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Silicate Melting and Vaporization During Rocky Planet Formation

Key Points:

- We calculate specific entropy on the principal Hugoniot of forsterite (Mg_2SiO_4) by thermodynamic integration of experimental data
- We derive new criteria for shock-induced melting and vaporization
- Vaporizing collisions between small bodies are an important component of terrestrial planet formation

Supporting Information:

- Supporting Information S1

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Abstract Collisions that induce melting and vaporization can have a substantial effect on the thermal and geochemical evolution of planets. However, the thermodynamics of major minerals are not well known at the extreme conditions attained during planet formation. We obtained new data at the Sandia Z Machine and use published thermodynamic data for the major mineral forsterite (Mg_2SiO_4) to calculate the specific entropy in the liquid region of the principal Hugoniot. We use our calculated specific entropy of shocked forsterite, and revised entropies for shocked silica, to determine the critical impact velocities for melting or vaporization upon decompression from the shocked state to 1 bar and the triple points, which are near the pressures of the solar nebula. We also demonstrate the importance of the initial temperature on the criteria for vaporization. Applying these results to N -body simulations of terrestrial planet formation, we find that up to 20% to 40% of the total system mass is processed through collisions with velocities that exceed the criteria for incipient vaporization at the triple point. Vaporizing collisions between small bodies are an important component of terrestrial planet formation.

Plain Language Summary During planet formation, collisions onto planets and between planetary building blocks, such as asteroids, can be fast enough to melt or vaporize rock. Melting and vaporization changes the chemical makeup of planets. However, until recently, the extreme pressures and temperatures reached during planetary collisions could not be reproduced in laboratory experiments. We were missing key measurements on major materials that make up Earth's mantle, such as the mineral forsterite (Mg_2SiO_4). Here, we used the Z Machine, a facility at Sandia National Laboratories that can launch projectiles up to 40 km/s (almost 90,000 miles per hour), to measure the properties of forsterite at extreme conditions. Based on these measurements, we calculated that collisions faster than 8.2 km/s (about 18,000 miles per hour) can completely melt and begin to vaporize the rocky portions of planets and their building blocks. We then analyzed computer simulations of planet formation to determine how much material could have been melted or vaporized during the growth of our rocky planets. We found that 20% to 40% of all the material that makes up the inner solar system could have been involved in collisions that melted and vaporized rock.

1. Introduction

Collisions are a key aspect of planet formation (e.g., Chambers, 2010). The outcomes of collisions are diverse and complex, ranging from growth of planetesimals to the formation of synestias (Leinhardt & Stewart, 2012; Lock & Stewart, 2017). Because collisions deposit energy and redistribute material, they can significantly affect the thermal and geochemical evolution of growing planets (Asphaug, 2010; Carter et al., 2015, 2018; Stewart & Leinhardt, 2012). The increase in internal energy from a collision is generated by the shock from the initial impact and secondary shocks because of the changes in the gravitational potential well. The impact conditions required to reach the onset of melting and vaporization depend on the equations of state (EOS) and ambient conditions of the constituent materials (Kraus et al., 2012, 2015; Stewart & Ahrens, 2005).

Forsterite (Mg_2SiO_4), the magnesium end-member of the olivine system, is a major silicate phase among the first solids in the solar nebula (Lodders, 2003), and olivine ($(\text{Mg},\text{Fe})_2\text{SiO}_4$) is an abundant phase in primitive meteorites. Thus, olivine is a major phase in the mantles of differentiated planetesimals and planets. Because terrestrial olivines are rich in Mg, with compositions near $(\text{Mg}_{0.9}\text{Fe}_{0.1})_2\text{SiO}_4$, and there is an

abundance of data on forsterite, simulations of planetary collisions often use forsterite as a proxy for the bulk silicate composition of differentiated bodies. While bridgmanite, $(\text{Mg,Fe})\text{SiO}_3$, is the primary silicate phase in the lower mantle of the Earth, the upper mantle and planetary building blocks are likely to be dominated by olivines. Hence, the criteria for melting and vaporization of forsterite form important constraints for the amount of impact-induced melting and vaporization during planet formation. Although an analysis based on single component systems neglects the complexity of incongruent melting and vaporization in multicomponent systems, at this time, impact-induced phase changes in single component systems can be calculated more robustly than in multicomponent systems.

Currently, one of the most widely used EOS for forsterite among impact modelers is an extension of the ANEOS (Analytic Equations of State) model (Thompson, 1990; Thompson & Lauson, 1974) to include molecular vapor species (M-ANEOS), which was developed for silica in Melosh (2007). ANEOS is a collection of analytic expressions that describe the Helmholtz-free energy across a wide range of pressures and temperatures. ANEOS requires approximately 40 input parameters to describe a material, and there are a few different sets of input parameters for forsterite in current use for impact modeling (Canup, 2012; Canup et al., 2013; Collins & Melosh, 2014; Ćuk & Stewart, 2012; Nakajima & Stevenson, 2014). These parameter sets were developed before experimental data above 250 GPa were available. Recent measurements of the principal shock Hugoniot of forsterite (Root et al., 2018) found that the previous ANEOS models diverge from the data in the liquid region. Stewart et al. (2019) have modified ANEOS to improve the fit in the liquid region and provide a revised set of input parameters for forsterite. Although the free energy expressions in ANEOS are sometimes poor approximations of the underlying physics, it has the capability to generate thermodynamically consistent EOS models over the extremely wide range of pressures and temperatures that are vital for numerical simulations of planetary collisions.

Our goal is to describe the thermodynamics of liquid forsterite. Recently, the principal Hugoniot of forsterite has been measured at the Z machine at Sandia and the Omega laser at the U. Rochester up to 950 GPa (Root et al., 2018). These measurements provide pressure, specific volume, and temperature (P-V-T) along a single line on the EOS surface. However, to develop a wide-ranging EOS, we need other thermodynamic information. One of the most important variables to predict phase changes is specific entropy (S). Here, we present an experimentally constrained thermodynamic integration to calculate specific entropy on the principal Hugoniot. The Grüneisen parameter is needed to determine thermodynamic states off the principal Hugoniot. We present new measurements of the Grüneisen parameter obtained from shallow release experiments on the Sandia Z machine.

To apply our work to planetary collisions, we use the *entropy method* to calculate phase changes after shock compression and isentropic decompression (Ahrens & O'Keefe, 1972). Figure 1 presents a schematic of the entropy method and the thermodynamic path of a parcel of shocked material. The shock wave increases the pressure, temperature, and specific entropy to a state on the Hugoniot. The shocked material will decompress from the Hugoniot to the pressure of the surrounding medium via a rarefaction wave. This rarefaction wave propagates at the speed of sound, is faster than thermal diffusion, and does no work on the material. Consequently, decompression via a rarefaction wave is approximately isentropic, and the isentropic assumption produces an absolute lower bound to the entropy of the release state. When there is no shear strength, as in a liquid, the decompression path is reversible and isentropic. Thus, the specific entropy of the shocked state corresponds to the specific entropy of the decompressed state. Some estimates of shock-induced melting and vaporization have reported the criteria for phase changes on decompression into Earth's atmosphere at 10^5 Pa (1 bar). However, during accretion, the pressure of the protoplanetary nebula is more appropriate, for example, 1 to 10 Pa at 1 au (Wood, 2000). These pressures are close to the triple points of silicates, approximately 2 Pa for silica (Mysen & Kushiro, 1988) and 5.2 Pa for forsterite (Nagahara et al., 1994). When the final specific entropy falls in a mixed phase region, the lever rule is used to determine the mass fraction of each phase of material. The lever rule defines the mass fraction of the second phase as

$$X_2 = \frac{S_B - S_1}{S_2 - S_1}, \quad (1)$$

where S_B is the specific entropy of the bulk parcel, S_1 is the specific entropy of the first phase, and S_2 is the specific entropy of the second phase. In the case of a decompression into a liquid-vapor mixture, the first phase is liquid, the second phase is vapor, and equation (1) gives the mass fraction of the parcel that is in the vapor phase.

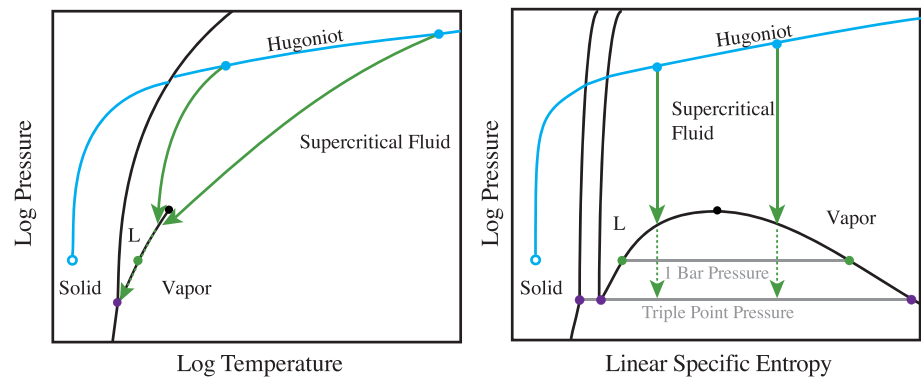


Figure 1. Schematic of a generalized single component phase diagram in T - P and S - P space. The solid, liquid (L), and vapor phase boundaries (black lines) are shown, along with the critical point (black point), triple point (purple points), and 1 bar pressure (green points). The blue curve is the Hugoniot, the locus of possible states reached by single shocks into material in a given initial state at 1 bar pressure and a particular temperature. The green lines show decompression along isentropes from specific shock states (blue points). (left) In temperature-pressure space, when the decompression isentropes intercept the liquid-vapor phase boundary, the bulk material is a mixture of liquid and vapor; however, it is difficult to determine the mass fraction of each phase. (right) In S - P space, the liquid-vapor phase boundary is represented by a dome that shows the specific entropy difference between liquid and vapor in the mixed phase region. The triple point pressure of silicates is similar to the fiducial pressure of the solar nebula at 1 au, about 10^{-4} bar.

Previous studies estimated shock-induced melting and vaporization of various silicates using the entropy method: $(\text{Mg}_{0.88}\text{Fe}_{0.12})_2\text{SiO}_4$ (Tonks & Melosh, 1993), Mg_2SiO_4 (Pierazzo et al., 1997), and SiO_2 (Kraus et al., 2012). In the case of Tonks and Melosh (1993) and Pierazzo et al. (1997), melting and vaporization were estimated using an equation of state model constructed with limited data for the vapor curve. For silica, Kraus et al. (2012) determined specific entropy on the principal Hugoniot via thermodynamic integration and constrained the liquid-vapor phase boundary from experimental data at 1 bar. Our work applies the techniques developed in Kraus et al. (2012) to derive wide-ranging EOS information.

We also investigate the process of impact-induced melting and vaporization in the context of terrestrial planet accretion. In general, numerical simulations of planet formation have focused on capturing the mechanical outcomes of collisions while neglecting the details of the thermal effects of collisions. N -body techniques are used to investigate the accretion of planets, and many studies assumed that each collision results in perfect merging (e.g., Chambers, 2001; Izidoro et al., 2016; Quintana et al., 2002; Raymond et al., 2009). However, collisional outcomes during planet accretion are diverse and include graze-and-merge, hit-and-run, erosion, and catastrophic disruption (Genda et al., 2011; Leinhardt & Stewart, 2012; Leinhardt et al., 2015). Recently, studies have modeled planet accretion in simulations that take into account imperfect merging and collisional fragmentation (Chambers, 2013; Quintana et al., 2016). To obtain more information about collisional ejecta, collisional fragmentation and re-accretion of escaped material were investigated in recent N -body simulations of planet formation where the outcomes of collisions between bodies were tracked during dynamically excited planet formation such as in the Grand Tack model (Carter et al., 2015). The Grand Tack is a model of Jupiter's evolution that was developed to explain the low mass of Mars. In this model, Jupiter migrates inward and then back outward to near its current orbit, dynamically scattering mass out of the Mars and asteroid belt region (Walsh et al., 2011). In these dynamically excited simulations, the majority of collisions during and immediately after the tack disrupt the projectile or erode the target. Even in calmer planet formation scenarios without giant planet migration, a significant fraction of collisions are erosive (Carter et al., 2015). Here, we examine the velocities of these collisions to determine how many reach the criteria for impact-induced melting and vaporization.

The total mass of material that is ejected during impacts and subsequently re-accreted can be a significant fraction of the mass of the final bodies (Bonsor et al., 2015). The effects of subsequent re-accretion of ejecta on planets have been studied: Carter et al. (2015) examined the bulk compositional effects on mantles and cores of growing planets, and Carter et al. (2018) investigated crustal erosion, but neither scrutinized the

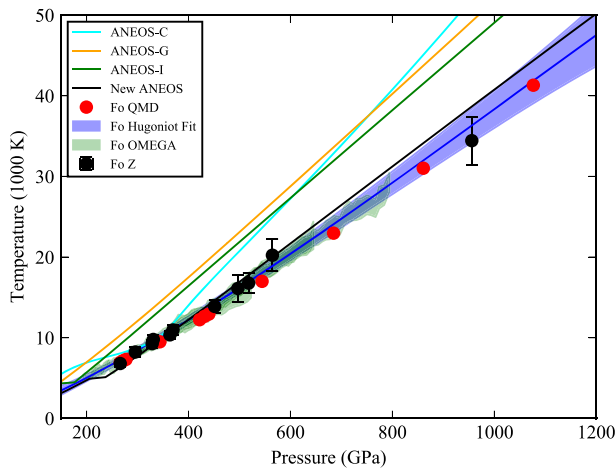


Figure 2. The principal shock Hugoniot of forsterite in pressure-temperature space. Red points are QMD calculations, Z Machine data are in black, and the green bands are from decaying shock experiments on the OMEGA laser. The blue line is a fit through the data, with the blue band representing the 1σ error envelope. QMD, Z, and Omega data are from Root et al. (2018). Z temperature measurements at about 500 and 560 GPa have been revised and are discussed in section S2. The cyan (ANEOS-C Canup et al. 2013), green (ANEOS-I Collins and Melosh ; 2014), and orange (ANEOS-G Nakajima and Stevenson ; 2014; Ćuk and Stewart ; 2012) lines are different forsterite input parameter sets for the ANEOS model (Melosh, 2007). They all overpredict the temperature over the entire measured range. The black line presents a revised ANEOS model formulation and forsterite parameter set developed using the newly available data (Stewart et al., 2019).

plate remains at a constant density, generating a planar shock wave upon impact with the target (Lemke et al., 2003; Lemke et al., 2005; Lemke et al., 2011). The OMEGA laser facility generates decaying shock waves in the mineral sample by laser ablation of a driver material. By carefully tuning the laser pulse shape (Boehly et al., 1994), thermal emission over a large, continuous range of shock pressures can be measured in a single experiment. These experiments measured shock velocity and thermal emission using a line VISAR (Celliers et al., 2004) and a streaked optical pyrometer (Miller et al., 2007). For more details of these experiments, refer to Root et al. (2018).

thermodynamic path of the ejected material. We know that impacts are capable of melting and vaporizing silicates, but we do not know the scale or conditions in which these phase changes occur during planet formation nor the potential of changing the dynamics of impact ejecta. If a significant fraction of planetary mass is processed as melted and vaporized ejecta, re-accretion of this thermally processed material may have a cumulative effect on the chemistry of the final planets.

In this work, we use the experimentally constrained forsterite principal Hugoniot from Root et al. (2018) and new measurements of the Grüneisen parameter to calculate specific entropy on the forsterite principal Hugoniot. We use specific entropy on the forsterite and silica Hugoniot to predict melt and vapor production induced by impacts. We then apply these predictions to N -body planet formation models to investigate the occurrence of impact-induced melting and vaporization during planet formation.

2. Absolute Entropy on the Forsterite Hugoniot

The principal Hugoniot for forsterite, shown in Figure 2, was measured up to 950 GPa on two platforms: Sandia National Laboratory's Z Machine and the OMEGA Laser at the University of Rochester (Root et al., 2018). The Z Machine and the OMEGA laser can reproduce most impact energies that are achieved during accretion.

The Z Machine is a pulsed power source that is capable of delivering 20 MA of current to the target assemblage over pulses of a few 100 ns (Matzen, 1997; Savage et al., 2007; Spielman et al., 1996). The current pulse is tailored to accelerate flyer plates, while the impact side of the

Table 1
Summary of Variables Used in This Work

Symbol	Description	Units
u_s	Shock velocity	km/s
u_p	Particle velocity	km/s
T	Temperature	K
P	Pressure	Pa
S	Specific entropy	J/K/kg
E	Specific internal energy	J/kg
C_x	Specific heat capacity at constant x	J/K/kg
ρ	Density	kg/m ³
V	Specific volume	m ³ /kg
f	Eulerian strain	—
γ	Grüneisen parameter	—
V_i	Impact velocity	km/s

Note. Specific heat capacity is typically measured at constant V or constant P .

Table 2
Thermodynamic Parameters for Mg_2SiO_4 Used in the Derivation of Specific Entropy on the Principal Hugoniot

Parameter	Description	Value	Units	Reference
ρ_0	Initial liquid density (1 bar and 3000 K)	2,597(\pm 11)	kg/m ³	Thomas and Asimow (2013)
γ_0	Grüneisen parameter at ρ_0	0.396	—	Thomas and Asimow (2013)
m_a	Molar mass	140.6931	mol/g	—
T_m	Melting temperature	2,174	K	Richet et al. (1993)
T_{ref}	Isentrope foot temperature	3,000	K	—
C_V	Liquid isochoric heat capacity	1,737.36	J/K/kg	Thomas and Asimow (2013)
C_P	Liquid isobaric heat capacity	1,926.18	J/K/kg	Thomas and Asimow (2013)
γ_∞	Infinite compression limit of the γ	2/3	—	—

Note. Where uncertainties are not reported in the primary reference, we describe assumed uncertainties in text.

The principal Hugoniot data and fitted Hugoniot for forsterite are shown in Figure 2. There is excellent agreement between the two experimental platforms and density functional theory based quantum molecular dynamics calculations that were also presented in Root et al. (2018). Two corrected Z temperature measurements are described in section S2. From these measurements, phase changes can be calculated upon isentropic release of shocked material to the reference pressure using the entropy method, as described in section 1. However, specific entropy cannot be measured directly and must be calculated from the thermodynamics of the system. Tables 1, 2, and 3 define the symbols, parameters, and Hugoniot equations used in this work, respectively.

To tie the known specific entropy of forsterite at standard temperature and pressure (STP) (Robie et al., 1982) to the measured principal shock Hugoniot, we calculated a thermodynamic integral using available thermodynamic data. Because state variables are independent of the thermodynamic path, the specific entropy on the principal Hugoniot can be calculated using any thermodynamic path. A schematic of the chosen thermodynamic path is shown in Figure 3. The path to the principal Hugoniot is separated into four steps,

Table 3
Equations for the Principal Hugoniots of Different Materials Used in This Work

Material	Hugoniot equations	Reference
Forsterite	$u_s(u_p) = 4.632 + 1.455u_p + (4.291E - 03)u_p^2 - (7.844E - 04)u_p^3$	Refit from Root et al. (2018)
	$T(u_s) = -183.188u_s + 15.605u_s^2 + 2.785u_s^3$	Root et al. (2018)
	$S(P) = -28691P^{-0.5} + 152.8P^{0.5} - 0.0208P^{1.5} + 3680$	This work
	$\gamma(\rho) = \gamma_\infty + (0.377 - \gamma_\infty)\left(\frac{\rho_0}{\rho}\right)^{3.705} + 0.657e^{-(\rho-4930)^2/1192^2}$	This work
α -quartz	$u_s(u_p) = 1.754 + 1.862u_p - (3.364E - 02)u_p^2 + (5.666E - 04)u_p^3$	Knudson and Desjarlais (2013)
	$S(P) = 2820(\pm 106)P^{-0.5} + 468.3(\pm 1.8)P^{0.5} - 6.68(\pm 0.37)P - 593(\pm 25.2)$	Refit from Kraus et al. (2012)
Fused Silica	$u_s(u_p) = 1.386 + 1.647u_p - (1.146E - 02)u_p^2 - (0.6952E - 04)u_p^3$	Root et al. (2019)
TPX	$u_s(u_p) = 1.795(\pm 0.018) + 1.357(\pm 0.003)u_p - 0.694(\pm 0.027)u_p e^{-0.273(\pm 0.011)u_p}$	Root et al. (2015)
All	$\frac{\rho}{\rho_0} = \frac{u_s}{u_s - u_p}$	Conservation of mass
	$P = \rho_0 u_s u_p$	Conservation of momentum
	$E - E_0 = \frac{1}{2}(P + P_0)\left(\frac{1}{\rho_0} - \frac{1}{\rho}\right)$	Conservation of energy

Note. Errors on the parameters for equations $S(P)$ for α -quartz and $u_s(u_p)$ for TPX are given. Uncertainty for $\gamma(\rho)$ of forsterite is 32%. Uncertainty for the others has covariance matrices included in the supporting information in Table S1. The forsterite principal Hugoniot equations have a range of validity from 200 to 950 GPa, and the γ function has a range from 2,597 to 6,500 kg/m³. The initial specific entropy is 669 J/K/kg, the initial temperature is 298.15 K, and the initial density (ρ_0) is 3,220 kg/m³. The quartz principal Hugoniot $u_s(u_p)$ equations have a range of validity 40 to 800 GPa, while the $S(P)$ equation has a validity between 110 and 800 GPa. The initial specific entropy is 660 J/K/kg, the initial temperature is 298.15 K, and the initial density is 2,651 kg/m³. Fused Silica principal Hugoniot equations have a range of validity 200 to 1,600 GPa, and the initial density is 2,200 kg/m³. TPX principal Hugoniot equations have a range of validity up to 985 GPa and an initial density of 833 kg/m³. The material-specific $u_s(u_p)$ equations can be transformed to pressure, specific volume, and specific internal energy via the Rankine-Hugoniot conservation equations.

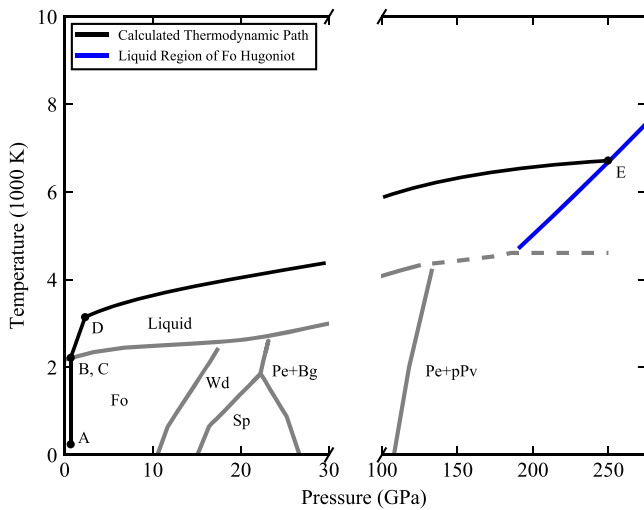


Figure 3. Simplified phase diagram of Mg_2SiO_4 and schematic thermodynamic path used to calculate specific entropy on the principal Hugoniot. Solid phases are forsterite (Fo), wadsleyite (Wd), spinel (Sp), and periclase (Pe) plus bridgmanite (Bg), from Presnall and Walter (1993). The postperovskite (pPv) plus periclase field is taken from Belonoshko et al. (2005). The melt curve (with dashed extrapolation) is an estimate from Mosenfelder et al. (2009). The black line shows the thermodynamic path chosen to integrate for specific entropy, with labels indicating points A through E. The isentrope from point D to E was chosen because it is likely to be completely in the melt region of the Mg_2SiO_4 system.

bounded by points labeled A through E. Point A is the ambient condition at STP, point B is the state after isobaric heating to the melting point of forsterite, point C accounts for the specific entropy of melting; point D is the state after isochoric heating to the foot of an isentrope, and point E is where this isentrope intersects the principal Hugoniot. Our selected path from point C to point D is in the liquid region and does not intersect the melt curve. While published thermodynamic data for forsterite exist and isentropes in the melt have been calculated (Asimow, 2018; de Koker et al., 2008; Thomas & Asimow, 2013), these isentropes are not experimentally constrained above 200 GPa. To be able to calculate isentropes at higher pressures and densities in the liquid region, we performed shallow release experiments concurrently with the shock Hugoniot measurements described above.

2.1. Shallow Release Experiments to Determine the Grüneisen parameter

The Grüneisen parameter, γ , is a thermodynamic parameter needed to calculate thermodynamic states off of the principal Hugoniot. The so-called Mie-Grüneisen approximation assumes that γ only depends on volume such that

$$\gamma = V \left(\frac{dP_{\text{th}}}{dE_{\text{th}}} \right)_V, \quad (2)$$

where the dP_{th} and dE_{th} are the thermal pressure and thermal specific internal energy. Furthermore, the thermodynamic relation,

$$\gamma = - \left(\frac{d \ln T}{d \ln V} \right)_S, \quad (3)$$

provides temperature along an isentrope. Thus, γ is the critical parameter that describes the intersection of the isentrope with the principal Hugoniot (points D to E in Figure 3).

In most solids, during compression, γ decreases with decreasing specific volume such that the following empirical relation is often used,

$$\gamma = \gamma_0 \left(\frac{V}{V_0} \right)^q, \quad (4)$$

where $q \sim 1$ for many materials (Asimow, 2018). At high compression, γ must approach a limiting value, which is often taken to be either $2/3$, for a free electron gas, or $1/2$, for the Thomas-Fermi limit (see discussion in Burakovsky & Preston, 2004). At STP, $\gamma_0 = 1.29(\pm 1)$ for solid forsterite (Gillet et al., 1991).

In contrast to solids, the γ for liquid silicates are known to increase with compression (Asimow, 2018; de Koker et al., 2008; Thomas & Asimow, 2013). The γ for initially 2273 K liquid forsterite has been inferred from shock experiments up to 114.3 GPa, with $\gamma = 0.396$ at ambient density and increasing to 1.3 at a density of 4.68 g/cm^3 (Thomas & Asimow, 2013). There is currently no data for the γ of liquid forsterite at higher compression. The value must decrease at the high-compression limit, so functional forms fit to lower pressures cannot be extrapolated. Here, we determine the Grüneisen parameter for liquid forsterite at high compression using shallow release experiments that constrain isentropic paths in $P - \rho$ space.

Shallow release experiments measure the $P - \rho$ change of release isentropes upon decompression from the shocked state. These experiments are performed by backing the forsterite sample with a transparent and lower-impedance material. The shock wave compresses the sample and transitions into the backing window that has a known Hugoniot. At the interface, the shock releases by isentropic decompression to the impedance of the window and a lower-pressure shock wave propagates forward into the window and a rarefaction wave partially decompresses the sample. At the material interface, the Rankine-Hugoniot conservation conditions are satisfied so P and u_p are identical in both materials. P and u_p are calculated from the measured u_s using known shock Hugoniots. The only measurement needed to determine the shocked state in the window, and therefore the partial release state in the forsterite sample, is u_s . For these experiments, three window materials were used, α -quartz, fused silica, and polymethylpentene, another commonly used

Table 4
Compiled Results of Experiments, Measurements, and Calculations for Partial Release of Forsterite at the Z-Machine

Shot #	Win.	V_i (km/s)	Sam. u_s (km/s)	Win. u_s (km/s)	ρ_s (kg/m ³)	ρ_r (kg/m ³)	γ
Z2792 N7	Qtz	14.82(±0.12)	14.75(±0.24)	14.15(±0.08)	6,132(±82)	5,882(±276)	1.03(±0.28)
Z2792 S7	Qtz	14.72(±0.07)	14.8(±0.19)	14.10(±0.06)	6,141(±78)	5,917(±245)	1.01(±0.28)
Z2868 N7	Qtz	15.83(±0.07)	15.46(±0.14)	14.9(±0.06)	6,264(±76)	6,015(±118)	0.95(±0.28)
Z2868 S7	Qtz	15.67(±0.06)	15.41(±0.15)	14.81(±0.05)	6,255(±76)	6,020(±128)	0.95(±0.28)
Z2879 N7	Qtz	12.95(±0.04)	13.50(±0.13)	12.83(±0.04)	5,887(±69)	5,680(±107)	1.02(±0.28)
Z2879 S7	Qtz	13.82(±0.08)	14.11(±0.16)	13.46(±0.06)	6,008(±74)	5,819(±91)	0.95(±0.27)
Z3033 N3	FS	18.89(±0.09)	17.50(±0.24)	16.94(±0.06)	6,636(±89)	6,208(±122)	0.81(±0.25)
Z3033 N5	TPX	18.84(±0.08)	17.49(±0.19)	19.17(±0.05)	6,634(±85)	5,119(±121)	—
Z3033 S3	FS	20.32(±0.06)	18.34(±0.18)	17.81(±0.04)	6,787(±88)	6,362(±105)	0.74(±0.23)
Z3033 S5	TPX	20.37(±0.11)	18.44(±0.24)	20.32(±0.04)	6,806(±94)	5,190(±133)	—
Z3044 N3	TPX	21.79(±0.07)	19.34(±0.18)	21.94(±0.05)	6,971(±94)	5,053(±163)	—
Z3044 N5	FS	21.75(±0.06)	19.34(±0.20)	19.13(±0.04)	6,971(±95)	6,432(±132)	0.88(±0.26)
Z3044 S3	TPX	23.2(±0.08)	20.11(±0.10)	23.07(±0.05)	7,115(±94)	4,105(±115)	—
Z3044 S5	FS	23.2(±0.05)	20.13(±0.10)	20.09(±0.04)	7,118(±94)	6,541(±145)	0.90(±0.26)
Z3101 N1	TPX	12.74(±0.08)	13.42(±0.11)	14.09(±0.05)	5,871(±67)	4,643(±151)	—
Z3101 S1	TPX	13.83(±0.11)	14.27(±0.12)	14.89(±0.06)	6,039(±70)	4,819(±132)	—
Z3172 N7	TPX	17.9(±0.10)	16.88(±0.15)	18.30(±0.06)	6,523(±80)	5,068(±122)	—
Z3172 S1	TPX	19.16(±0.06)	17.57(±0.16)	18.83(±0.07)	6,647(±83)	5,268(±194)	—
Z3201 S1	TPX	23.14(±0.05)	20.10(±0.19)	23.02(±0.09)	7,112(±99)	5,119(±142)	—

Note. All experiments are a forsterite sample backed by a window to tamper the release. Shock velocities in the forsterite samples are corrected from and in the same manner as Root et al. (2018), based on acceleration of the shock front at the sample-window interface. Densities at the interface before decompression are recalculated based on the corrected shock velocity. γ 's are given at the density of the release state of forsterite.

standard window referred to as TPX that have known principal shock Hugoniot (Knudson & Desjarlais, 2013; Root et al., 2015, 2019). The experimental configuration was otherwise similar to Root et al. (2018) to generate planar shock experiments at Z. The measured shock states in forsterite and the partial release states at the windows are given in Table 4.

The change in volume during isentropic decompression is given by the Riemann integral,

$$V_r = V_H + \int_{u_{pr}}^{u_{pH}} \left(\frac{du_p}{dP} \right)_S du_p, \quad (5)$$

where subscript r is the released state, subscript H is the shock Hugoniot state, and subscript S denotes constant specific entropy (Rice et al., 1958). This integral is valid only for isentropic processes. Equation (5) requires a fitted curve between the shock and release states to evaluate the integral. For this work, we found that the function $P = Au_p^{-2} + B$, where A and B are fitting parameters, could fit the release path data and be simple to evaluate in the integral above. Experiments with identical shock states are grouped and fit together to measure multiple points on the same isentrope. An example of such a fit and calculated volumes is shown in Figure 4.

To obtain γ , we calculate an isentrope that connects the shocked state to the released state by a Mie-Grüneisen release isentrope. The isentrope requires a reference curve with known P - V - E , and we reference the forsterite principal Hugoniot in a stepwise fashion as in McQueen et al. (1970). Because the density of the release states in TPX is below the range of validity of the reference curve, it is not considered to calculate γ . Pressure on the isentrope for each step is calculated via finite difference,

$$P_i = \frac{P_H - \left(E_H - E_{i-1} + P_{i-1} \frac{\Delta V}{2} \right) \left(\frac{\gamma(\rho)}{V} \right)_i}{1 + \frac{\Delta V}{2} \left(\frac{\gamma(\rho)}{V} \right)_i}, \quad (6)$$

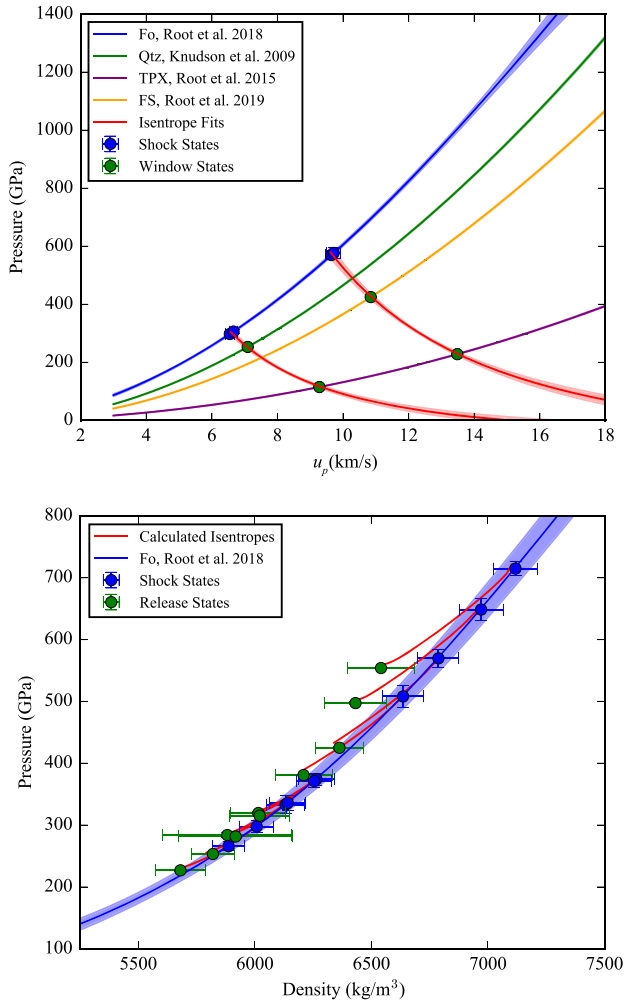


Figure 4. Results from shallow release experiments at the Z machine. (Top) Example fitted release isentropes with uncertainty envelopes (red) shown between the principal Hugoniot of forsterite (Fo), quartz (Qtz), TPX, and fused silica (FS) in pressure-particle velocity space. We derive the densities where the isentropes intersect the partial release states. Residuals of states from the P - u_p fit are incorporated into the uncertainty of the calculated density. (Bottom) All shallow release states used to calculate the Grüneisen parameter for liquid forsterite and calculated release isentropes (red).

where the specific internal energy is

$$E_i = E_{i-1} - \frac{P_i + P_{i-1}}{2} \Delta V. \quad (7)$$

The subscript i is the current element on the isentrope, indexed with density, subscript H is the reference Hugoniot state at the same density, and $\gamma(\rho)$ is some function of the Grüneisen parameter dependent only on density. We examine several different formulations for γ to solve equation (6): linear, exponential, and Gaussian, which are described in detail in section S3. Because the range of densities between the shock and release states is limited to 5,500–7,500 kg/m^3 , the formulations are intentionally simple. All formulations include a single fitting parameter, q , which is iterated until the pressure and density on the isentrope match the shock and release states.

We found that the formulation with the smallest residuals to calculate the isentropic release paths, shown in Figure 4, was the Gaussian formulation (section S3). The calculated values of γ are shown in Figure 5. The whole data set is given in Table 4.

Our γ values were combined with the lower-density data from Thomas and Asimow (2013). We find that the values for γ decrease with increasing density over our measured range. The values at the highest densities are consistent with theoretical limits (Burakovsky & Preston, 2004). None of the standard formulations for $\gamma(\rho)$ fit the entire liquid density range. To fit both sets of data, a fit of the form $\gamma(\rho) = \gamma_\infty + (A - \gamma_\infty) \left(\frac{\rho_0}{\rho}\right)^B + C e^{-(\rho-D)^2/E^2}$ is used. Our fit is only valid between the densities of 2,597–6,500 kg/m^3 where data exist. This form is a combination of previously used exponential forms (Thomas & Asimow, 2013) for low densities and the Al'tshuler form for high-density limits of γ . All parameters are fit except ρ_0 , which is the initial liquid density and the infinite compression limit, γ_∞ , which we take as 2/3. A more robust discussion on the mechanics of γ is beyond the scope of this study. The fit is given in Table 3 with 1-sigma uncertainty estimated at 32%. The data and fit are shown in Figure 5.

2.2. Entropy on the Principal Hugoniot

We calculated specific entropy change from STP to the principal Hugoniot along the path (A–E) in Figure 3. The total specific entropy is given by

$$S_{\text{Total}} = S_{\text{STP}} + \Delta S_{\text{Solid-Heating}} + \Delta S_{\text{Melting}} + \Delta S_{\text{Liquid-Heating}}. \quad (8)$$

Starting from the specific entropy at STP ($S_{\text{STP}} = 669(1) \text{ J/K/kg}$, Table 5), we calculated the increase in S due to heating forsterite isobarically to the melting point at 1 bar (points A to B). For heating at constant pressure, the change in specific entropy is given by

$$\Delta S_{\text{Solid-Heating}} = \int_{298.15}^{T_{\text{Melting}}} \frac{C_P(T)}{T} dT, \quad (9)$$

where T (K) is the temperature and $C_P(T)$ (J/mol/K) is taken from Gillet et al. (1991),

$$C_P(T) = -402.753 + 74.29 \ln(T) + \frac{87.588E3}{T} - \frac{25.913E6}{T^2} + \frac{25.374E8}{T^3}. \quad (10)$$

The total change in specific entropy from the isobaric heating to 2174 K is $\Delta S_{\text{Solid-Heating}} = 2339(\pm 195) \text{ J/K/kg}$. The specific entropy associated with melting (points B to C) is taken from Richet et al. (1993), where

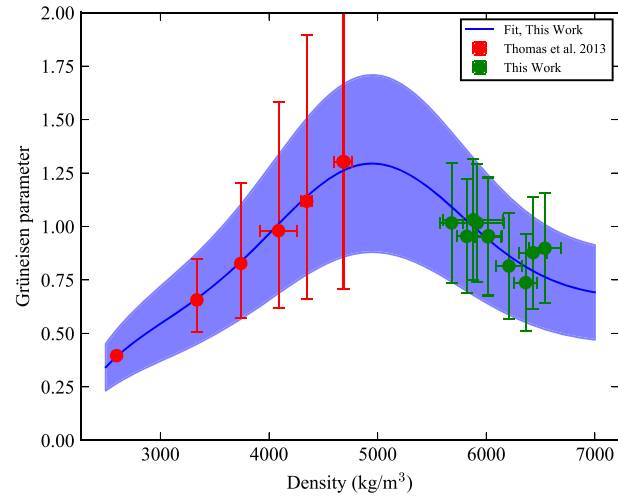


Figure 5. Experimentally derived values and our fit to the Grüneisen parameter for liquid forsterite.

$\Delta S_{\text{Melting}} = 464(\pm 4.3)$ J/K/kg. From the melting point, we calculate the specific entropy of isochorically heating liquid forsterite to a reference isentrope at 3,000 K (points C to D). This is a similar integral as before, where the only difference is that volume is held constant instead of pressure,

$$\Delta S_{\text{Liquid-Heating}} = \int_{T_{\text{Melting}}}^T \frac{C_V}{T} dT. \quad (11)$$

The heat capacity at constant volume is taken from Thomas and Asimow (2013), giving $\Delta S_{\text{Liquid-Heating}} = 559(\pm 80)$ J/K/kg.

Table 5 summarizes the specific entropy at each step and the total at the base of the isentrope. Now that we have specific entropy at the foot of the isentrope, we calculate the intersection with the principal Hugoniot (points D to E) in temperature and density.

The Grüneisen parameter along an isentrope is given in equation (3). Approximating infinitesimal steps along the isentrope and substituting specific volume with density, the temperature on the isentrope is given stepwise by

$$T_i = \exp(\gamma(\rho_i)[\ln \rho_i - \ln \rho_{i-1}] + \ln T_{i-1}), \quad (12)$$

where γ is defined in Table 3. The liquid forsterite is compressed along this isentrope until it intersects the principal shock Hugoniot, giving the specific entropy at one pressure-temperature-density point on the Hugoniot. The intersection and calculated path are shown in Figure 6. The thermodynamic values of the intersection and each step of the thermodynamic integral are summarized in Table 5.

Table 5

Pressure, Density, Temperature, and Specific Entropy Values for Each Step of the Thermodynamic Integral From STP (A), Isobaric Heating to the Melting Point (B), Melting (C), Isochoric Heating to 3000 K (D), to an Intersection Point on the Principal Hugoniot (E)

Point	Description	P (GPa)	ρ (kg/m ³)	T (K)	S (J/K/kg)	ΔS	Entropy reference
A	STP	10^{-4}	3,220(± 9.7)	298.15	669(± 1)	—	Robie et al. (1982) and Robie and Hemingway (1995)
B	Isobaric heating	10^{-4}	2,995(± 10) ^a	2,174(± 100)	3,008(± 195)	2,339(± 195)	Equation (9)
C	Melting	10^{-4}	2,597(11)	2,174(100)	3,474(± 195)	464(± 4.3)	Richet et al. (1993)
D	Isochoric heating	2.3(± 0.4)	2,597(± 11)	3,000	4,033(± 211)	559(± 80)	Equation (11)
E	Isentropic compression	231(± 54)	5,737(± 249)	6,013($\pm 1,803$)	4,033(± 211)	—	

^aDensity prior to melting is from Bouhifd et al. (1996).

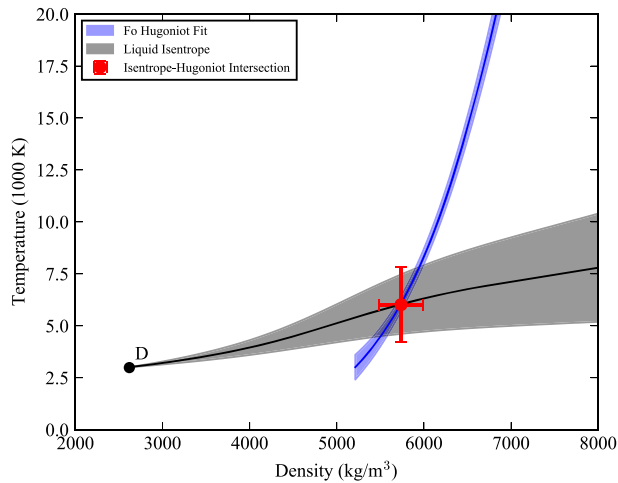


Figure 6. In density and temperature space, the principal Hugoniot of forsterite (blue line) and the calculated isentrope (black line) used for relating a known specific entropy at STP to a point along the Hugoniot. The intersection between the principal Hugoniot and isentrope is shown by the red point. Uncertainty on the intersection is dominated by uncertainty of the Grüneisen parameter. Point D is the foot of the defined isentrope as in Table 5.

Once specific entropy is known at one point on the principal Hugoniot, the specific entropy along the principal Hugoniot can be found via the first and second laws of thermodynamics,

$$dS = \frac{dE}{T} + \frac{PdV}{T} \quad (13)$$

where dS is the differential specific entropy and dV is the differential specific volume. dE , the differential specific internal energy, is known by the Rankine-Hugoniot conditions, and temperature on the principal Hugoniot is known between 200 and 950 GPa (Root et al., 2018). We fit a polynomial to our derived values for specific entropy on the principal Hugoniot, which is presented in Figure 7 and Table 3.

Uncertainty is propagated through all calculations using a *Monte Carlo* uncertainty analysis technique. All variables with a measured uncertainty are randomly perturbed according to a normal distribution about their 1-sigma uncertainty. The calculations are repeated with such random perturbations until the resulting data cloud converges to a Gaussian, typically 10,000 steps. The mean of the Gaussian is taken as the fitted value and the standard deviation to be the one sigma uncertainty. For this work, the forsterite principal Hugoniot was refit using a cubic polynomial for the purposes of facilitating *Monte Carlo* uncertainty analysis. We refit using data from Root et al. (2018) and available gas gun data above

$u_p > 4$ km/s (Jackson & Ahrens, 1979; Lyzenga & Ahrens, 1980; Mosenfelder et al., 2007; Watt & Ahrens, 1983). The largest source of uncertainty in this work is the uncertainty on the Grüneisen parameter, and future work should focus on decreasing the uncertainty between 4,000 and 5,000 kg/m³.

2.3. Comparison to Hydrocode Models for Forsterite

Our new thermodynamic data on the liquid region of the forsterite EOS provide new constraints for development of revised EOS models. Our newly calculated principal Hugoniot is compared to different model Hugoniots in Figure 7. Here, we refer to three different versions of the ANEOS model for forsterite as ANEOS-C (Canup et al., 2013), ANEOS-I (Collins & Melosh, 2014), and ANEOS-G (Čuk & Stewart, 2012; Nakajima & Stevenson, 2014). These model parameter sets were developed before the availability of the high-pressure data used in this work. All previous ANEOS models consistently overpredict temperature in shocks above 200 GPa, with greater divergence in specific entropy at higher pressures. Root et al. (2018) showed that the model principal Hugoniot from ANEOS-G falls within the experimental uncertainties in pressure-density space. Further comparisons are made between the new high-pressure data and ANEOS models in Stewart et al. (2019).

From the Rankine-Hugoniot conservation equations, the change in specific internal energy between shock states is only dependent on changes in density and pressure. Given the first law of thermodynamics, if the change in temperature is overpredicted, then the change in entropy is off as well. The new ANEOS model improves the fit in the liquid region of the Hugoniot because of the implementation of adjustable limit to the specific heat capacity. The new model parameters were focused on fitting the liquid and vapor regions of the phase diagram; see Stewart et al. (2019).

3. Shock-Induced Phase Changes

Using our new calculations of specific entropy of shocked forsterite, we revisit the question of the onset of melting and vaporization during planetary collisions. The criteria for the onset of a phase change depend on the ambient pressure and initial conditions. For applications to planetary collisions, we use two ambient reference pressures: 1 bar and the triple point. Example isentropic release paths that decompress to these pressures are shown in Figure 1.

Kraus et al. (2012) calculated the specific entropy on the silica principal Hugoniot. Here, we have extended the entropies to lower pressures and revised the fit (see section S1). We present revised criteria for shock-induced phase changes of silica, a less refractory phase, for comparison to our new forsterite results.

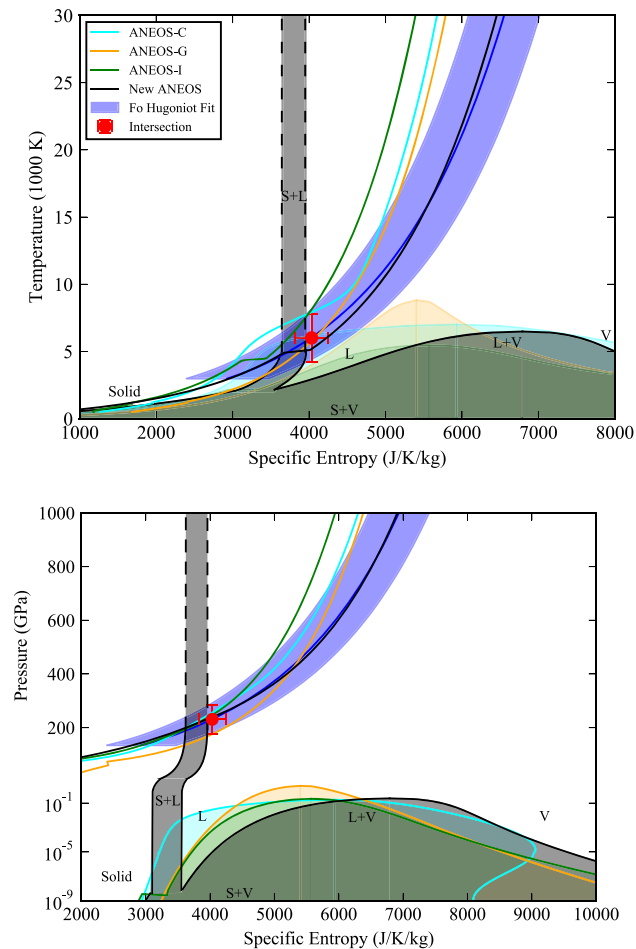


Figure 7. The principal Hugoniot is shown in blue with its associated error envelope in specific entropy space. Other solid lines are the predicted Hugoniot and liquid-vapor dome from the various ANEOS parameterizations. The red point is the calculated intersection between the isentrope and Hugoniot. The dashed black lines represent a schematic of the melt curve. (top) Temperature against specific entropy and (bottom) pressure against specific entropy of the principal Hugoniot and the ANEOS models. The divergence between the previously predicted and the measured Hugoniot should not be ignored; however, the models do lie within the extremes of the uncertainty. The new ANEOS model Hugoniot reproduces the forsterite principal Hugoniot better than the previous models.

3.1. Reference States and Initial Conditions

To use the entropy method, we require the specific entropies of phase boundaries at the ambient reference pressures: 1 bar and the triple points for forsterite (5.2 Pa; Nagahara et al., 1994) and α -quartz (2 Pa; Mysen & Kushiro, 1988).

For forsterite, the specific entropy of complete melting at 1 bar is taken from our integration, point C in Table 5, and we assume that the specific entropy of complete melting at the triple point is within uncertainty of the 1 bar value. There is no experimental data for the boiling point of forsterite at 1 bar. Here, we estimate the 1 bar boiling temperature as 3300 ± 300 K. This is based on evaporation experiments (Nagahara et al., 1994) extrapolated to 1 bar, 3265 ± 473 K, the new ANEOS boiling point, 3375 K (Stewart et al., 2019), and that forsterite's boiling point must be higher than that of silica, 3127 K (Chase et al., 1998). We calculate the specific entropy of incipient vaporization at 1 bar by thermodynamic integral similar to equation (9), using a constant isobaric heat capacity from Thomas and Asimow (2013), from the melting temperature to the boiling point. The specific entropy at incipient vaporization at the triple point is simply the specific entropy at complete melting. Specific entropies at 50% vaporization are not experimentally constrained and vary greatly among the model calculations. For this work, we use the model vapor curve from Stewart et al. (2019),

Table 6

Critical Pressures, Impact Velocities, and Specific Entropies Required for Complete Melt, Incipient Vaporization, and 50% Vapor for Two Ambient Pressures

Initial temp. (K)	Release P (Pa)	Variables	Complete melt	Incipient vap.	50% vap.
Experimental forsterite (Mg_2SiO_4), this work					
298	10^5	P (GPa)	189(± 44)	270(± 73)	856(± 189)
		V_i (km/s)	9.8(± 1.4)	12.2(± 2.0)	24.3(± 3.2)
		S (kJ/K/kg)	3.474(± 0.197)	4.270(± 0.279)	6.635
298	5.2	P	189(± 44)	189(± 44)	1,126
		V_i	9.8(± 1.4)	9.8(± 1.4)	28.6
		S	3.474(± 0.197)	3.474(± 0.197)	7.616
1200	10^5	P	142	220	791
		V_i	8.2	10.9	23.4
		S	3.474(± 0.197)	4.270(± 0.279)	6.635
1200	5.2	P	142	142	1397
		V_i	8.2	8.2	32.4
		S	3.474(± 0.197)	3.474(± 0.197)	7.616
Experimental α -quartz (SiO_2), revised from Kraus et al. (2012)					
298	10^5	P	57.6(± 2.34)	90.9(± 1.98)	260(± 5.27)
		V_i	5.89(± 0.31)	7.83(± 0.29)	14.27(± 0.39)
		S	2.950	3.560	5.394
298	2.68	P	57.6(± 2.34)	57.6(± 2.34)	371(± 8.49)
		V_i	5.89(± 0.31)	5.89(± 0.31)	17.43(± 0.45)
		S	2.950	2.950	6.100

Note. The 1 bar forsterite incipient vaporization entropy is estimated based on evaporation experiments (Nagahara et al., 1994) and the new ANEOS boiling point. The 50% vaporization entropies for forsterite are derived from the new ANEOS model vapor curve. All other values are experimentally constrained; see text for details.

labeled new ANEOS in Figure 7. While this work experimentally constrains the specific entropy of the principal forsterite Hugoniot, the pressure-volume-temperature states at higher pressures on the liquid-vapor dome still need experimental validation. For α -quartz, we use the entropies for the phase boundaries given in Kraus et al. (2012). The specific entropy of complete melting at the triple point and 1 bar is assumed to be identical. Table 6 presents all of the specific entropies of all of the phase changes considered in this work.

The interiors of differentiated bodies in the early solar system are generally warmer than room temperature. To illustrate the effects of the initial temperature on the criteria for shock-induced phase changes, we calculate two different initial temperatures for forsterite. The first initial condition in both materials is STP. For warmer initial conditions, we chose an initial temperature of 1200 K for forsterite. For this warmer initial condition, we use the new ANEOS model (Stewart et al., 2019). The 1200 K was chosen for forsterite to be near the low-pressure solidus temperature of terrestrial mantle composition. Uncertainties are only calculated for forsterite and α -quartz with initial conditions at STP, and we have not propagated uncertainties for the warm initial condition of forsterite.

3.2. Criteria for Phase Changes

Table 6 presents the critical pressures, impact velocities, and entropies to reach melting and vaporization upon release to 1 bar and the triple point reference pressures for forsterite and silica. We calculate the corresponding impact velocities by assuming that the projectile and target are both the same material and using impedance matching. The difference between forsterite, one of the most refractory silicate phases, and α -quartz, a comparatively less refractory phase, is clearly seen by the large differences in the shock pressures and impact velocities required to initiate melting and vaporization.

Because of the topology of the vapor curve in pressure-entropy space (Figure 1), the onset of vaporization occurs at a lower-specific entropy when decompressing to a lower pressure. Depending on the curvature of the liquid-vapor dome, lowering the ambient pressure can dramatically lower the specific entropy at incipient vaporization. In the cases of forsterite and silica, releasing to the triple point instead of 1 bar reduces

the incipient vaporization impact velocity by approximately 2 km/s. Overall, impact-induced phase changes require lower pressures for the onset of melting and vaporization for collisions within the pressure of the solar nebula, which is near the triple points of silicates, compared to a 1 bar reference pressure.

At larger entropies, the effect of the final pressure on the final vapor fraction can reverse because of the asymmetric topology of the vapor dome. The impact criteria for 50% vaporization increase at the triple point pressure compared to 1 bar, approximately 4 km/s for forsterite and 3 km/s for silica. We note that the shock pressures required for 50% vaporization of forsterite upon release to the triple point fall beyond the range of validity of our Hugoniot entropy calculation. The entropies are extrapolated for shock pressures beyond 1 TPa.

The initial temperature is an important factor. For the warmer initial conditions, forsterite melts and vaporizes at much lower shock pressures upon release compared to an initial temperature of 298 K. Warm forsterite (1200 K) can reduce the impact velocity required to begin vaporizing by about 1.5 km/s. It is vital not to neglect the thermal state of impacting materials especially when considering problems in the early solar system.

4. Implications for Planet Formation

Melting and vaporization can have significant effects on the thermal and chemical evolution of planetesimals and planetary bodies. However, the degree to which silicates are melted and vaporized during planet formation is unknown. Here, we assess the importance of vaporizing collisions that occur during the growth of terrestrial planets.

Considering the critical impact conditions in Table 6, cold silica and warm forsterite begin to vaporize when impact velocities exceed about 6 and 8 km/s, respectively. These velocities correspond to the escape velocities of approximately Mars mass and larger planetary embryos. Thus, essentially all collisions onto warm, differentiated planetary embryos are in or approaching the vaporization regime. However, during terrestrial planet formation, most of the mass that impacts onto the largest bodies is accreted because of their deep gravitational potential wells (Asphaug, 2010; Leinhardt & Stewart, 2012). Giant impacts involving planetary embryos are extremely high energy events that cause substantial vaporization. In these giant impacts, the vaporized mantles generate transient silicate vapor atmospheres (Carter et al., 2020; Lock & Stewart, 2017). In giant impacts, variable amounts of vaporizing material escape the growing body as ejecta. In contrast, on smaller bodies, the velocities required for vaporization fall in the erosive or catastrophic disruption regimes (Carter et al., 2019a; Leinhardt & Stewart, 2012). Next, we examine a range of collisions that occur during the growth of the rocky planets in our solar system, starting with a population of small planetesimals.

4.1. Vaporizing Collisions During Planet Formation

To investigate the importance of melting and vaporization during accretion, N -body planet formation simulations that track collisional fragmentation and re-accretion (Carter et al., 2015) were postprocessed to extract the impact conditions and outcomes.

Here we consider a high-resolution, dynamically excited N -body simulation based on Carter et al. (2015). This new simulation is a high-resolution version of their simulation 27, which includes migration of Jupiter and aerodynamic drag from the nebula. The accretion simulation began with 100,000 particles distributed between 0.5 and 3 au, with the smallest bodies having masses of $\sim 6 \times 10^{-5} M_{\oplus}$ (corresponding to radii of ~ 200 km). The mass of collision fragments smaller than this mass limit was placed into an “unresolved debris” annulus corresponding to the location of the collision. Resolved planetesimals and embryos re-accreted this debris as they passed through these annuli, thus recycling the small ejecta (Carter et al., 2015; Leinhardt & Richardson, 2005; Leinhardt et al., 2015). These simulations used particle radius inflation to reduce the computation time, so the impact velocities were first corrected for the additional acceleration associated with infall from the inflated radius to a radius calculated using an assumed density of 3 g/cm³ (the radius given above is the uninflated value). This correction is only significant for low-velocity impacts and so has a negligible effect on our results.

Using the critical impact velocities calculated above for triple point release pressures, we selected all of the collisions in the simulation that exceed the requirements for phase changes upon release to the triple point pressure. We calculated the mass of unresolved “debris” ejected during collisions that exceeded the critical impact velocities. The mass ejected in impacts that cause vaporization is shown as a cumulative sum versus

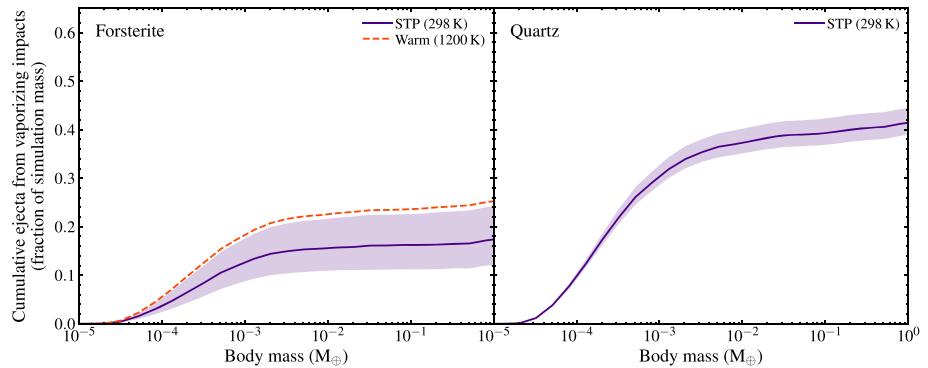


Figure 8. Cumulative fraction of total simulation mass ejected in collisions that exceed the critical impact velocities for the onset of vaporization of forsterite (left) and α -quartz (right), upon release to the triple point, against the mass of the target body. Purple curves indicate initial conditions at STP, while the dashed orange curve indicates a warm initial temperature. Warm initial conditions are used to show the dependence on temperature for the amount of melt and vapor that is induced for impacts of the same velocity. A significant fraction of mass is melted and/or vaporized, whether the material is forsterite or α -quartz. Up to 20–40% of system mass can be processed through impact-induced melting and vaporization, and the majority of such collisions occur between planetesimals. Warm initial conditions dramatically increase the amount of mass that is partially vaporized. The shaded regions indicate the range of curves calculated by adding/subtracting the corresponding uncertainties to the critical values given in Table 6. This figure uses the impact data from the new, high-resolution version of simulation 27 from Carter et al. (2015).

target mass for this example simulation of dynamically excited planet formation in Figure 8. The steepness of the curves for target bodies between 10^{-4} and $10^{-3} M_{\oplus}$ indicates that most of collisions that exceed the threshold for melting and vaporization occur between planetesimals.

Figure 8 shows that the effects of warmer initial conditions of impacts are dramatic. Collisions between warm, differentiated planetesimals (dashed orange line) substantially increase the amount of mass processed through melting and vaporization compared to cold planetesimals (purple lines). However, the parameters for cold (298 K) forsterite provide a conservative estimate for the mass of material that is processed through vaporizing collisions. Less refractory phases, such as α -quartz, melt and vaporize upon release at much lower impact velocities. Warm initial conditions only reduce the required impact velocity further.

The limitations of calculating melting and vaporization in N -body simulations need to be addressed. Even in the high-resolution N -body simulations used in this work, there is not enough information to determine the exact fraction of solid remnant in unresolved or resolved mass. Furthermore, in the calculation of ejecta, the thermal state is not taken into account nor is it evolved over time. There is also no differentiating between mass that has been processed multiple times from mass being processed for the first time. Subsequent reprocessing of the same mass inflates the total mass that is summed in Figure 8. Lastly, resolving more mass and smaller planetesimals would allow for more accurate tracking of processed material as compared to large bins of unresolved mass. Resolution is currently computationally limited.

Overall, we expect that up to 20–40% of total inner solar system mass can be involved in melting and vaporizing collisions during dynamically excited planet formation such as a Grand Tack scenario. At present, the cumulative effects of many vaporizing collisions between planetesimals are not well understood. In a Grand Tack scenario, some of the vaporizing collisions would occur within the nebular gas and some after dispersal of the gas. Recent studies found that the outcomes of vaporizing collisions with and without nebular gas will be different in terms of dispersal of material and mass of the largest remnant (Carter et al., 2019b; Davies et al., 2019; Stewart et al., 2019). In current N -body simulations the presence of nebular gas imparts a drag onto the particles, but the differences in collision outcomes are not yet taken into account. In general, collisions between small bodies that produce silicate vapor will also exceed the criteria for catastrophic disruption, where the largest fragment is less than half the total mass. As a result, the material from the colliding planetesimals is dispersed as many smaller fragments. Note that the available criteria for catastrophic disruption have been calculated for collisions in the absence of a surrounding gas (Leinhardt & Stewart, 2012).

Our results also have implications for observations of planet formation in exoplanetary systems. Several works (e.g., Meng et al., 2014; Su et al., 2019) have suggested that variable emission seen from dust in protoplanetary disks (extreme debris disks) is the result of condensed vapor produced by high-energy collisions between growing planets. The critical impact velocities for vaporizing collisions during planet growth that we have calculated here provide key information for interpreting such observations. Our results will allow constraints to be placed on the impact velocities and types of impacts occurring in time-variable disks.

While most of the material involved in vaporizing collisions will ultimately be accreted onto the final planets, partial vaporization offers the opportunity for some chemical and isotopic fractionation. Although our calculations of total mass involved in vaporizing collisions are not the same as the amount of vapor produced, our work demonstrates that major silicate phases can be vaporized. As a result, more volatile components will also be vaporized during these energetic collisions. Even in cases of small vapor fractions, the volume increase is many orders of magnitude, so the onset of vaporization is key for changing the dynamics of the ejecta. The total mass of vapor produced over time, and the composition of that vapor, requires additional information about the distribution of shock pressures within the colliding bodies and the compositions of the bodies. Whether or not a bulk chemical and isotopic fractionation would be imparted onto the final planets remains to be determined, as it depends on how efficiently the shock-processed planetesimal material is accreted. This work provides the motivation to pursue these more detailed investigations.

5. Conclusions

New experimental capabilities in shock physics provide thermodynamic data that are essential for understanding planetary materials and planet formation processes. Precise P-V-T Hugoniot measurements provide powerful constraints on the equation of state of materials, which enable a more robust understanding of material properties during planet formation. In this work, we have presented key thermodynamic data on forsterite and used our constraints on the EOS to evaluate the conditions required for shock-induced melting and vaporization during planetary collisions. Previously developed ANEOS forsterite model EOS currently in use in hydrocode simulations of collisions generally underpredicts production of entropy on the forsterite principal Hugoniot over the shock pressures encountered during planet formation. A revised ANEOS model for forsterite has been developed with our new data.

The calculations presented here offer constraints on the amount of melting and vaporization during planet formation. Cycling of mass through ejecta and re-accretion is a common process throughout planet formation, and impact velocities above the vaporization criteria are common as well. We have shown that the cycled mass can be a significant fraction of the total mass of the final terrestrial planets. Therefore, processing of planetesimals via impact-induced melting and vaporization must be considered prevalent during terrestrial planet formation. Further investigations are required to determine the full effects of vaporizing collisions.

To make more precise calculations of partially vaporized materials, we need experiments that measure the liquid-vapor boundary for forsterite. The vapor curve has been measured for silica and iron using shock-and-release techniques (Kraus et al., 2012, 2015), and similar experiments on forsterite are in progress. The addition of temperature dependence on the Grüneisen parameter model would further refine this work as well. Future coupling of equations of state data to N -body simulations would allow for the tracking of specific entropy and energy throughout planet formation.

References

- Ahrens, T. J., & O'Keefe, J. D. (1972). Shock melting and vaporization of lunar rocks and minerals. *The Moon*, 4(1-2), 214–249.
- Asimow, P. D. (2018). Melts under extreme conditions from shock experiments, *Magnas under pressure* (pp. 387–418). Amsterdam: Elsevier.
- Asphaug, E. (2010). Similar-sized collisions and the diversity of planets. *Chemie der Erde-Geochemistry*, 70(3), 199–219.
- Belonoshko, A. B., Skorodumova, N., Rosengren, A., Ahuja, R., Johansson, B., Burakovskiy, L., & Preston, D. (2005). High-pressure melting of MgSiO₃. *Physical Review Letters*, 94(19), 195701.
- Boehly, T., Craxton, R., Hinterman, T., Jaanimagi, P., Kelly, J., Kessler, T., et al. (1994). The upgrade to the omega laser system. *Fusion Technology*, 26(3P2), 722–729.
- Bonsor, A., Leinhardt, Z. M., Carter, P. J., Elliott, T., Walter, M. J., & Stewart, S. T. (2015). A collisional origin to Earth's non-chondritic composition? *Icarus*, 247, 291–300.
- Bouhifd, M. A., Andrault, D., Fiquet, G., & Richet, P. (1996). Thermal expansion of forsterite up to the melting point. *Geophysical Research Letters*, 23(10), 1143–1146.

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- Burakovsky, L., & Preston, D. L. (2004). Analytic model of the Grüneisen parameter all densities. *Journal of Physics and Chemistry of Solids*, 65(8-9), 1581–1587.
- Canup, R. M. (2012). Forming a moon with an Earth-like composition via a giant impact. *Science*, 338(6110), 1052–1055.
- Canup, R., Barr, A., & Crawford, D. (2013). Lunar-forming impacts: High-resolution SPH and AMR-CTH simulations. *Icarus*, 222(1), 200–219.
- Carter, P., Davies, E., Lock, S., & Stewart, S. (2019a). High collision velocities between planetesimals during planet growth and migration. *Lunar and planetary science conference* (Vol. 50).
- Carter, P., Davies, E., Lock, S., & Stewart, S. (2019b). Collapsing impact vapor plumes: A new planetesimal formation environment. *Lunar and planetary science conference* (Vol. 50).
- Carter, P. J., Leinhardt, Z. M., Elliott, T., Stewart, S. T., & Walter, M. J. (2018). Collisional stripping of planetary crusts. *Earth and Planetary Science Letters*, 484, 276–286.
- Carter, P. J., Leinhardt, Z. M., Elliott, T., Walter, M. J., & Stewart, S. T. (2015). Compositional evolution during rocky protoplanet accretion. *The Astrophysical Journal*, 813(1), 72.
- Carter, P. J., Lock, S. J., & Stewart, S. T. (2020). The energy budgets of giant impacts. *Journal of Geophysical Research: Planets*, 125, e2019JE006042. <https://doi.org/10.1029/2019JE006042>
- Celliers, P., Bradley, D., Collins, G., Hicks, D., Boehly, T., & Armstrong, W. (2004). Line-imaging velocimeter for shock diagnostics at the OMEGA laser facility. *Review of Scientific Instruments*, 75(11), 4916–4929.
- Chambers, J. (2001). Making more terrestrial planets. *Icarus*, 152(2), 205–224.
- Chambers, J. (2010). *Terrestrial planet formation* (pp. 297–317). Tucson, Arizona, United States: University of Arizona Press.
- Chambers, J. (2013). Late-stage planetary accretion including hit-and-run collisions and fragmentation. *Icarus*, 224(1), 43–56.
- Chase, M. W., Davies, C. A., Downey, J. R., Frurip, D. J., McDonald, R. A., & Syverud, A. N. (1998). NIST-JANAF thermodynamic tables. *Journal of Physical and Chemical Reference Data Monographs*, 9, 1717–1756.
- Collins, G. S., & Melosh, H. J. (2014). Improvements to ANEOS for multiple phase transitions. *Lunar planet. sci. conf* (Vol. 2664).
- Cuk, M., & Stewart, S. T. (2012). Making the Moon from a fast-spinning Earth: A giant impact followed by resonant despinning. *Science*, 338(6110), 1047–1052.
- Davies, E., Carter, P., Duncan, M., Root, S., Spaulding, D., Kraus, R., et al. (2019). Impact generated vapor plumes after dispersal of the solar nebula. *Lunar and planetary science conference* (Vol. 50).
- de Koker, N. P., Stixrude, L., & Karki, B. B. (2008). Thermodynamics, structure, dynamics, and freezing of Mg₂SiO₄ liquid at high pressure. *Geochimica et Cosmochimica Acta*, 72(5), 1427–1441.
- Genda, H., Kokubo, E., & Ida, S. (2011). Merging criteria for giant impacts of protoplanets. *The Astrophysical Journal*, 744(2), 137.
- Gillet, P., Richet, P., Guyot, F., & Fiquet, G. (1991). High-temperature thermodynamic properties of forsterite. *Journal of Geophysical Research*, 96(B7), 11,805–11,816.
- Izidoro, A., Raymond, S. N., Pierens, A., Morbidelli, A., Winter, O. C., & Nesvorný, D. (2016). The asteroid belt as a relic from a chaotic early solar system. *The Astrophysical Journal*, 833(1), 40.
- Jackson, I., & Ahrens, T. J. (1979). Shock wave compression of single-crystal forsterite. *Journal of Geophysical Research*, 84(B6), 3039–3048.
- Knudson, M., & Desjarlais, M. (2013). Adiabatic release measurements in α -quartz between 300 and 1200 GPa: Characterization of α -quartz as a shock standard in the multimegabar regime. *Physical Review B*, 88(18), 184107.
- Kraus, R. G., Root, S., Lemke, R. W., Stewart, S. T., Jacobsen, S. B., & Mattsson, T. R. (2015). Impact vaporization of planetesimal cores in the late stages of planet formation. *Nature Geoscience*, 8(4), 269.
- Kraus, R., Stewart, S., Swift, D., Bolme, C., Smith, R., Hamel, S., et al. (2012). Shock vaporization of silica and the thermodynamics of planetary impact events. *Journal of Geophysical Research*, 117, E09009. <https://doi.org/10.1029/2012JE004082>
- Leinhardt, Z. M., Dobinson, J., Carter, P. J., & Lines, S. (2015). Numerically predicted indirect signatures of terrestrial planet formation. *The Astrophysical Journal*, 806, 23.
- Leinhardt, Z. M., & Richardson, D. C. (2005). Planetesimals to protoplanets. I. Effect of fragmentation on terrestrial planet formation. *The Astrophysical Journal*, 625, 427–440.
- Leinhardt, Z. M., & Stewart, S. T. (2012). Collisions between gravity-dominated bodies. I. Outcome regimes and scaling laws. *The Astrophysical Journal*, 745, 79.
- Lemke, R., Knudson, M., Bliss, D., Cochran, K., Davis, J.-P., Giunta, A., et al. (2005). Magnetically accelerated, ultrahigh velocity flyer plates for shock wave experiments. *Journal of Applied Physics*, 98(7), 73530.
- Lemke, R. W., Knudson, M. D., & Davis, J.-P. (2011). Magnetically driven hyper-velocity launch capability at the Sandia Z accelerator. *International Journal of Impact Engineering*, 38(6), 480–485.
- Lemke, R., Knudson, M., Hall, C., Hail, T., Desjarlais, P., Asay, J., & Mehlhorn, T. (2003). Characterization of magnetically accelerated flyer plates. *Physics of Plasmas*, 10(4), 1092–1099.
- Lock, S. J., & Stewart, S. T. (2017). The structure of terrestrial bodies: Impact heating, corotation limits, and synestias. *Journal of Geophysical Research: Planets*, 122, 950–982. <https://doi.org/10.1002/2016JE005239>
- Lodders, K. (2003). Solar system abundances and condensation temperatures of the elements. *The Astrophysical Journal*, 591(2), 1220.
- Lyzenga, G. A., & Ahrens, T. J. (1980). Shock temperature measurements in Mg₂SiO₄ and SiO₂ at high pressures. *Geophysical Research Letters*, 7(2), 141–144.
- Matzen, M. K. (1997). Z pinches as intense x-ray sources for high-energy density physics applications. *Physics of Plasmas*, 4(5), 1519–1527.
- McQueen, R., Marsh, S., Taylor, J., Fritz, J., & Carter, W. (1970). The equation of state of solids from shock wave studies. *High Velocity Impact Phenomena*, 293, 294–417.
- Melosh, H. (2007). A hydrocode equation of state for SiO₂. *Meteoritics & Planetary Science*, 42(12), 2079–2098.
- Meng, H. Y. A., Su, K. Y. L., Rieke, G. H., Stevenson, D. J., Plavchan, P., Rujopakarn, W., et al. (2014). Large impacts around a solar-analog star in the era of terrestrial planet formation. *Science*, 345, 1032–1035. <https://doi.org/10.1126/science.1255153>
- Miller, J., Boehly, T., Melchior, A., Meyerhofer, D., Celliers, P., Eggert, J., et al. (2007). Streaked optical pyrometer system for laser-driven shock-wave experiments on OMEGA. *Review of Scientific Instruments*, 78(3), 34903.
- Mosenfelder, J. L., Asimow, P. D., & Ahrens, T. J. (2007). Thermodynamic properties of Mg₂SiO₄ liquid at ultra-high pressures from shock measurements to 200 GPa on forsterite and wadsleyite. *Journal of Geophysical Research*, 112, B06208. <https://doi.org/10.1029/2006JB004364>
- Mosenfelder, J. L., Asimow, P. D., Frost, D. J., Rubie, D. C., & Ahrens, T. J. (2009). The MgSiO₃ system at high pressure: Thermodynamic properties of perovskite, postperovskite, and melt from global inversion of shock and static compression data. *Journal of Geophysical Research*, 114, B01203. <https://doi.org/10.1029/2008JB005900>

- Mysen, B. O., & Kushiro, I. (1988). Condensation, evaporation, melting, and crystallization in the primitive solar nebula; experimental data in the system MgO-SiO₂-H₂ to 1.0 × 10⁻⁹ bar and 1870 degrees C with variable oxygen fugacity. *American Mineralogist*, 73(1-2), 1–19.
- Nagahara, H., Kushiro, I., & Mysen, B. O. (1994). Evaporation of olivine: Low pressure phase relations of the olivine system and its implication for the origin of chondritic components in the solar nebula. *Geochimica et Cosmochimica Acta*, 58(8), 1951–1963.
- Nakajima, M., & Stevenson, D. J. (2014). Investigation of the initial state of the Moon-forming disk: Bridging SPH simulations and hydrostatic models. *Icarus*, 233, 259–267.
- Pierazzo, E., Vickery, A., & Melosh, H. (1997). A reevaluation of impact melt production. *Icarus*, 127(2), 408–423.
- Presnall, D. C., & Walter, M. J. (1993). Melting of forsterite, Mg₂SiO₄, from 9.7 to 16.5 GPa. *Journal of Geophysical Research*, 98(B11), 19,777–19,783.
- Quintana, E. V., Barclay, T., Borucki, W. J., Rowe, J. F., & Chambers, J. E. (2016). The frequency of giant impacts on Earth-like worlds. *The Astrophysical Journal*, 821(2), 126.
- Quintana, E. V., Lissauer, J. J., Chambers, J. E., & Duncan, M. J. (2002). Terrestrial planet formation in the α Centauri system. *The Astrophysical Journal*, 576(2), 982.
- Raymond, S. N., O'Brien, D. P., Morbidelli, A., & Kaib, N. A. (2009). Building the terrestrial planets: Constrained accretion in the inner solar system. *Icarus*, 203(2), 644–662.
- Rice, M., McQueen, R. G., & Walsh, J. (1958). Compression of solids by strong shock waves. *Solid State Physics*, 6, 1–63.
- Richet, P., Leclerc, F., & Benoist, L. (1993). Melting of forsterite and spinel, with implications for the glass transition of Mg₂SiO₄ liquid. *Geophysical Research Letters*, 20(16), 1675–1678.
- Robie, R. A., & Hemingway, B. S. (1995). *Thermodynamic properties of minerals and related substances at 298.15 K and 1 bar (10⁵ pascals) pressure and at higher temperatures* (Vol. 2131). USA: US Government Printing Office.
- Robie, R. A., Hemingway, B. S., & Takei, H. (1982). Heat capacities and entropies of Mg₂SiO₄, Mn₂SiO₄, and Co₂SiO₄ between 5 and 380 K. *American Mineralogist*, 67(5-6), 470–482.
- Root, S., Mattsson, T. R., Cochrane, K., Lemke, R. W., & Knudson, M. D. (2015). Shock compression response of poly (4-methyl-1-pentene) plastic to 985 GPa. *Journal of Applied Physics*, 118(20), 205901.
- Root, S., Townsend, J. P., Davies, E. J., Lemke, R. W., Bliss, D. E., Fratanduono, D. E., et al. (2018). The principal Hugoniot of forsterite to 950 GPa. *Geophysical Research Letters*, 45, 3865–3872. <https://doi.org/10.1029/2017GL076931>
- Root, S., Townsend, J. P., & Knudson, M. D. (2019). Shock compression of fused silica: An impedance matching standard. *Journal of Applied Physics*, 126(16), 165901.
- Savage, M., Bennett, L., Bliss, D., Clark, W., Coats, R., Elizondo, J., et al. (2007). An overview of pulse compression and power flow in the upgraded Z pulsed power driver. *Pulsed power conference, 2007 16th IEEE international* (Vol. 2, pp. 979–984): IEEE.
- Spielman, R., Breeze, S., Deeney, C., Douglas, M., Long, F., Martin, T., et al. (1996). PBFA Z: A 20-MA z-pinch driver for plasma radiation sources, *1996 11th international conference on High-power particle beams* (Vol. 1, pp. 150–153): IET.
- Stewart, S. T., & Ahrens, T. J. (2005). Shock properties of H₂O ice. *Journal of Geophysical Research*, 110, E03005. <https://doi.org/10.1029/2004JE002305>
- Stewart, S., Carter, P., Davies, E., Lock, S., Kraus, R., Root, S., et al. (2019). Impact vapor plume expansion and hydrodynamic collapse in the solar nebula. *Lunar and planetary science conference* (Vol. 50).
- Stewart, S. T., Davies, E. J., Duncan, M. S., Lock, S. J., Root, S., Townsend, J. P., et al. (2019). The shock physics of giant impacts: Key requirements for the equations of state. arXiv preprint arXiv:1910.04687v2.
- Stewart, S. T., & Leinhardt, Z. M. (2012). Collisions between gravity-dominated bodies. II. The diversity of impact outcomes during the end stage of planet formation. *The Astrophysical Journal*, 751(1), 32.
- Su, K. Y. L., Jackson, A. P., Gáspár, A., Rieke, G. H., Dong, R., Olofsson, J., et al. (2019). Extreme debris disk variability: Exploring the diverse outcomes of large asteroid impacts during the era of terrestrial planet formation. *The Astronomical Journal*, 157(5), 202. <https://doi.org/10.3847/1538-3881/ab1260>.
- Thomas, C. W., & Asimow, P. D. (2013). Direct shock compression experiments on premolten forsterite and progress toward a consistent high-pressure equation of state for CaO-MgO-Al₂O₃-SiO₂-FeO liquids. *Journal of Geophysical Research: Solid Earth*, 118, 5738–5752. <https://doi.org/10.1002/jgrb.50374>
- Thompson, S. L. (1990). ANEOS analytic equations of state for shock physics codes input manual. Albuquerque, NM (USA): Sandia National Labs.
- Thompson, S. L., & Lauson, H. S. (1974). Improvements in the Chart D radiation-hydrodynamic CODE III: Revised analytic equations of state (SC-RR-71-0714). Albuquerque, NM: Sandia Labs.
- Tonks, W. B., & Melosh, H. J. (1993). Magma ocean formation due to giant impacts. *Journal of Geophysical Research*, 98(E3), 5319–5333.
- Walsh, K. J., Morbidelli, A., Raymond, S. N., O'Brien, D. P., & Mandell, A. M. (2011). A low mass for Mars from Jupiter's early gas-driven migration. *Nature*, 475(7355), 206.
- Watt, J. P., & Ahrens, T. J. (1983). Shock compression of single-crystal forsterite. *Journal of Geophysical Research*, 88(B11), 9500–9512.
- Wood, J. A. (2000). Pressure and temperature profiles in the solar nebula. In *From dust to terrestrial planets* (pp. 87–93). Springer.