergers.

We wish to dedicate this paper to Gerry Neugebauer (1932–2014).

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