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Characterising the target selection pipeline for the Dark Energy Spectroscopic Instrument Bright Galaxy Survey

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ABSTRACT

We present the steps taken to produce a reliable and complete input galaxy catalogue for the Dark Energy Spectroscopic Instrument (DESI) Bright Galaxy Sample (BGS) using the photometric Legacy Survey DR8 DECam. We analyze some of the main issues faced in the selection of targets for the DESI BGS, such as star-galaxy separation, contamination by fragmented stars and bright galaxies. Our pipeline utilizes a new way to select BGS galaxies using *Gaia* photometry and we implement geometrical and photometric masks that reduce the number of spurious objects. The resulting catalogue is cross-matched with the Galaxy and Mass Assembly (GAMA) survey to assess the completeness of the galaxy catalogue and the performance of the target selection. We also validate the clustering of the sources in our BGS catalogue by comparing with mock catalogues and SDSS data. Finally, the robustness of the BGS selection criteria are assessed by quantifying the dependence of the target galaxy density on imaging and other properties. The largest systematic correlation we find is a 7 per cent suppression of the target density in regions of high stellar density.

Key words: Surveys – Catalogues – large-scale structure of Universe – Galaxies

1 INTRODUCTION

The Dark Energy Spectroscopic Instrument¹ (DESI) (DESI Collaboration et al. 2016) is a multi-fibre spectrograph that will be used

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¹ <http://desi.lbl.gov/>

to carry out a number of wide-field surveys of galaxies and quasars to map the large-scale structure of the Universe. These surveys will probe the form of dark energy by allowing high precision measurements of the baryon acoustic oscillation (BAO) scale and the growth rate of structure using redshift-space distortions (RSD). The characterisation and definition of the target list for each DESI survey is a critical step for efficient survey execution and to allow reliable measurements of galaxy clustering. Here we describe this process for the DESI bright galaxy survey (hereafter BGS), a flux limited sample of around 10 million galaxies, using photometry from a new imaging survey, the Legacy Surveys² (LS).

DESI is a robotically-actuated, fibre-fed spectrograph that is capable of collecting 5 000 spectra simultaneously.

The spectra cover the wavelength range 360 to 980 nm, with a spectral resolution of $R = \lambda/\Delta\lambda$ between 2 000 and 5 500, depending on the wavelength. DESI will be used to conduct a five-year survey starting in 2020, with the aim of measuring redshifts over a solid angle of 14 000 deg². More than 30 million spectroscopic targets will be selected for four different tracer samples drawn from the imaging data. These are (i) luminous red galaxies (LRGs) in the redshift range $z = 0.3$ to $z = 1$, (ii) emission line galaxies (ELGs) to $z = 1.7$, (iii) quasars to higher redshifts ($2.1 < z < 3.5$), and (iv) a magnitude-limited BGS out to $z \approx 0.6$ with a median redshift of $z \approx 0.2$ which is the focus of this paper.

DESI observations are divided into two main programmes: the Bright Time Survey (BTS) and the Dark Time Survey (DTS). The BGS will be part of the BTS and is conducted when the Moon is above the horizon and the sky is too bright to allow efficient observation of fainter targets. The BTS excludes the few nights closest to full Moon and BGS always targets fields that are at least 40 – 50 deg away from the Moon. BGS alone will be ten times larger than the SDSS-I and SDSS-II main galaxy samples (MGS) of 1 million bright galaxies that were observed over the time period 1999 – 2008 (Abazajian et al. 2003).

The target sample for the BGS is intended to be a galaxy sample that is flux-limited in the r -band. The magnitude limit is determined by the total amount of bright observing time and the exposure times required to achieve the desired redshift efficiency. This target selection is, in essence, a deeper version of the target selection for the SDSS MGS (Strauss et al. 2002).

To make predictions for BGS target sample we make use of the mock galaxy catalogue created from the Millennium-XXL (MXXL) N -body simulation of Angulo et al. (2012) by Smith et al. (2017). This mock is tuned match the luminosity function, colour distribution, and clustering properties of the SDSS MGS at low redshift, and the evolution of these statistics to redshift $z \approx 0.5$ as measured from the GAMA survey (Driver et al. 2012; Liske et al. 2015; Baldry et al. 2017).

The DESI BGS is expected to have a target density of just over 800 galaxies per square degree in a primary sample defined by a faint r -band magnitude limit of 19.5. Then, in a lower priority sample, a secondary sample of ~ 600 galaxies deg⁻² defined by the magnitude range $19.5 < r < 20$ (DESI Collaboration et al. 2016). From hereon in we will refer to these BGS samples as BGS BRIGHT and BGS FAINT respectively. A few per cent of galaxies in the DESI BGS will be lost due to deblending errors, superposition with bright stars, and other artifacts that typically affect imaging catalogues. Our aim is to provide a reliable input galaxy catalogue for the DESI BGS and to characterize its properties, such as the

surface density of galaxies and their clustering. A complementary study by Kitaniidis et al. (2020) examined the impact of imaging systematics on the selection and clustering of targets in the LRG, ELG and QSO DESI surveys, using an earlier release of the Legacy Surveys imaging data (Dey et al. 2019).

Here, we define and characterized the BGS target selection based on the latest DECaLS release, DR8, which covers $\sim 2/3$ of the full 14 000 deg² of DESI footprint. The resulting catalogue is defined in Ruiz-Macias et al. (2020) and here we present the details of that selection and associated analysis of the catalogue. This BGS catalogue was used by DESI in the commissioning stage of the early survey validation observations. It is planned that the final BGS catalogue will be based on the next, DR9, Legacy Survey data release. This release will include better modelling of large galaxies and the light in bright star haloes. More discussion of DR9 and planned subsequent characterization of the BGS selection can be found in Section 6.

This paper is organised as follows: in Section 2 we describe the Legacy Surveys imaging data used to select our targets and the secondary datasets used to tune the selection. In Sections 3 and 4 we define the spatial and photometric cuts used to select BGS targets and to get rid of artifacts that might become problematic for DESI observations plus the removal of poor quality imaging data. In Section 4 we define our star-galaxy classification using *Gaia* DR2. In Section 5 we compare the BGS catalogue with its overlap of the GAMA DR4³ (Driver et al. 2012; Liske et al. 2015; Baldry et al. 2017) to assess the completeness and contamination of the BGS and to quantify its expected redshift distribution. In Section 5.2 we look at eight potential systematics that might be affecting our BGS target selection and try to mitigate these effects with linear weights determined using the stellar density. Section 5.3 shows the clustering of our BGS selection before and after applying the weights and we compare it with SDSS and the MXXL lightcone catalogue (Smith et al. 2017). Finally, in Section 6, we summarize our results and present our conclusions.

2 PHOTOMETRIC DATA SETS

During the BGS target selection process we make use of several catalogues. The main data set used is the Legacy Surveys DR8 (hereafter LS DR8) imaging catalogue from which we select our targets. We also make use of secondary catalogues for masking purposes, such as the Tycho-2 star catalogue (Høg et al. 2000), the *Gaia* DR2 (Gaia Collaboration et al. 2016a), the Siena Galaxy Atlas - 2020 (SGA-2020) (Moustakas in prep.) and globular clusters from the OpenNGC⁴ catalogue. We also use a combination of *Gaia* DR2 and LS photometry to perform star-galaxy separation.

³ This is an unreleased version of GAMA catalogue that the GAMA collaboration made available to us. It is essentially the same as GAMA DR3, but with more redshifts.

⁴ OpenNGC, <https://github.com/mattiaverga/OpenNGC>, is a database containing positions and main data of NGC (New General Catalogue) and IC (Index Catalogue) objects constructed by the GAVO data center team by merging data from NED, HyperLEDA, SIMBAD, and several databases available at HEASARC (<https://heasarc.gsfc.nasa.gov/>).

² <http://legacysurvey.org/>

2.1 Legacy Survey DR8 (DECam)

The Dark Energy Camera Legacy Survey (DECaLS), the Beijing-Arizona Sky Survey (BASS), and the Mayall z -band Legacy Survey (MzLS) together constitute the DESI Legacy Imaging Survey (hereafter the Legacy Survey). The imaging Legacy Survey was created with the aim of attaining photometry with the necessary target density, coverage and depth required for DESI. The SDSS MGS (Strauss et al. 2002) and Pan-STARRS1 (Chambers et al. 2016) catalogues are both too shallow to be used to reliably select the DESI survey targets. The DES survey (The Dark Energy Survey Collaboration 2005) does reach the target depth for DESI, but only covers 5000 deg², mostly in the South Galactic Cap (SGC), with only ~ 1130 deg² observable with DESI.

This work is based on the eighth release of the Legacy Survey project (LS DR8) which is the first release to integrate data from all of the individual components of the Legacy Surveys (BASS, DECaLS and MzLS). However, this paper focuses only on DECaLS data.

The DECaLS data in the LS DR8 data release comprises observations from 9th August 2014 through 7th March 2019. DECam images come from the Dark Energy Camera (DECam Flaugher et al. 2015) at the 4-m Blanco telescope at the Cerro Tololo Inter-American Observatory. DECam has 62 2048 \times 4096 pixel format 250 μ m-thick LBNL CCDs arranged in a roughly hexagonal ~ 3.2 deg² field of view. The pixel scale is 0.262 arcsec/pix and the camera has high sensitivity across a broad wavelength range of $\sim 400 - 1000$ nm. Since LS DR8 data goes beyond the intended DESI footprint⁵ of $\sim 14\,000$ deg², we are going to consider only data within the DESI footprint. This corresponds to $\sim 9\,717$ deg² of DECaLS data of which $\sim 1\,114$ deg² are covered by DECam data coming from the DES (The Dark Energy Survey Collaboration 2005). We essentially have two DECam data sets, i) DECam imaging taken for the LS programme which we refer to as DECam LS and ii) the DECam data coming from the DES programme which we refer to as DECam DES. DECam LS and DECam DES combine to form the DECaLS data set. Fig. 1 shows the sky map coverage of DECaLS imaging indicating the DECaLS imaging that lies within the DESI footprint. DECaLS is the only survey that covers the entire SGC (4394 deg²) and the NGC (5323 deg²) regions of the DESI survey at declination $\delta \leq +32.375^\circ$.

In order to fulfil the target selection required for the different DESI surveys (BGS, LRGs, ELGs and QSOs), it was concluded that a three-band g , r and z optical imaging programme, complemented by *Wide-field Infrared Survey Explorer* (WISE) W1 and W2 photometry, would be sufficient. The minimal depth⁶ required is $g = 24.0$, $r = 23.4$ and $z = 22.5$. DECam LS reaches these required depths in total exposure times of 140, 100 and 200 sec in g , r , z respectively in nominal⁷ conditions, typically in a minimum of two visits per field.

All data from the Legacy Surveys are first processed at the NSF's National Optical-Infrared Astronomy Research Laboratory

in Tucson (NSF's OIR Lab) through the NSF's OIR Lab Community Pipeline⁸ (CP). The CP takes raw data as an input and provides detrended and calibrated data products such as instrumental calibration (e.g. bias subtraction and flat fielding), astrometric calibration (e.g. mapping the distortions and providing a world coordinate system, or WCS), photometric characterization (e.g. magnitude zero point calibration) and artifact identification, masking and/or removal (e.g. removal of cross-talk and pupil ghosts, and identification and masking of cosmic rays).

The source catalogues for the Legacy Surveys are constructed using the LEGACYPIPE⁹ software, which uses the TRACTOR¹⁰ (Lang et al. 2016) code for pixel-level forward-modelling of astronomical sources. This is a statistically rigorous approach to fitting the differing point spread functions (PSF) and pixel sampling of these data, which is particularly important as the optical data have a typical PSF width of ~ 1 arcsec.

The steps in the LEGACYPIPE processing are described in Dey et al. (2019); we briefly summarize relevant parts here.

After initial source detection and defining the contiguous set of pixels associated with each detection (termed a blob), LEGACYPIPE proceeds to fit these pixels with models of the surface brightness, including a point-source and a variety of galaxy models. These fits are performed on the individual optical images (in g , r and z bands), taking into account the different PSF and sensitivity of each image, using TRACTOR.

Besides the PSF model, TRACTOR fits four other light profile models to sources: a round exponential with a variable radius (referred to as REX), an exponential profile (EXP), a de Vaucouleurs profile (DEV), and a composite of DEV and EXP profiles (COMP). The decision as to whether or not to retain an object in the catalogue and the choice of the model to best describe its light profile is treated as a penalized- χ^2 model selection problem.

This process results in object fluxes and colours that are consistently measured across the wide-area imaging surveys that form the input into the DESI target selection. In general, TRACTOR improves the target selection for all DESI surveys by allowing information from low resolution and low signal-to-noise measurements to be combined with those from high resolution and high signal-to-noise data. The TRACTOR catalogues include source positions, fluxes, shape parameters, and morphological quantities that can be used to discriminate extended sources from point-sources, together with errors on these quantities. The BGS is flux limited in the r -band. However, since TRACTOR performs simultaneous fits in g , r and z we also chose to impose quality cuts in the other bands as well as those in the r band when selecting the BGS targets.

The main TRACTOR outputs required for the BGS are the total fluxes¹¹ corresponding to the best-fitting source model (i.e., PSF, REX, EXP, DEV or COMP) in all three bands (g , r and z), the number of observations (NOBS) in the three bands, the predicted flux (in the r -band only) within the aperture of a fibre which is around 1.5 arcsec diameter (FIBERFLUX¹²) in 1 arcsec Gaussian seeing. The Galactic extinction values are derived from the SFD98 maps (Schlegel et al. 1998) and are reported in linear units of transmission (MW_TRANSMISSION) in the g , r and z bands, with a value of

⁵ Current LS DR8 imaging covers around $\sim 20\,332$ deg² of which 15 174 deg² corresponds to DECaLS.

⁶ The depths are defined as the optimal-extraction (forced-photometry) depths for a galaxy near the limiting depth of DESI, where that galaxy is defined to be an exponential profile with a half-light radius of $r_{\text{half}} = 0.45$ arcsec.

⁷ Here 'nominal' is defined as photometric and clear skies with seeing FWHM of 1.3 arcsec, airmass of 1.0, and sky brightness in g , r and z of 22.04, 20.91 and 18.46 AB mag arcsec⁻², respectively.

⁸ https://www.noao.edu/noao/staff/fvaldes/CPDocPrelim/PL201_3.html

⁹ <https://github.com/legacysurvey/legacypipe>

¹⁰ <https://github.com/dstndstn/tractor>

¹¹ The fluxes output by TRACTOR are in units called NANOMAGGIES. A flux of 1 NANOMAGGIE corresponds to an AB magnitude of 22.5.

¹² The FIBERFLUX is in units of NANOMAGGIES

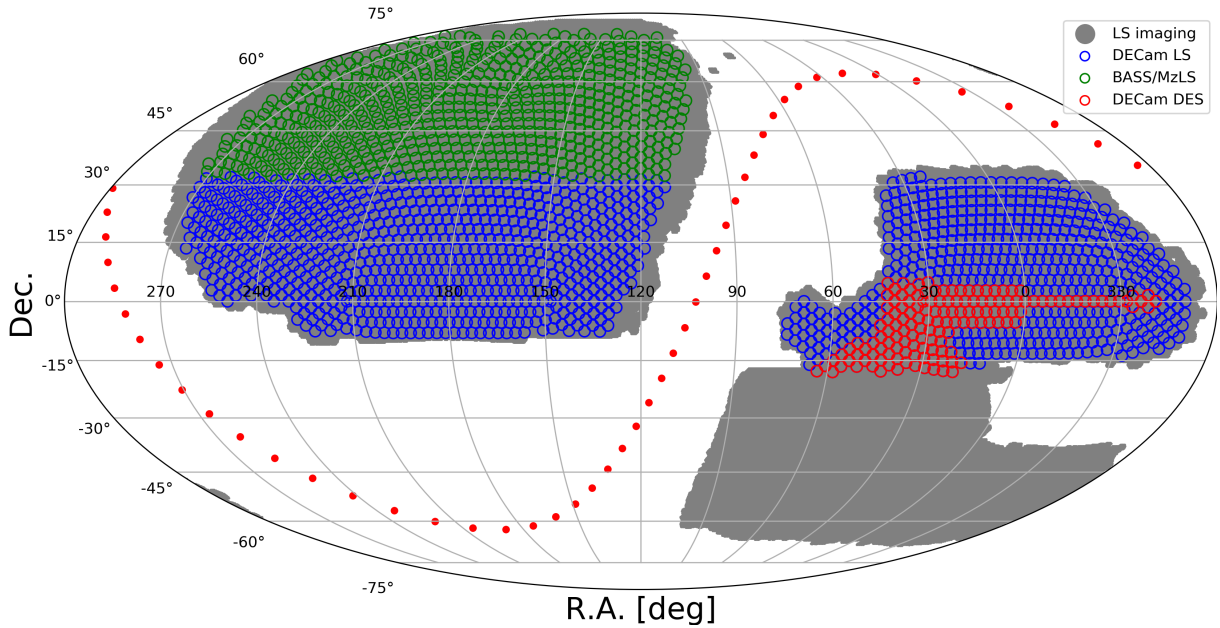


Figure 1. The sky map of the footprint of all the LS imaging used in DECALS and in BASS and MzLS is shown in gray. The red and blue circles show the DESI tiles that define the portion of DESI survey footprint that lies within DECALS. The blue tiles are those for which the data comes from the DECam LS imaging while the red tiles come from DECam DES imaging. The green tiles show the northern DESI footprint whose imaging data comes from the BASS and MzLS surveys which are not the focus of this paper. The red dots show the locus of the Galactic plane.

unity representing a fully transparent region of the Milky Way and 0 indicating a fully opaque region. The extinction coefficients for the DECam filters were computed through an airmass of 1.3, for a source with a 7000 K thermal spectrum (Schlafly & Finkbeiner 2011). The resulting coefficients are $A/E(B - V) = 3.995, 3.214, 2.165, 1.592, 1.211, 1.064$ in $ugrizY$. These are then multiplied by the SFD98 $E(B - V)$ values at the coordinates of each object to derive the g, r and z MW_TRANSMISSION values. Finally, in each band, there is a set of quality measures called FRACMASKED, FRACFLUX and FRACIN that quantify the quality of the data in each profile fit. We describe these in more detail in Section 4.4.

The fluxes returned by TRACTOR can be transformed into AB magnitudes as follows:

$$\begin{aligned} \text{mag}_r &= 22.5 - 2.5 \log_{10}(\text{FLUX}), \\ \text{mag}_z &= 22.5 - 2.5 \log_{10}(\text{FLUX}/\text{MW_TRANSMISSION}), \end{aligned} \quad (1)$$

where Eqn. (1) does not include the correction for Galactic extinction, unlike Eqn. (2). The r in Eqn. (1) stands for raw .

Table 1 shows the area covered by photometry in each of the three bands of DECALS DR8 with 1, 2 or 3 passes. These values are just for the data within the DESI footprint, as shown in Fig. 1. This DECALS footprint covers a total of 9717 deg². Expressed in percentages, 99.5 per cent of this area has at least one pass in all of the three bands grz , 95.3 per cent has at least two passes and 70.7 per cent has at least three passes in all three bands.

2.2 Secondary catalogues

Here we list other catalogues that are used either to exclude regions of the sky in which the extraction of galactic sources is compromised by the presence of other objects, or to perform star-galaxy separation.

Table 1. The area, in square degrees, of DECALS DR8 covered by at least 1, 2 or 3 passes in each of the three filters (grz) individually (first three rows), and combined (i.e. at least 1, 2 or 3 passes in each of the 3 bands; bottom row). We have restricted our results to observations within the DESI footprint as shown in Fig. 1.

Band/Number of Passes	≥ 1	≥ 2	≥ 3
g -band	9 687	9 454	7 769
r -band	9 686	9 422	7 569
z -band	9 686	9 487	8 036
combined	9 669	9 257	6 870

2.2.1 Tycho 2

Bright stars can impinge upon the estimation of the photometric properties of nearby galaxies or may even lead to the generation of spurious sources. Hence, it is prudent to simply exclude or veto regions close to known bright stars to avoid such problems. Regions near bright stars are masked out of the target catalogue using the Tycho-2 catalogue (Høg et al. 2000). The Tycho-2 catalogue contains positions, proper motions, and two-colour photometry for 2 539 913 of the brightest stars in the Milky Way.

2.2.2 Gaia DR2

Gaia (Gaia Collaboration et al. 2016a) is a European Space Agency mission that was launched in 2013 with the aim of observing ≈ 1 per cent of all the stars in the Milky Way, measuring accurate positions for them along with their proper motions, radial velocities, and optical spectrophotometry. The wavelength coverage of the astrometric instrument, defined by the white-light photometric G -band magnitude, is 330 - 1050 nm (Carrasco et al. 2016). These photometric data have a high signal-to-noise ratio and are particularly suitable for variability studies.

Since the first release of *Gaia* data (Gaia Collaboration et al.

2016b), this survey has been widely used by the DESI LS (i.e. for astrometric calibrations, proper motions, bright star masking) and is also ideal for constructing a star-galaxy separator for the BGS. There are 1.7 billion stars in the second *Gaia* data release (DR2)¹³, over the whole sky to $G = 20.7$, which is sufficiently deep to detect all stars that might contaminate the BGS FAINT sample. We describe how we use a combination of *Gaia* and LS photometry to perform star-galaxy separation in Section 4.1.

2.2.3 Globular clusters and planetary nebulae

Globular clusters and planetary nebulae are bright extended sources that can affect the identification of extragalactic sources in a similar way to bright stars. In the LS, an area of sky around such objects is excluded to minimize their impact on target selection. The OpenNGC catalogue¹⁴ is used to provide a list of such sources. The extent and impact of masking around globular clusters and planetary nebulae is discussed in Section 3.1.3.

2.2.4 The Siena Galaxy Atlas

Large galaxy images can be broken up by photometric pipelines, which, for example, could mistake H II regions inside the galaxy for individual extended sources. Also, spurious sources could be generated around the boundaries of large galaxies. The Siena Galaxy Atlas - 2020 (SGA-2020)¹⁵ is an ongoing project to select the largest galaxies in the LS using optical data from the HyperLeda catalogue¹⁶ (Makarov et al. 2014) and infrared data from the ALLWISE catalogue (Secrest et al. 2015). Currently the catalogue contains 535 292 galaxies that have an angular major axis (at the 25 mag/arcsec² isophote) larger than 20 arcsec. The use of the SGA-2020 in the spatial mask of the BGS is described in Section 3.1.2.

3 SPATIAL MASKING

Our main goal is to produce a reliable BGS input catalogue that fulfils the DESI science requirements. If the target list contains spurious objects, these will mistakenly be allocated fibres leading to a reduction in the efficiency and completeness of the redshift survey. Furthermore, spurious objects could imprint a systematic effect in the measured clustering.

A step towards minimising the number of spurious objects is to mask out regions of the sky around bright stars, since features such as extended halos, ghosts, bleed trails and diffraction spikes around the stars can compromise the measurement of the photometry of neighbouring objects. Similarly we must remove areas around very large galaxies and globular clusters and planetary nebulae; such objects can also affect the photometric measurements of their neighbours, leading to incorrect properties or spurious objects.

Within the same framework, we have to propagate instrumental effects such as saturated pixels, bad pixels, bleed trails, etc. that

the NSF's OIR Lab CP tracks and TRACTOR reports in the LS catalogue¹⁷

One way to avoid contamination of the catalogue with spurious objects is to exclude regions around bright stars and galaxies. This can be done with a simple but effective circular mask for stars and by using elliptical masks for galaxies. In Section 3.1 we set out the geometrical masking functions we have applied around bright stars, large galaxies and globular clusters to minimize the number of spurious targets in our BGS catalogue. In Section 3.2 we describe the masks applied to reduce the number of spurious targets due to imaging artifacts such as bad pixels resulting from saturation and bleed trails.

For subsequent analysis (e.g. estimating clustering statistics), it is very important to keep a record of the areas of the survey that are removed by these masks. For this purpose we have made use of the randoms catalogue developed by the DESITARGET¹⁸ team. The randoms catalogue has a total density of 50 000 objects/deg² divided into 10 subsets, each with density of 5 000 objects/deg². Each random carries with it some of the DECam imaging information computed from the image pixel (in each band and exposure) in which it is located and supplementary information such as the dust extinction extracted from HEALPix¹⁹ maps (Zonca et al. 2019). These imaging attributes include the number of observations (NOBS_G, NOBS_R, NOBS_Z), galactic extinction (EBV), the bitwise mask for optical data (MASKBITS), etc²⁰.

In Fig. 2 we show a flow chart which summarizes the spatial masking applied when constructing the BGS catalogue. The spatial masking is broken down into two classes: *geometrical masking* and *pixel masking*. The blue boxes of the flow chart report the survey area (in deg²) and mean target densities (in objects/deg²) after successively applying each mask (gray hexagonal boxes). The red boxes record the same information for the rejected area and objects. The final BGS catalogue does not depend on the order in which the masks are applied, but as some areas and targets are rejected by more than one mask the information in the red boxes depends on the ordering. For example, the area and number of objects shown as being rejected by the pixel masking excludes what would be rejected by this mask if the geometric masks had not been applied first. Overall, for the DECaLS footprint of 9 717 deg², the spatial masking removes 3.25 per cent of the area.

3.1 Geometrical masking

3.1.1 Bright star mask (BS)

The bright star (BS) mask is based on the locations of stars from *Gaia* DR2 (Gaia Collaboration et al. 2018) and the Tycho-2 (Høg et al. 2000) catalogue after correcting for epoch and proper motions. This mask consists of the union of circular exclusion regions around each star, where the radius of the exclusion region, estimated from an earlier stacking analysis, depends on the magnitude of the star in the following way:

¹⁷ In the LS DR8 catalogue information on whether or not the photometric parameters measured for an object have the possibility of being influenced by a bad pixel is flagged by the ALLMASK MASKBITS.

¹⁸ <https://github.com/desihub/desitarget>

¹⁹ <http://healpix.sourceforge.net>

²⁰ For more information on the properties of randoms see: <http://legacysurvey.org/dr8/files/#random\protect\discretionary{\char\hyphenchar\font}{-}{-}catalogs>

¹³ DR2 covers 22 months of observations and was released on 25 April 2018.

¹⁴ <https://github.com/mattiaverga/OpenNGC>

¹⁵ <https://github.com/moustakas/SGA>

¹⁶ <http://leda.univ-lyon1.fr/>

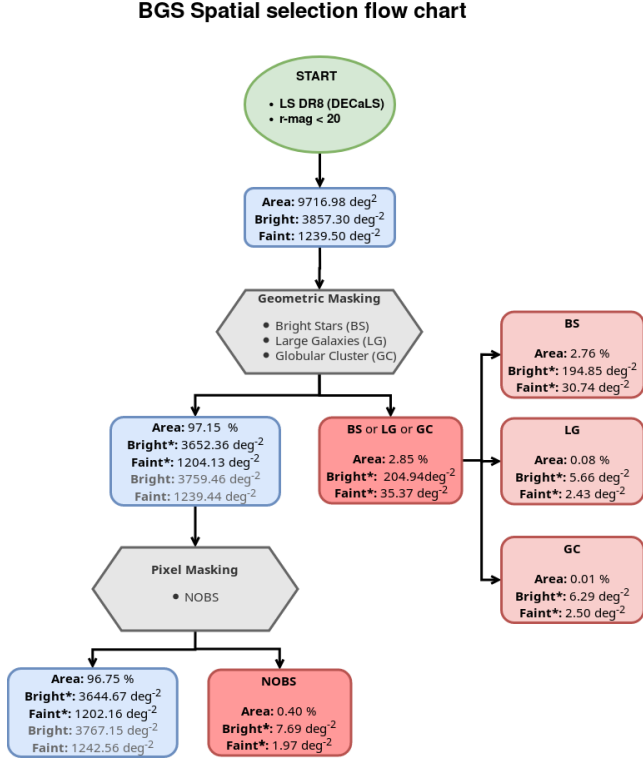


Figure 2. The flow chart shows the effects of the spatial masks that are applied as part of BGS target selection for the DECaLS DR8 data. The spatial masking is divided into two classes, one defined by the geometrical cuts which exclude regions around bright sources (bright stars, large galaxies and globular clusters), and the other by pixel-based cuts which use information such as the number of observations (NOBS). The boxes in the flow chart show the survey area (in deg^2) and the target number density (per square deg) split into BGS BRIGHT and BGS FAINT after each mask is applied. The blue boxes give this information for the portion of the survey that is retained while the red boxes give this information for the areas removed. If more than one mask is combined at a single stage (as indicated within the gray hexagonal boxes), then the dark-red boxes show the results for the combination of these masks and the light-red boxes shows the results for each individual mask. As some of the masks can overlap the numbers in the light-red boxes do not necessarily add up to those in the dark-red boxes. The target densities with the (*) superscript are computed without correcting for the area removed by the masking while those without the (*) superscript are corrected for the masked area. The gray hexagonal boxes describe the different masks. Note that star-galaxy separation is not yet applied here and this is why we have a high target density in the blue boxes.

$$\begin{aligned}
 R_{\text{BS}}(m) &= 39.3 \times 2.5^{(11-m)/3} \text{ arcsec}, & m > 2.9 & \quad (3) \\
 &= 471.6 \text{ arcsec}, & m < 2.9. &
 \end{aligned}$$

Here m is either *Gaia* G -mag or Tycho-2 MAG_{VT} with *Gaia* G -mag being used when both are available. Stars fainter than $m = 13$ have no exclusion zone around them.

The BS masking uses a total of 773 673 *Gaia* DR2 objects (82 objects/deg^2) with *Gaia* G -mag brighter than 13, while from Tycho-2, we have a total of 3 349 objects ($\sim 0.36 \text{ objects/deg}^2$) to a Tycho-2 visual magnitude brighter than $\text{MAG}_{\text{VT}} = 13$. In order to avoid overlaps both catalogues have been matched after applying proper motions to bring *Gaia* objects to the same epoch as Tycho-2 and keeping only the Tycho-2 objects that are not found in *Gaia*.

These Tycho-2 stars represents only a 0.4% of total stars used for the BS masking. Then the magnitude, m , used to compute the mask radius in equation (3) is the *Gaia* G -band magnitude for the *Gaia* stars and the Tycho-2 visual magnitude, MAG_{VT} , for the retained Tycho-2 stars. The overall median difference between the Tycho-2 and *Gaia* magnitude is 0.4 with Tycho-2 being fainter. This 0.4 magnitude difference translates into a median decrease in masking radius of 50 arcsecs for *Gaia* stars with magnitude of 3 and a decrease of 2 arcsecs for *Gaia* stars with magnitude of 13 from equation 3. Within $R_{\text{BS}}(m)$ TRACTOR forces all the sources it detects to be fit with the PSF profile to avoid artificially fitting diffraction spikes and stellar haloes as large extended sources. Thus any galaxies detected within R_{BS} will have their fluxes underestimated. Consequently to define a reliable galaxy catalogue we must veto all sources within R_{BS} of a bright star. In Fig. 2 we show that this Bright star mask covers 2.76 per cent of the initial footprint and rejects ~ 195 potential BGS BRIGHT objects/ deg^2 and ~ 31 potential BGS FAINT objects/ deg^2 when averaged over the full initial footprint. It should be noted that most of these objects are stars as star-galaxy separation has not been applied at this stage in the flow chart shown in Fig. 2. An alternative ordering of the flow chart with star-galaxy separation applied first is shown in Fig. A1. There we see that for galaxies the corresponding numbers are 13.7 galaxies/ deg^2 for BGS BRIGHT and 8.5 galaxies/ deg^2 for BGS FAINT.

To determine if the bright star mask is adequate or whether the effects of stellar haloes causes a systematic error in the photometry of neighbouring galaxies that extends to larger radii, we plot in Fig. 3 the average density of BGS galaxies in the vicinity of bright stars prior to applying the bright star mask. If the photometry of galaxies has been compromised in any means, this can be seen in the galaxy number density to a fixed magnitude due to the strong dependence of galaxy number density on apparent magnitude. The term BGS galaxy refers to the BGS sample after applying the star-galaxy separation and the spatial and photometric cuts down to the r -band magnitude of 20, which will be covered in the subsequent subsections of Section 3 and in Section 4. The stacks are made by expressing the angular separation, r , of the BGS galaxies prior to apply the bright star mask from their nearest bright star in units of the bright star masking radius R_{BS} , as given by Eqn. 3. In these rescaled coordinates, $R = r/R_{\text{BS}}$, galaxies within a radius of unity, shown by the black circle, are within the BS masking zone. We show stacks for two magnitude bins defined by the G -mag and visual magnitude MAG_{VT} for *Gaia* DR2 and Tycho-2 stars respectively, one with bright stars of magnitude between 8 to 12 and one fainter with magnitude between 12 to 13. The radial profile (red solid line) shows the variation in the target density, defined as $\Delta\rho(R) \equiv \log_2(\eta(R)/\bar{\eta})$ where $\eta(R)$ is the target density in an annulus at radius R of width $\Delta R \sim 0.06$, and $\bar{\eta}$ is the mean target density evaluated over the region $1.1 < R < 3$. This means that $\Delta\rho(R) = 0$ corresponds to the mean density, $\Delta\rho(R) \geq 1$ to an overdensity at least twice the mean density, and $\Delta\rho(R) < 0$ to an underdensity. The large underdensity at radius $R \leq 1$ is due to TRACTOR forcing all objects within this region to be fit by the PSF model. In Section 4.1 we will see how stars and galaxies are defined for BGS target selection, which does not depend on TRACTOR PSF designation, therefore, galaxies in the region $R < 1$ are allowed. In the left panel of Fig. 3, we see a spike of spurious galaxies for $R < 0.2$. In contrast the right panel shows a strong deficit of galaxies at $R < 0.2$. For $R > 1$, the stacks show uniform density close to mean, suggesting the star mask is working. There is a small bump just outside the masking radius where a ~ 6 per cent excess is seen in both panels. This may need to

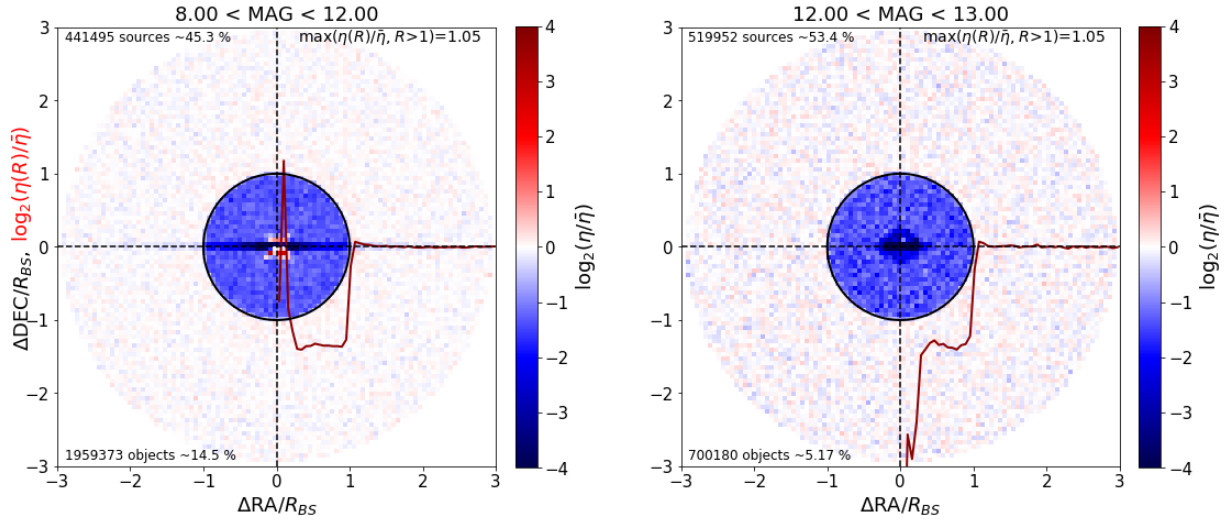


Figure 3. 2D histograms of the positions of BGS objects relative to their nearest Bright Star (BS) taken from the *Gaia* and Tycho-2 sources down to G -mag and visual magnitude MAG_{VT} of 13 respectively. These stacks are performed in magnitude bins in the BS catalogue from magnitude 8 to 12 (left) and 12 to 13 (right). The stacks are made using angular separations rescaled to the masking radius function given in Eqn 3, which means that objects within a scaled radius of 0 to 1 will be masked out by the BS veto while objects with $R = r/R_{BS} > 1$ will not (here $r^2 = (\Delta RA)^2 \cos(DEC)^2 + \Delta DEC^2$). The colour scale shows the ratio of the density per pixel (η) to the mean density ($\bar{\eta}$) within the shell $1.1 < r/R_{BS} < 3$. The density ratio is shown on a \log_2 scale where red shows overdensities, blue corresponds to underdensities and white shows the mean density. The black solid circle shows extent of the BS exclusion zone. The red solid line shows the radial density profile on the same scale as the colour distribution $\log_2(\eta/\bar{\eta})$ where $\eta(R)$ is the target density within the annulus at radius R of width $\Delta R \sim 0.06$.

be revisited for accurate clustering studies, but is not large enough to be a concern for the efficiency of target selection.

3.1.2 Large galaxies mask (LG)

Without special treatment, large galaxies in which spiral arms and other structures such as H II regions are resolved would be artificially fragmented by TRACTOR into multiple sources. To avoid this and to achieve more accurate photometry for large galaxies in the SGA-2020 catalogue (see §2.2.4), TRACTOR is seeded with different priors, and within an elliptical mask centred on the large galaxy TRACTOR fits secondary detections using only the PSF model. This reduces the spurious fragmentation of large galaxy images, but also means that genuine neighbouring galaxies within the masked area have compromised photometry. The elliptical mask that is used has the same position, 25 mag/arcsec² isophotal major axis angular diameter, D_{25} , semi-minor to semi-major ratio, B/A and position angle, PA as the ones used to define the large galaxies in the SGA-2020 catalogue. Defining an effective masking radius of $r = \sqrt{ab}$, where a and b are the semi-major and semi-minor axes of the elliptical mask, the median masking radius for the LG galaxies is 10.8 arcsecs.

We apply these same masks to reject objects from the BGS catalogue but then we reinstate the large galaxies provided they are not also masked by the bright star or globular cluster mask. The area covered by the combined LG mask amounts to only 0.08 per cent of the initial area and the number of objects removed amounts to 5.7 objects/deg² BGS BRIGHT and 2.4 objects/deg² BGS FAINT objects over the full initial area.

3.1.3 Globular cluster mask (GC)

The globular cluster (GC) mask works in a similar way to the BS mask, by applying a circular exclusion zone around the GC. The masking radius is defined by the major axis attribute for the object in the OpenNGC catalogue.

The GC mask has the smallest impact of the geometric masks, rejecting only 0.01 per cent of the initial area, accounting for densities of 6.3 objects/deg² in BGS BRIGHT and 2.5 objects/deg² in BGS FAINT. TRACTOR also force fits as PSFs everything within this mask.

3.2 Pixel masking

Some of the effects that compromise the photometry on a pixel basis and the model fitting include bad pixels, saturation, cosmic rays, bleed trails, transients. The NSF's OIR Lab DECam CP identifies these instrumental effects during its various calibrations²¹ (see Table 5 in Dey et al. (2019) for a list of the calibrations) and these are passed through TRACTOR and compiled in the ALLMASK BITMASK²². ALLMASK denotes a source blob that overlaps with any of the mentioned bad pixels in all of the overlapping images.

Besides the bad pixels which arise due to instrumental defects, the BGS requires a complete sample in the three bands (g , r and z). We therefore impose a requirement that there is at least one observation in each of the bands through the NOBS parameter. NOBS stands for Number of Observations, and is defined as the number of images that contributes to the source detected central pixel in each of

²¹ The document that lists all the calibrations and which includes details about the various maskings can be found at: https://www.noao.edu/noao/staff/fvaldes/CPDocPrelim/PL201_3.html

²² Details of this BITMASK can be found here: <http://www.legacysurvey.org/dr8/bitmasks/#allmask-x-anymask-x>

the bands. Both ALLMASK and NOBS are pixel-based and hence this information is also available in the random catalogue. However, we find that virtually all of the area (97 per cent) (and hence virtually all of the randoms) rejected by ALLMASK is also rejected by using NOBS = 0 (in any band). In addition, ALLMASK rejects a significant number of objects (196 objects/deg²) but with a small associated area (0.01 per cent of the full area). Virtually all the objects rejected by ALLMASK and many others are already rejected by the quality cuts in FRACMASKED, FRACIN and FRACFLUX (in any band); these cuts will be reviewed in Section 4.

In conclusion, there is little to be gained from using ALLMASK and we have therefore decided to use only NOBS as our pixel level mask, shrinking the area by 0.4 per cent and reducing the target density by 7.7 objects/deg² in BGS BRIGHT and 2 objects/deg² in BGS FAINT.

4 PHOTOMETRIC SELECTION

Following the spatial masking described in the previous section, the next step in the construction of the BGS target list is to incorporate information about photometric measurements into the selection process. According to the science requirements of the BGS and the mock BGS catalogues made by Smith et al. (2017), the survey is expected to have a target density of 800 galaxies deg⁻² to an r -band limit of 19.5. For the faint sample ($19.5 < r < 20$), which is second priority in BGS, a density of 600 galaxies deg⁻² is expected.

One of the major challenges for the BGS is the separation of stars and galaxies. In Section 4.1 we describe how we compare high angular resolution point source magnitudes from *Gaia* DR2 (Gaia Collaboration et al. 2018) with total magnitudes from the best-fitting light profile model selected by TRACTOR to distinguish point sources from extended sources.

In Section 4.2 we describe how we reject spurious objects that have incongruous light profiles by comparing their total magnitudes with the fibre magnitude that TRACTOR computes from the fitted profile assuming 1 arcsec Gaussian seeing and 1.5 arcsec fibre diameter. We place a cut in the fibre magnitude versus total magnitude plane that is motivated by the locus of confirmed galaxies from the GAMA DR4 survey.

Further posterior cuts which use photometry include removing colour outliers in $g-r$ and $r-z$ (see § 4.3), and applying quality cuts that indicate low accuracy in the flux measurement for an object (see § 4.4). The quality cuts make use of the quantities FRACMASKED, FRACFLUX and FRACIN measured by TRACTOR for each object in each of the three bands (grz). These are defined and discussed in § 4.4.

In Fig. 4 we show the second part of the BGS target selection flow chart. This flow chart focuses on the photometric selection cuts and starts from where the previous flow chart (Fig. 2), showing the spatial cuts, left off. The BGS catalogue, in the DECaLS subregion, ends up having a reduced area of 9401 deg² out of the initial 9717 deg², and target densities of 846 objects/deg² and 578 objects/deg² for BGS BRIGHT and BGS FAINT respectively.

4.1 Star-galaxy separation

The classification of images as star or galaxies is an old problem that is of great importance when defining target catalogues for the efficient use of multi-object spectrographs. Sophisticated techniques are employed which include algorithms using machine learning methods applied to both colour and morphological information e.g.

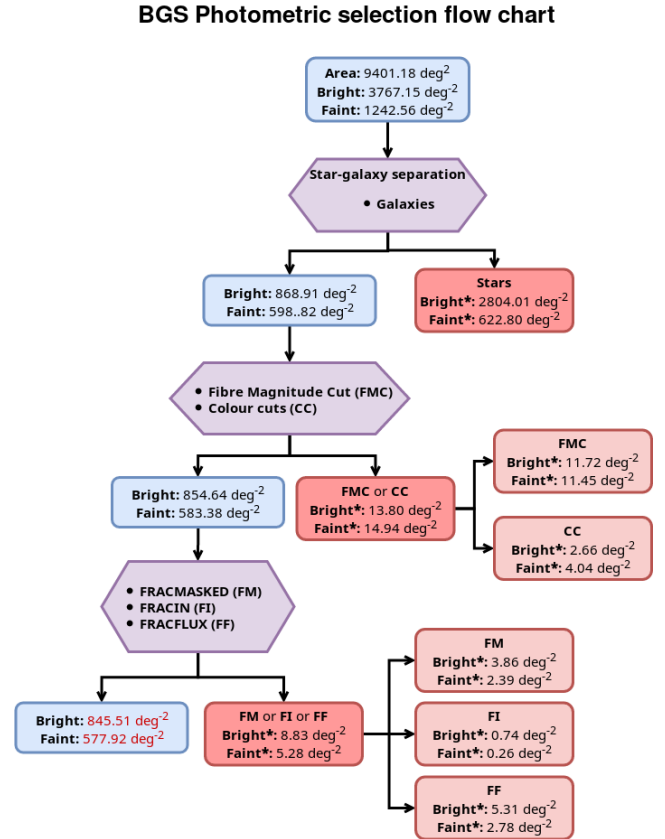


Figure 4. Flow chart of the BGS target selection in the Legacy Surveys DR8 based on photometric considerations. The photometric selection of BGS targets is divided into four stages; star-galaxy separation, fibre magnitude cut (FMC), colour cuts (CCs) and quality cuts (QCs). The photometric selection flow chart is a continuation of the spatial cut flow chart (Fig. 2) and therefore we start from the area and object densities reported at the end of the spatial cut flow chart. We report densities for the bright and faint samples separately, showing in blue boxes the values for the sources remaining after each of the BGS cuts. The densities of the removed objects are shown in red/pink boxes. The different cuts applied are shown in purple hexagonal boxes.

artificial neural networks (Odewahn et al. 1992; Bertin & Arnouts 1996), support vector machines (Fadely et al. 2012) and decision trees (Weir et al. 1995). TRACTOR uses a rigorous statistical approach to determine the best fitting light profile model to each object. In this way it classifies objects as either point sources (PSF) or extended sources (DEV, EXP, COMP or REX). However, this pipeline is not infallible and it is inevitable with ground based seeing that some compact galaxies will be misclassified as being of PSF type rather than extended. As we want to avoid incompleteness that depends on the variable seeing of the images we have instead made use of the space based high angular resolution *Gaia* photometry to distinguish point sources from extended sources. This is possible for the BGS as virtually²³ all stars brighter than the BGS magnitude limit of $r < 20$ are bright enough to be detected by *Gaia*.

The *Gaia* DR2 catalogue (Gaia Collaboration et al. 2018) that we use is primarily a catalogue of stars but has some galaxy and quasar contamination as reported by Bailer-Jones et al. (2019). This means we cannot simply classify all of the BGS objects that are

²³ *Gaia* DR2 is complete between $12 < G\text{-mag} < 17$.

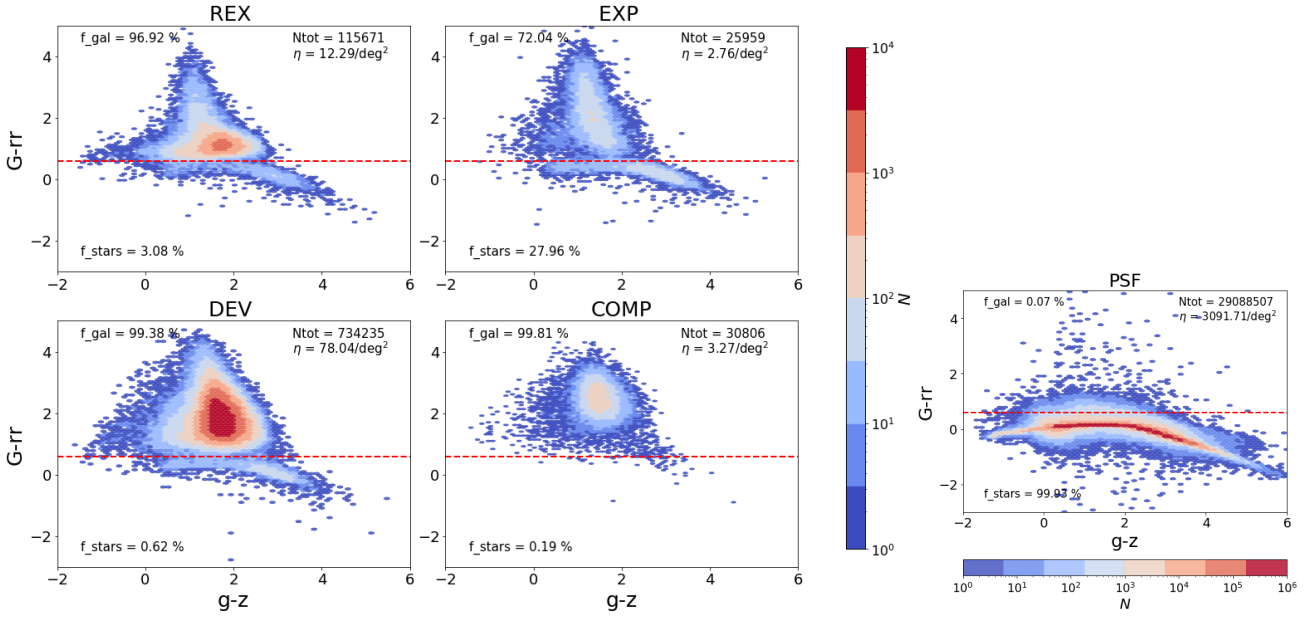


Figure 5. Separately for objects classified by TRACTOR as type REX, EXP, DEV COMP and PSF we show the difference between the *Gaia* (PSF) magnitude G and total non-dust corrected r -band model magnitude measured by TRACTOR, rr versus TRACTOR extinction corrected $g - z$ colour. All the objects plotted have passed the geometrical and pixel cuts detailed in Fig. 2, and all but the star-galaxy classification cut of the photometric-based cuts detailed in Fig. 4. The plots show objects that have been cross-matched between LS DR8 objects and *Gaia* DR2. Each panel shows a different morphological class, as labelled, according to the best-fitting light profile assigned by TRACTOR. The red-dashed line indicates our adopted division at $G - rr = 0.6$ with stars below and galaxies above the line. The colour in the plots shows the number counts of objects in an hexagonal cell, ranging from 1 to 10000, except for the case of PSF-type objects, in which case the colour scale covers the range from 1 to 1 million as indicated in the colour bars. We display the fraction of galaxies and stars according to this classification at the top-left corner and bottom-left corner respectively. The total number of objects (N_{tot}) in each plot and the target density (η) this represents is displayed in the top-right corner.

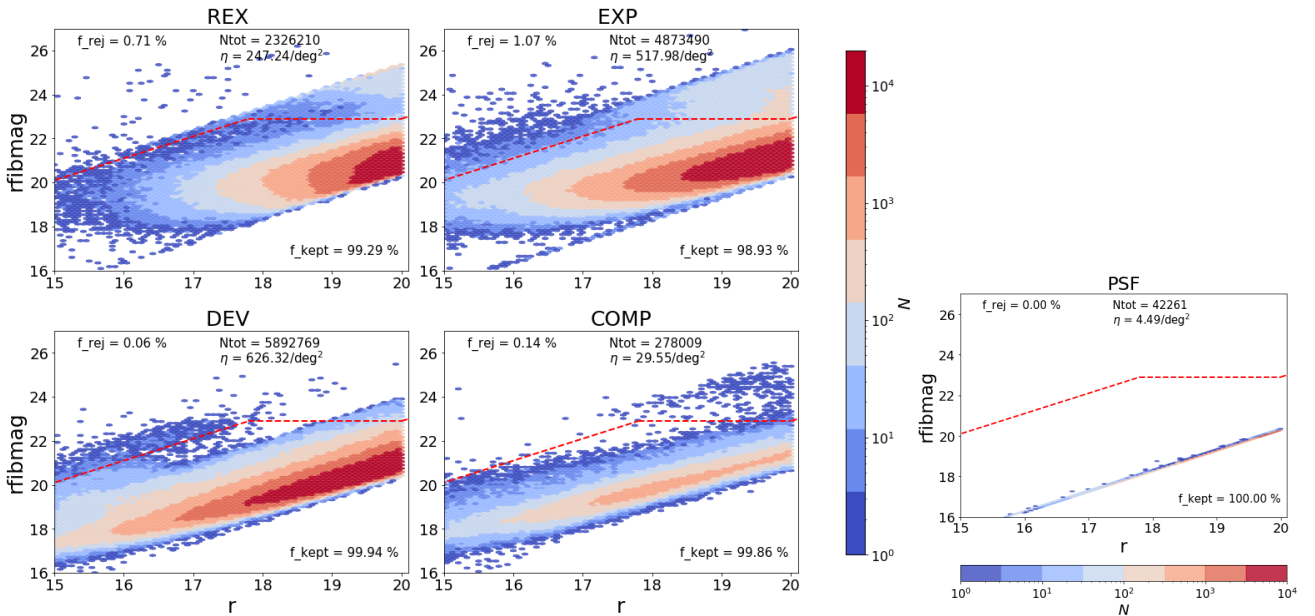


Figure 6. BGS galaxies in the r -band total magnitude (x-axis) versus r -band fibre magnitude (y-axis) plane in the LS DR8. The results are divided into the five different TRACTOR best-fitting light profile models, as labelled at the top of each panel. The colour bar shows the number counts of objects in an hexagonal cell covering the range from 1 to 20000 for four of the light profile models with the exception of PSF-type galaxies, in which case the scale covers 1 to 10000. The red-dashed line shows the fibre magnitude cut (FMC): we reject every object that is above this threshold. The numbers shown in top-left and bottom-right corners give the fraction of galaxies rejected and kept, respectively, while the number in the top-right corner shows the total number of galaxies (N_{tot}) and the corresponding target density (η).

in *Gaia* as stars. However, by comparing TRACTOR magnitude measurements with the higher spatial resolution magnitude measurements from *Gaia* we can determine which objects have extended light profiles. The *Gaia* magnitudes are computed assuming all objects are point sources. This results in accurate magnitudes for stars but magnitudes that are systematically fainter than the associated total magnitudes for sources that are extended compared to the ~ 0.4 arcsec PSF achieved by *Gaia*. In contrast, the model magnitudes computed by TRACTOR should capture more fully the total magnitude of the object. Consequently, if *Gaia* and TRACTOR magnitudes were measured in the same band, we would expect them to agree for point sources but for the TRACTOR magnitude to be brighter than the *Gaia* magnitude for extended sources. We would even expect this to be true for extended objects that TRACTOR misclassifies as PSF since the wide, ground-based PSF of TRACTOR would capture more of the total flux than the narrow PSF of *Gaia*. The complication is that the *Gaia* G band is a much wider filter than the DESI r band, but as we shall see, the colour dependence is weak.

Based on these considerations we define TRACTOR objects with $r < 20$ as being galaxies if either of the following two conditions is met:

- The object is not in the *Gaia* catalogue.
- The object is in the *Gaia* catalogue but has $G - rr > 0.6$.

In the above, the G -band is the G photometric *Gaia* magnitude and rr is the raw r -band magnitude from the LS DR8 *without* applying a correction for Galactic extinction. This choice is made because the *Gaia* magnitude is not corrected for Galactic extinction. The discussion above explains that G and rr magnitudes are measured in different effective apertures and so the quantity $G - rr$ should be thought of as a measure of how spatially extended an object is and not its colour. The first criterion above is satisfied by most (93 per cent) of the BGS objects. It leaves very little stellar contamination in the BGS, as essentially any star brighter than $r = 20$ is bright enough to be detected and catalogued by *Gaia*. The second criterion is required to keep the BGS completeness high by not rejecting galaxies that are in the *Gaia* catalogue.

In Fig. 5 we show the $G - rr$ versus $g - z$ plane for objects in *Gaia* DR2 that are matched with objects in the LS DR8. The panels show different objects as classified by the TRACTOR model fits (i.e., PSF, COMP, DEV, EXP, REX). The cross-matched objects have been subject to all the BGS cuts (i.e. both spatial and photometric) with the exception of the star-galaxy separation itself. For objects classified by TRACTOR as PSF-type, we can see the stellar locus around $G - rr = 0$ with a weak colour dependence. For the extended sources (i.e., COMP, DEV, EXP, REX), we see part of the galaxy locus²⁴ in the upper part of the plot, just above $G - rr = 0$.

From Fig. 5 we can see that the assignment of the best fitting TRACTOR model supports our *Gaia* classification using $G - rr > 0.6$, but we can still see some remnants of the stellar locus for objects that have not been assigned PSF-type by TRACTOR. For the objects classified PSF-type by TRACTOR we see in the right-most panel of Fig. 5 that 99.93 per cent fall on the stellar side of our $G - rr$ cut. For the objects classified by TRACTOR as the extended types (REX, DEV and COMP) the stellar contamination (i.e. objects with $G - rr < 0.6$) is at most 3.1 per cent. However, the contamination of the EXP-type objects is approaching 30 per cent.

²⁴ We have to remember that Fig. 5 only includes stars and galaxies that are cross-matched between LS DR8 and *Gaia* DR2.

The BGS target selection has the expected surface density after applying the star-galaxy separation. From the spatial cut flow chart in Fig. 4, we find a bright target density of 868.91 objects/deg² and a faint target density of 598.82 objects/deg². Rejected *Gaia* stars have a target density of 2 804.01 objects/deg² bright stars and 622.80 objects/deg² faint stars.

4.2 Fibre magnitude cut

In order to reduce the number of image artefacts and fragments of ‘shredded’ galaxies that would otherwise be classified as BGS targets we apply a cut on the fibre magnitude that is defined as a function of r -band magnitude as follows:

$$r_{\text{fibmag}} < \begin{cases} 22.9 + (r - 17.8) & \text{for } r < 17.8 \\ 22.9 & \text{for } 17.8 < r < 20 \end{cases} \quad (4)$$

where r_{fibmag} is the magnitude of the predicted r -band fibre flux and r is the total r -band magnitude, both extinction corrected. The location of this cut was guided by inspecting postage stamp images of a selection of the objects with the faintest fibre magnitudes with the aim of rejecting objects that appear to be artefacts while retaining nearly all of the genuine galaxies. In addition, at the bright end our threshold was guided by the location of spectroscopically confirmed GAMA galaxies, as discussed further in Section 5.1. Fig. 6 shows the distribution of the BGS objects in the r_{fibmag} vs. r_{mag} plane, with a separate panel for the different TRACTOR classes, and a red-dashed line indicating the location of the *fibre magnitude cut* (hereafter FMC). In the first four panels we can see that the galaxy locus has a tight core and, in general, is well below the FMC. The FMC removes 1.2 per cent of the objects classified as EXP and even smaller fractions of the other light profile classes.

All BGS objects in the PSF class lie on a stellar locus. Whether all these objects are stars or whether this is an artefact of TRACTOR only fitting the PSF model to *Gaia* sources with low astrometric excess noise (AEN) is revisited in Section 5.1, where we compare our classification with that of the GAMA DR4 survey. The stellar locus is also visible in the other photometric classes indicating there is some stellar contamination in our sample, but it is at a very low level.

In summary the adopted FMC rejects a further 23.17 objects/deg² of which 11.72 are in BGS BRIGHT and 11.45 are in BGS FAINT from the objects that have passed the previous cuts which include the rejection of stars by our star-galaxy classifier.

4.3 Colour cuts

An efficient way of rejecting further spurious targets from the BGS is to reject objects with bizarre colours. The limits we impose to reject outliers are:

$$\begin{aligned} -1 &< g - r < 4 \\ -1 &< r - z < 4. \end{aligned} \quad (5)$$

Fig. 7 shows the $g - r$ vs. $r - z$ colour-colour distribution of the objects retained in BGS if all but the colour cut (CCs) were applied. The red box indicates the colour range we keep. We can see from this plot that the locus of normal galaxy colours lies well within the range we retain and the cuts are only removing objects/artefacts with bizarre colours. It is evident that some stellar contamination remains as the stellar locus can be seen as a spur of objects with very red $r - z$ colours. However the density of objects in this spur, and its blueward extension which overlaps the galaxy locus, is no

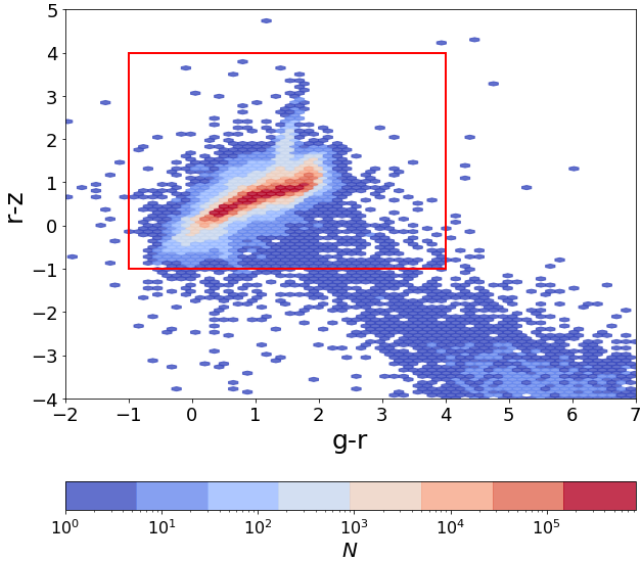


Figure 7. Colour-colour distribution showing $g - r$ vs. $r - z$ for BGS objects without applying the CCs. The colour bar shows the number counts of objects in an hexagonal cell covering the range from 1 to 800 000. The solid red box shows CCs defined in Equation 5. Sources outside of this box are excluded from the BGS.

more than a few objects/deg² as we shall see in Section 5.1. The colour cuts (CCs) we apply reject an additional 6.7 objects/deg², with 2.66 in BGS BRIGHT and 4.04 in BGS FAINT.

4.4 Quality cuts

Each object in the TRACTOR catalogue has three measures of the quality of its photometry recorded in each of the three bands (grz). These are:

- FRACKMASK (FM): The profile-weighted fraction of pixels masked in all observations of the object in a particular band. This quantity lies in the range $[0, 1]$. High values indicate that most of the flux of the fitted model lies in pixels for which there is no data due to masking and so the measurement is unreliable.
- FRACIN (FI): The fraction of the model flux that lies within the set of contiguous pixels (termed a ‘blob’) to which the model was fitted. FRACIN is close to unity for most real sources. Low values indicate that most of the model flux is an extrapolation of the model into regions in which no data was available to constrain it.
- FRACFLUX (FF): The profile-weighted fraction of the flux from other sources divided by the total flux of the object in question. FRACFLUX is zero for isolated objects but can become large for faint objects detected in the wings of brighter objects that are nearby.

Once the other cuts have been applied, in particular, the cut on NOBS and the BS mask, the distribution of each of these quantities is tightly peaked around the favoured values of FRACKMASK ≈ 0 , FRACIN ≈ 1 and FRACFLUX ≈ 0 . However, each quantity has a distribution with a fairly featureless tail that extends out to less desirable values. There are also clear correlations between the three quantities for a given photometric band and in some cases between photometric bands. The choice of the best set of thresholds to reject outliers is not trivial. We have adopted the following quality cuts

Table 2. The BGS target densities for each of the TRACTOR best-fitting photometric models. The first column labels the photometric model. The next three columns list the surface density of objects per deg² for the BGS BRIGHT and BGS FAINT samples separately and their combined sum. The area covered by the DECaLS portion of the BGS is 9, 401 deg².

Model	η_{bright} [deg ⁻²]	η_{faint} [deg ⁻²]	η_{overall} [deg ⁻²]
DEV	427	202	629
EXP	284	230	514
REX	104	141	246
COMP	27	3	31
PSF	3	2	5
Total	846	578	1423

(QCs):

$$\begin{aligned} \text{FRACMASKED}_i &< 0.4, \\ \text{FRACIN}_i &> 0.3, \\ \text{FRACFLUX}_i &< 5, \end{aligned} \quad \text{where } i = g, r \text{ or } z, \quad (6)$$

based on visual inspection of postage stamp images.

As mentioned in Section 3.2, we find that the objects flagged by the TRACTOR quantity ALLMASK are essentially a subset of the objects that are rejected by applying the quality cuts listed in Eqn. 6. While cutting on ALLMASK would have the advantage that it could also be applied to the randoms, we find that it is important to apply the QCs to remove spurious objects that are missed by the other cuts. For instance, some spurious objects that are outliers in either the fibermag vs. mag plane or in the colour-colour space that just pass the FMC and CCs are removed by considering FRACMASKED or FRACIN.

As shown in the flow chart, Fig. 4, the QCs reject an additional 14.11 objects/deg² of which ~ 60 per cent are removed by FRACFLUX, ~ 45 per cent by FRACMASKED and ~ 7 per cent due to FRACIN. The overlap between the FRACMASKED, FRACIN and FRACFLUX cuts is minimal, with only 1.05 objects/deg² for objects with $r < 19.5$, and in round 0.15 objects/deg² for objects with $19.5 < r < 20$ being rejected by more than one of the cuts. Separately for BGS BRIGHT and BGS FAINT, we show the target density of objects rejected by these cuts after applying all the previous cuts. The largest overlap between these cuts is between FRACMASKED and FRACFLUX for BGS BRIGHT, but even here it amounts to less than 1 object/deg². For BGS FAINT this overlap is small, 0.11 object/deg², and there is no overlap with FRACIN.

In Appendix A we present another version of the selection cut flow chart in which the cuts are applied in a different order. There we give a galaxy view of the target selection by first applying the star-galaxy classification so that all the subsequent cuts apply only to galaxies. The final selected sample which comprises of 845.5 galaxies/deg² in BGS BRIGHT and 577.9 galaxies/deg² in BGS FAINT, is exactly the same, as the order of the cuts does not matter. The objects rejected by each filter, however, does change as many objects are rejected by more than one filter. To illustrate this point we have also swapped the order of the FMC and QCs cuts so one can see how these influence one another.

5 CATALOGUE PROPERTIES

The final BGS catalogue in the DECaLS region in the South Galactic Cap (SGC) covers the declination range $-17 \lesssim \text{DEC} \lesssim 32$ degrees,

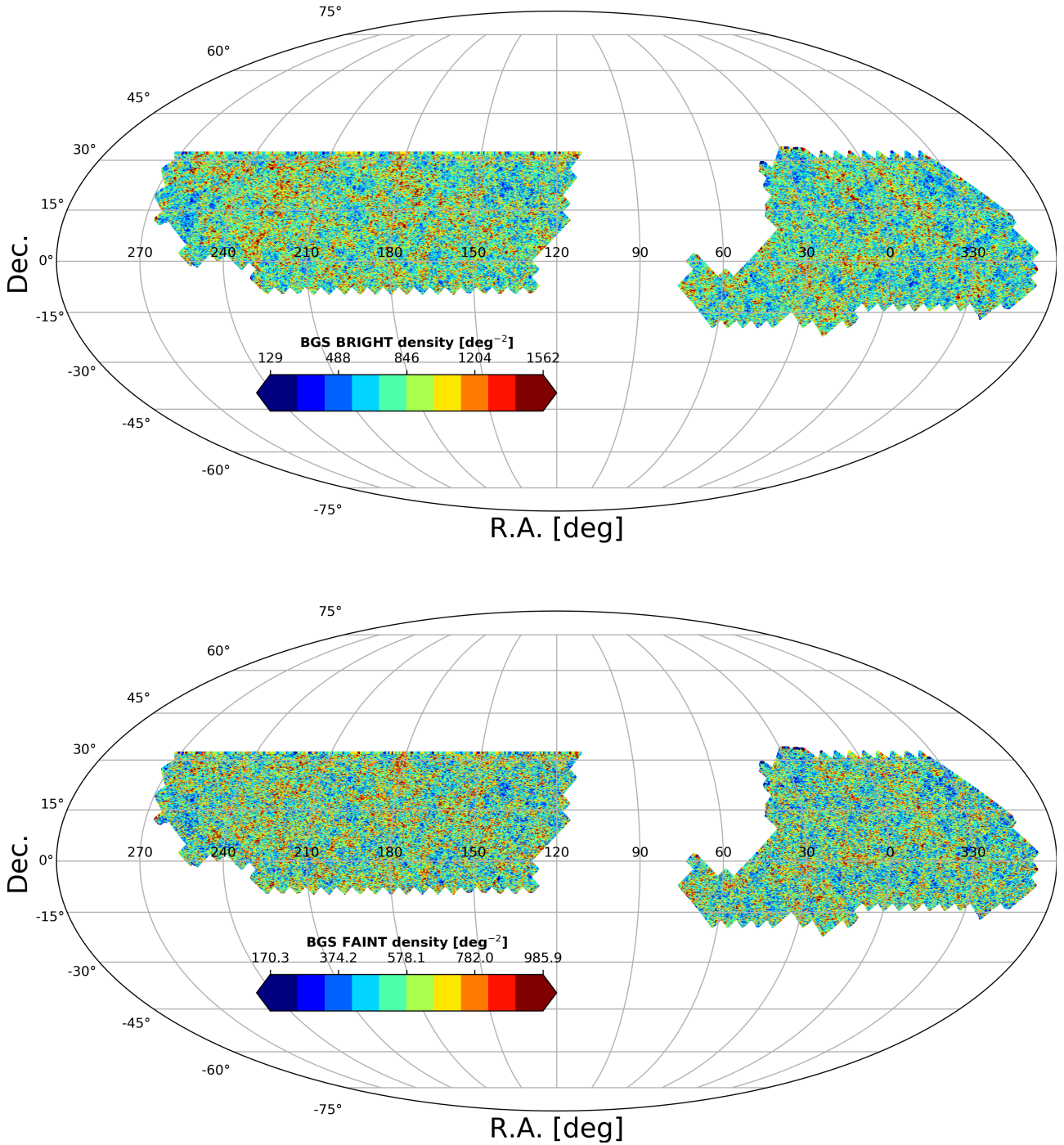


Figure 8. The distribution on the sky of the BGS BRIGHT (upper map) and BGS FAINT (bottom map) target density in objects/deg², computed on a HEALPix grid with a resolution of $N_{\text{side}} = 256$. The mean densities are 846 and 579 objects/deg² for the bright and faint BGS respectively.

and in the North Galactic Cap (NGC) the range $-10 \lesssim \text{DEC} \lesssim 32$ degrees. The BGS has a total of 13,378,062 galaxies of which 7,944,975 are in BGS BRIGHT and 5,433,087 are in BGS FAINT. The total area covered by the BGS in the DECaLS subregion defined by the footprint of the tiles in Fig.1 and after accounting for the spatial cuts described in Section 3 is 9401 deg². In Table 2 we list the target density of the BGS catalogue for each of the best-fitting photometric models used in TRACTOR.

In Fig. 8 we show the BGS BRIGHT and BGS FAINT sky

map densities computed with the HEALPix scheme using

$$\begin{aligned} \eta_i &= N_i^{\text{BGS}} / A_{\text{eff}}, \\ A_{\text{eff}} &= N_i^{\text{R}} / \eta^{\text{R}}, \end{aligned} \quad (7)$$

where for each pixel N_i^{BGS} is the number of BGS targets, A_{eff} is the effective area computed from the number of randoms, N_i^{R} , and the total surface density of the randoms, without any masking, is $\eta^{\text{R}} = 15,000$ objects/deg². We use a HEALPix grid of $N_{\text{side}} = 256$ giving a pixel area of $A_{\text{pix}} = 0.052$ deg². The appearance of the

density fluctuations is very similar in the two disjoint regions and show no variation with galactic latitude. We look more closely at systematic variations in the target density in Section 5.2.

5.1 Cross-comparison with GAMA

The main target sample in GAMA (Baldry et al. 2017) is a complete sample of galaxies with SDSS Petrosian r -band magnitude brighter than $r = 19.8$. The Petrosian magnitude is measured within a circular aperture of twice the Petrosian radius, where the radius is computed using the r -band surface brightness profile (Adelman-McCarthy et al. 2008). The GAMA photometric selection is very similar to that of DESI BGS and so we expect a very similar redshift distribution as GAMA which has median of $z = 0.2$ and a 90 percentile value of $z = 0.5$.

Star-galaxy separation in GAMA was conservative in that it aimed for very high completeness at the expense of some stellar contamination. These properties combined with its very high spectroscopic completeness (high quality redshift has been obtained for more 98.85 per cent of the GAMA targets) make it a nearly ideal "truth table" from which to assess the completeness of the BGS target selection and measure the expected redshift distribution of the BGS BRIGHT sample. Below we make use of GAMA to examine various aspects of our BGS catalogue. In Sec. 5.1.1 we compare the r -band photometry of the matched objects and determine the redshift distribution of the BGS galaxies that match with galaxies in the GAMA survey. Section 5.1.2 explores an issue related to TRACTOR only providing PSF photometry for some of the BGS galaxies. In Section 5.1.3 we assess incompleteness in BGS relative to GAMA and quantify how much is caused by each of the various geometric and photometric selections.

5.1.1 Magnitude definition and redshift distribution

We match the GAMA Main Survey DR4 galaxy catalogue (Driver et al. 2012; Liske et al. 2015; Baldry et al. 2017), which is defined by a Petrosian magnitude (R_{PETRO}) limit of $r = 19.8$, to the BGS target catalogue. We use a maximum linking length of 1 arcsec to match them. The mean separation of the matches we find is 0.093 arcsec with a 1σ dispersion of 0.091 arcsec. We focus on three of the five GAMA fields: G09, G12, G15. We omit G02 as this GAMA field is only partially within the DECaLS footprint, and G23 is far to the south. The redshift completeness of the main GAMA survey is extremely high in the sense that 98.85 per cent of the objects in the catalogue yield redshifts with a quality flag $\text{NQ} \geq 3$.

The GAMA spectroscopic redshifts can be used to reliably reject stars with a cut at $z = 0.002$. In what follows we restrict our GAMA catalogue to the spectroscopically confirmed galaxies (~ 98 per cent of the full catalogue). The area of each of the GAMA fields considered is 59.98 deg^2 which means that our matched sample has a total area of $\sim 180 \text{ deg}^2$. The overall density of sources that are cross-matched between BGS and GAMA galaxies is ~ 970 objects/ deg^2 with a mean redshift of $z = 0.224$.

For this matched catalogue, Fig. 9 compares the DR8 r -band total magnitude (r_{LS}) with the Petrosian r -band magnitude from GAMA (r_{GAMA}) by plotting $r_{\text{LS}} - r_{\text{GAMA}}$ vs r_{GAMA} . To see how this difference depends on galaxy morphology, we divide the LS galaxies into the five photometric classes assigned by TRACTOR. In each panel we show the fraction of matched galaxies in each TRACTOR model fit class; DEV and EXP classes together make up 80 per cent of the sample and the PSF class just 2.5 per cent.

We mark on the plot the $r_{\text{LS}} < 20$ limit of BGS, but note this has not been applied when defining the LS sample that was matched to GAMA.

Differences in the effective passbands of the r -band filters of the LS and SDSS result in offsets in $r_{\text{LS}} - r_{\text{GAMA}}$ of around -0.05 and -0.1 for blue and red galaxies respectively (Dey et al. 2019). One also has to consider the difference in magnitude definitions which contributes the more to this magnitude offset. To the extent that the best fit profiles accurately describe the actual light profiles of the objects, LS provides total magnitudes. In contrast, the SDSS Petrosian magnitudes used by GAMA quantify only the flux within twice the Petrosian radius (Blanton et al. 2001). The fraction of the flux within this aperture depends on the light profile. For EXP profile it captures 99.4 per cent, but for the DEV profile which, is more sharply peaked but with broader wings, only 82 per cent is captured. It is these differences in definition which largely drive the differences in median offsets we see in the DEV, EXP, REX and COMP classes. In all these cases the LS magnitude is brighter (more negative) than the GAMA magnitude with median offsets being -0.085 magnitudes for EXP and -0.188 magnitudes for DEV. In contrast for the PSF case the median $r_{\text{LS}} - r_{\text{GAMA}}$ is positive, which means that the LS PSF model magnitude captures less flux than the GAMA Petrosian magnitude. For true point sources we would expect these two magnitudes to be almost equal. The positive difference appears to happen because TRACTOR force fits PSF models to sources that are actually extended (deemed extended by our *Gaia* based star-galaxy separation) and consequently underestimates their fluxes. The reason this happens is discussed in Section 5.1.2.

If we take account of the scatter between the BGS and GAMA magnitudes we can use GAMA to assess the level of contamination in the BGS catalogue. If we treat GAMA as being a 100 per cent complete galaxy catalogue then any objects in BGS that are not in GAMA would be contamination in the form of stars or image artefacts. This is not true at $r = 20$ as here some BGS objects will not be in GAMA simply because of the $r_{\text{petro}} < 19.8$ magnitude limit in GAMA. This can be seen in Fig. 9 from the location of the $r_{\text{LS}} = 20$ dashed line relative to where the GAMA data truncates at $r_{\text{GAMA}} = 19.8$. To avoid this problem if we apply a brighter magnitude limit $r < r_{\text{lim}}$ to BGS then for a broad range of $18.5 \leq r_{\text{lim}} \leq 19.3$ we find that ~ 3 per cent of BGS objects are not matched with GAMA galaxies. This sets an upper limit (in this magnitude range) of 3 per cent contamination in BGS as GAMA itself may not be 100 per cent complete.

Fig. 10 shows the distribution of redshifts for BGS objects that have been cross-matched with GAMA galaxies. The overall distribution is shown along with those for the BGS FAINT and BGS BRIGHT. We expect this distribution to be representative of the BGS BRIGHT sample as we can see from Fig. 9 that incompleteness caused by the GAMA magnitude limit to be very small. However the redshift distribution plotted for BGS FAINT is more strongly affected by the GAMA magnitude limit and its true redshift distribution is expected to be more extended.

5.1.2 Galaxies with TRACTOR type PSF

To avoid stars being classified as extended sources TRACTOR uses a catalogue of stars from *Gaia* to pre-select a set of objects on which it will only allow PSF fits. The *Gaia* objects for which it does this are based on the following cut on the *Gaia* astrometric excess noise

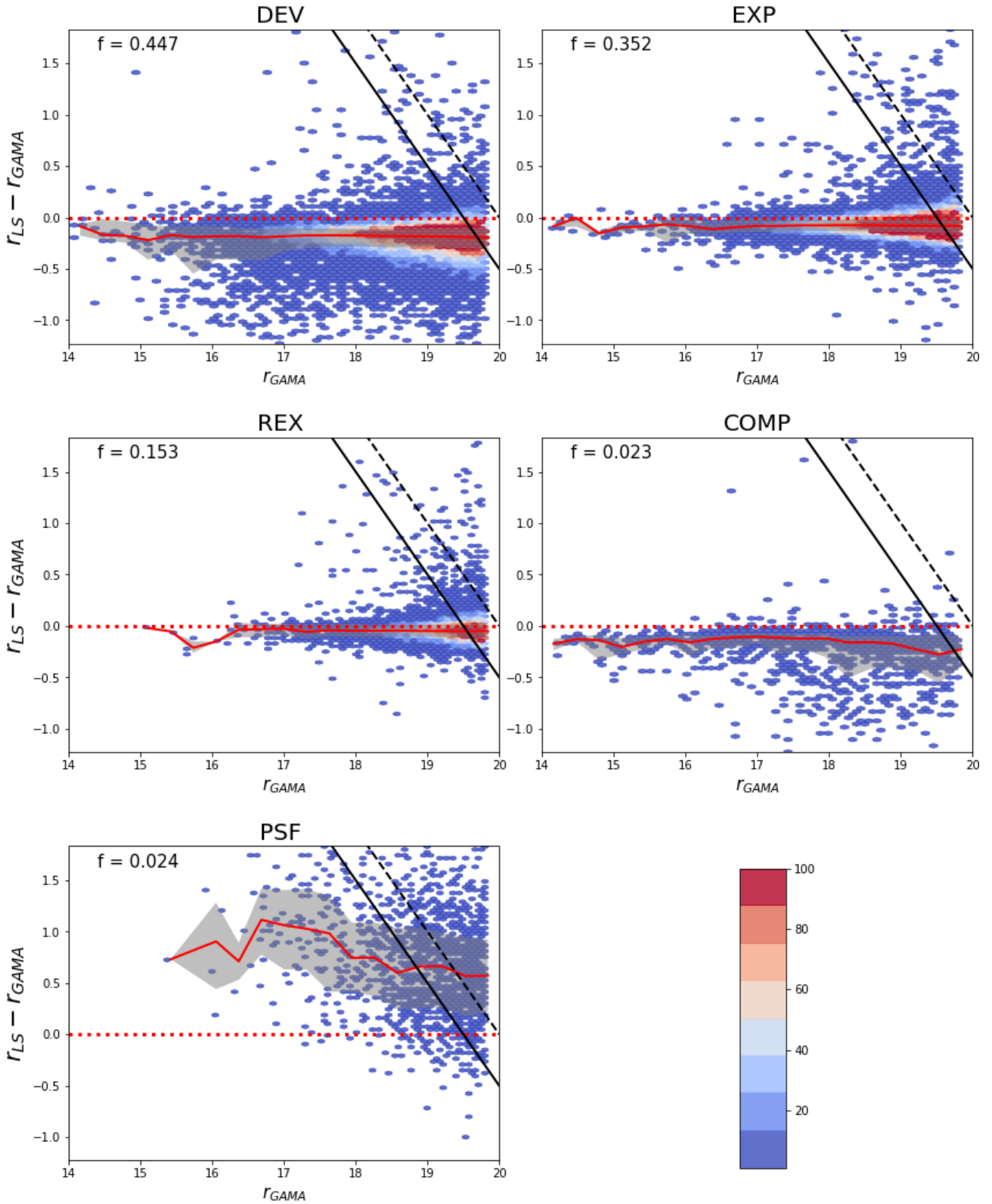


Figure 9. The r -band total magnitude in the LS (r_{LS}) vs the SDSS r -band Petrosian magnitude in GAMA (r_{GAMA}) for LS DR8 objects cross-matched with GAMA. Each plot corresponds to one of the five photometric model fits assigned by TRACTOR. The red solid line shows the median value of $r_{LS} - r_{GAMA}$ as function r_{LS} ; the gray shading shows the 20 to 80 percentile range; the dashed black line shows the limiting magnitude of $r_{LS} = 20$ for BGS and the solid black line shows limiting magnitude of $r_{LS} = 19.5$ for BGS. The colour bar shows the number counts of objects in an hexagonal cell covering the range from . The fraction of LS DR8 objects plotted out of the total number matched with GAMA is shown in the top-left corner of each panel.

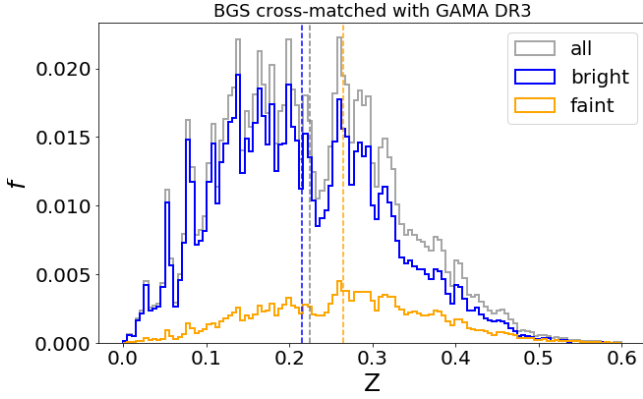


Figure 10. The redshift distribution of BGS objects cross-matched with GAMA DR4 broken into bright ($r < 19.5$, blue) and faint ($19.5 < r < 20$, orange) galaxies according to the BGS r -band. The gray histogram shows the overall redshift distribution of BGS galaxies cross-matched with GAMA. The mean redshift values for each distribution are: 0.215 for the bright sample (dashed blue line), 0.265 for the faint sample (dashed orange) and 0.224 for all galaxies (dashed gray).

parameter, AEN,

$$\begin{aligned} \text{AEN} < 10^{0.5}, \quad G \leq 19 \\ \text{AEN} < 10^{0.5+0.2(G-19)}, \quad G \geq 19, \end{aligned} \quad (8)$$

where G is the *Gaia* photometric G -band. The AEN can be used as measure of whether a source is extended as for extended sources the astrometric measurements are noisier than one would expect for a point source.

In contrast, in the BGS we use the difference between the *Gaia* G -band magnitude and the TRACTOR raw r -band magnitude, rr , (not corrected for extinction) as a measure of how extended the object is (see Section 4.1). In Fig. 11 we have plotted $\log(\text{AEN})$ versus G separately for objects classified as stars and galaxies by our $G - rr$ classifier. The threshold adopted by TRACTOR can be seen to separate the bulk galaxies from the stars. For 96 objects/deg² the two criteria agree the object is a galaxy, but the distributions are extended and the agreement is not perfect. There are 36 objects/deg² that the AEN criterion classifies as galaxies which $G - rr$ classifies as stars. More problematic are the 5 objects/deg² that the AEN criterion classifies as stars which $G - rr$ classifies as galaxies. This is an issue as it means some objects that are classified as galaxies in the BGS are treated by TRACTOR as stars and only have a PSF light profile fitted. Overall in the BGS there are 5 objects/deg² with PSF type within the DECaLS footprint (see Table 2). These objects have fibre magnitudes that are consistent with the locus of stars in Fig. 6 which makes us question if they really are galaxies. We investigate this below by making use of GAMA to determine whether or not they are galaxies.

First, we restrict our attention to the 180 deg² of our matched GAMA catalogue. The BGS PSF-type galaxies (main sample) have a density of 4.10 objects/deg², somewhat less than the 5 objects/deg² which is the average over the full DECaLS area. This reduces further to 1.76 objects/deg² after cross-matching with GAMA. We further subdivide these two cases (BGS PSF type and BGS PSF type cross-matched with GAMA) into three disjoint sub samples: i) those that are not in *Gaia*, ii) those that are in *Gaia* and which are classified using the AEN value as stars, and iii) those that are in *Gaia* and which are classified using the AEN value as galaxies.

Table 3. The surface density of PSF-type objects in the BGS in the G09, G12 and G15 GAMA fields combined before (η_{BM}) and after (η_{AM}) cross-matching with GAMA (top half of table). The bottom half of the table shows the surface density and percentage of objects in disjoint subsamples of the PSF-type BGS sample, as listed in the first column: objects that are not in *Gaia*, objects that the AEN scheme classifies as stars and those that the AEN scheme classifies as galaxies.

Sample	η_{BM} [deg ⁻²]	η_{AM} [deg ⁻²]		
PSF-type BGS	4.10	1.76		
Subsample	η_{BM} [deg ⁻²]	% _{BM}	η_{AM} [deg ⁻²]	% _{AM}
Not in <i>Gaia</i>	1.72	42.0	0.04	2.3
<i>Gaia</i> AEN star	2.26	55.2	1.69	96.4
<i>Gaia</i> AEN galaxy	0.11	2.8	0.02	1.3

The subsample sizes are reported in Table 3, where we give the surface density of objects before and after the cross-match with GAMA (η_{BM} and η_{AM}) along with the percentage of the total number of objects represented by each subsample. This shows that ~ 96 per cent of the BGS PSF-type cross-matched with GAMA are *Gaia* AEN stars, which represents the ~ 55 per cent in the non-matched sample. For the remained 45 per cent in the non-matched sample, GAMA is not reliable to assess this as only 3.6 per cent of those are cross-matched with GAMA. Fig. 12 shows the GAMA redshift distribution for the BGS PSF-type cross-matched with GAMA broken into the three classes shown in Table 3. These objects shown a redshift distribution very similar to that of the full BGS sample. The reason for this mis-classification lies in the fact that for objects classified by the *Gaia* AEN criterion as stars TRACTOR only fits PSF models. For the galaxies that this *Gaia* AEN criterion falsely classifies as stars TRACTOR underestimates the total flux of the galaxy resulting in the offset with the GAMA photometry we saw in the PSF panel of Fig. 9 and putting these galaxies close to the stellar locus in Fig. 6.

5.1.3 Incompleteness of BGS relative to GAMA

To the depth of GAMA we can assess the completeness of the BGS catalogue by cross-matching the full depth LS DR8 catalogue with GAMA DR4. This cross-match yields a catalogue of 1011 objects/deg² which represents of 99.6 per cent of the GAMA catalogue. Visual inspection reveals some of the remaining 0.4 per cent are debundling issues where GAMA fragments a galaxy into two objects while TRACTOR keeps it as a single object. Of the matched objects 970 objects/deg² are in BGS while the other 41 objects/deg² are excluded from the BGS catalogue by one or other of our selection cuts.

Due to the scatter between SDSS r -band Petrosian magnitude used by GAMA and the TRACTOR model magnitude used by BGS (see Fig. 9), the BGS $r_{\text{LS}} = 20$ magnitude limit excludes 20 faint GAMA galaxies per square degree. This leaves 20.8 objects/deg² in GAMA that are missing from the BGS. Whether this represents potential problematic incompleteness in BGS or just a difference in sample definition depends on which selection cuts remove the objects. We quantify and discuss this below.

The diagonal elements in Fig. 13 indicate the number density of spectroscopically confirmed GAMA galaxies missing from the BGS catalogue as result of each of the following spatial and photometric cuts: the bright star mask (BS); the large galaxy mask (LG); the number of observations (NOBS); star-galaxy classification (SG); fibre magnitude cut (FMC); colour cut (CCs); the FRAC-

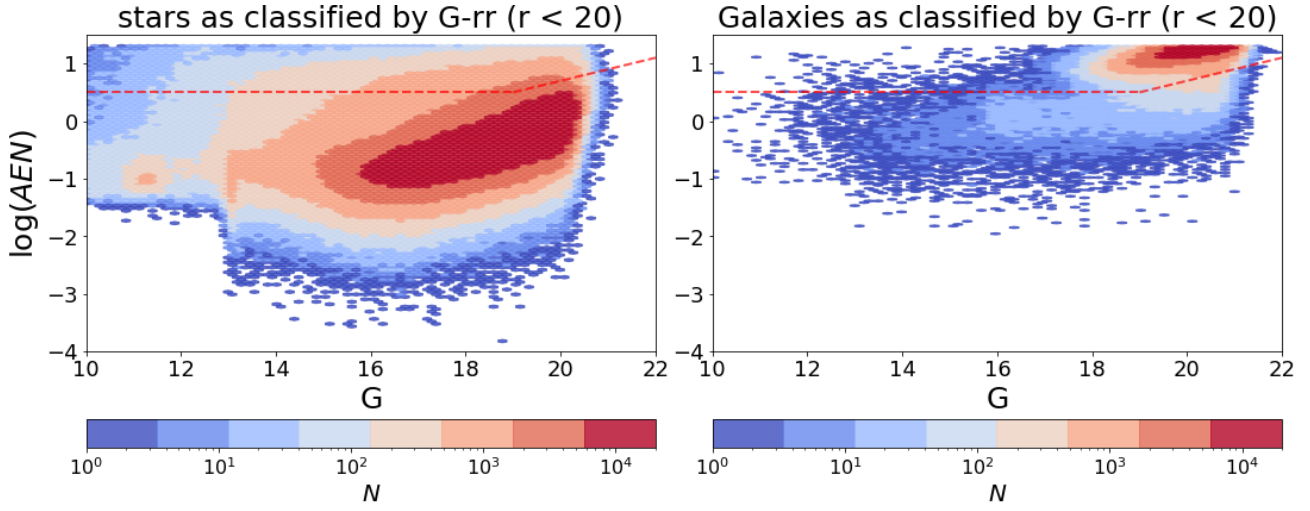


Figure 11. The *Gaia* Astrometric Excess Noise parameter (AEN) versus *G*-band magnitude. The top panel shows *Gaia* objects classified as stars by BGS and the bottom those classified as galaxies. Both plots only show *Gaia* objects with magnitudes limit of $r < 20$. The red dashed-line represents the threshold limit for the AEN classification used in TRACTOR, therefore everything below the line is a star and everything above is a galaxy according to the AEN classification. The colour bar shows the number counts of objects in an hexagonal cell covering the range from 1 to 20 000.

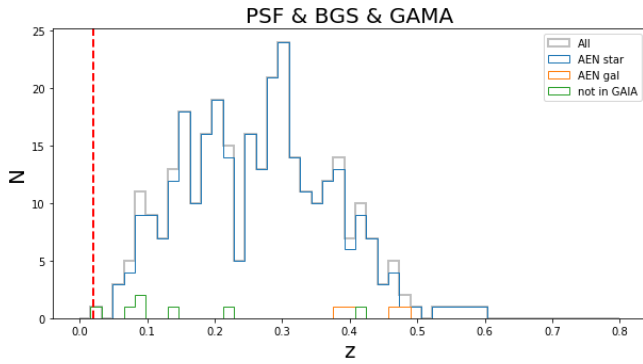


Figure 12. Redshift distribution of PSF-type BGS galaxies cross-matched with galaxies from three GAMA fields (G09, G12, G15). Redshifts are taken from GAMA DR4. The four distributions correspond to the matched sample (gray) and the disjoint subsamples comprising galaxies not in *Gaia* (green), and stars (blue) and galaxies (red), as defined by the AEN classification. The red dashed line marks the redshift $z = 0.002$; objects with redshifts smaller than this are stars.

MASKED quality cut (QCs FM); the FRACIN quality cut (QCs FI); the FRACFLUX quality cut (QCs FF). The off-diagonal entries in Fig. 13 show the surface density of GAMA galaxies that are removed by both of the two cuts indicated by the labels on the x and y axes.

The objects removed by the spatial BS and NOBS cuts are benign in that they do not affect BGS clustering measurements. These spatial masks are uncorrelated with BGS galaxy positions and so can be fully accounted for in clustering analyses by applying the same masks to the random catalogue. The values given in Fig. 13 show that these two masks have no overlap and together remove $9.36 \text{ objects/deg}^2$.

Applying these two spatial cuts leaves us with $11.43 \text{ galaxies/deg}^2$ that are in GAMA but are missed by BGS. The cuts that remove these objects are almost completely independent. $5.36 \text{ objects/deg}^2$ are removed by the our SG classification. These objects are close to the cut imposed for the

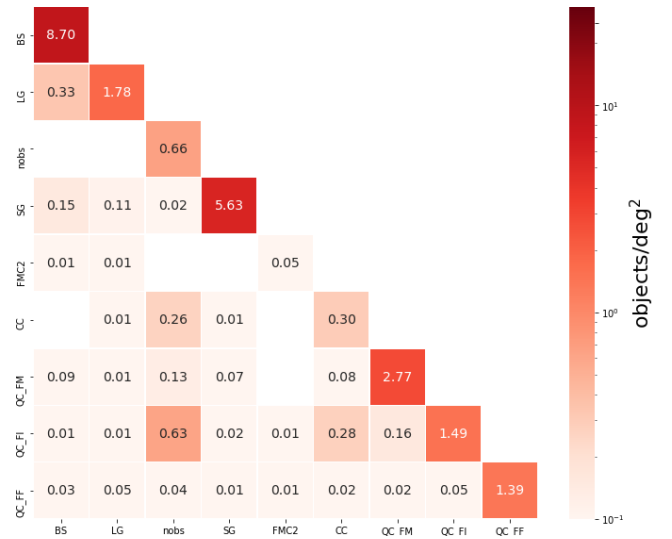


Figure 13. Heatmap showing the target density of GAMA galaxies ($z > 0.002$) that are missed in the BGS. The diagonal shows the number of objects per square degree removed by each of the individual spatial and photometric cuts applied in the BGS while the off-diagonal entries show the densities of objects removed by both cuts labelled on the x and y axes.

Gaia star-galaxy separation ($G - rr = 0.6$), but fall on the stellar side. We find that 98 per cent of these missed GAMA galaxies are classified as stars according to the *Gaia* AEN condition, which means that their photometry has been compromised as TRACTOR only fitted PSF models. If these are extended objects, then their flux as reported by TRACTOR is a fraction of what it should be and hence their rr -magnitude is shifted to fainter values. This results in BGS galaxies shifting to lower values of $G - rr$, moving them out of the galaxy locus and into the stellar one. If this were fixed we would expect the residual incompleteness to be $6.07 \text{ galaxies/deg}^2$, equivalent to $6.07/970 = 0.62$ per cent. The proportions of this produced by the LG QCs FM, QCs FI and QCs FF cuts are 23.5,

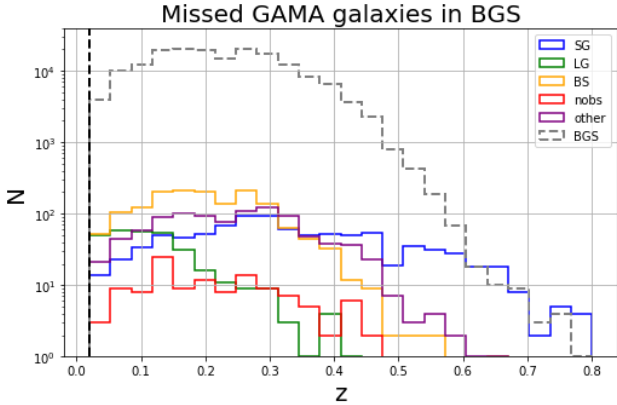


Figure 14. Redshift distribution of the GAMA galaxies that are not included in the BGS, with objects rejected by different cuts indicated by different line colours as labelled: blue shows GAMA objects missed due the star-galaxy separation applied (SG), green due to large galaxy masking (LG), yellow – bright star masking (BS), red – number of observations (NOBS) and purple due to the remaining cuts (CCs, FMC and all the QCs). The dashed gray line shows the redshift distribution of BGS galaxies cross-matched with GAMA. The vertical black dashed line marks the redshift boundary between stars ($z < 0.002$) and galaxies.

41.2, 13.8 and 21.4 per cent respectively with a negligible fraction removed by the CCs and FMC.

In Fig. 14 we show the redshift distribution of the GAMA galaxies that are not present in the BGS. The solid coloured lines show the distribution for GAMA galaxies rejected by different BGS cuts, as labelled in the figure. We also plot the overall redshift distribution of BGS galaxies for comparison. GAMA galaxies removed by the bright star masking and by the restrictions on the number of observations have a similar redshift distribution to the overall BGS. GAMA galaxies that are removed by the large galaxy mask have a distribution that is shifted to lower redshifts than the overall BGS distribution. GAMA galaxies can be found within the geometric BGS mask as GAMA does not use masking to deal with large galaxies, and so GAMA galaxies can be found in the regions that the BGS rejects around large galaxies. However, GAMA does perform masking around bright stars but this is less aggressive than the LS DR8 bright star masking. This can be seen from the areas rejected: the bright star masking in GAMA removes ~ 1 object/deg² (Baldry et al. 2010) whereas LS DR8 removes ~ 5 objects/deg².

5.2 Potential systematics

Here we look at potential systematic effects that could influence the homogeneity of the BGS catalogue and show how to mitigate these. As in any survey, the density of BGS targets is affected by observational effects which arise for a number of reasons. These include astrophysical foregrounds such as Galactic extinction, variations in the density of stars in the Milky Way, as well as variations in depth for the different imaging surveys and uncertainties in the data calibration.

To study the impact of these systematics on the observed galaxy density, we use a HEALPix map that divides the whole sky into $12N_{\text{side}}^2$ equal area pixels, adopting $N_{\text{side}} = 256$. Each pixel contains the median value of the systematics values within the pixel and the

BGS target density. The corresponding BGS target density in each pixel, η_i , is defined in Equation 7.

We study the effect of eight systematics on the BGS target density:

- Stellar density: we use stars from the *Gaia* DR2 catalogue with $12 < G < 17$ to construct the stellar density in each HEALPix pixel.
- Galactic extinction: the extinction values were computed using the SFD98 dust maps as reviewed in Section 2.1.
- PSF size (seeing) in the *grz* bands: the PSF size measures the full width at half maximum (FWHM) of the point spread function (PSF) which determines how much the transmission of light through turbulence in the Earth’s atmosphere blurs the observed images. The seeing varies across the multiple observations.
- Photometric depth in the *grz* bands: the depth of the photometry, as characterised by the 5σ AB magnitude detection limit for a 0.45 arcsec round exponential galaxy profile, varies across the survey due to changes in the observing conditions.

To determine if the BGS target density has a systematic dependence on any of these quantities, we bin the HEALPix pixels according to the value of the quantity and for each bin determine the mean target density, η_i , and the error on the mean, $\sigma_i/\sqrt{N_i}$. In Fig. 15 we show how the mean BGS target density, η , varies with respect to each of the quantities listed above. Each panel shows the mean and error on the mean for three samples, BGS BRIGHT, BGS FAINT and the combined BGS sample (labelled bgs_any). The histogram below the curves in each panel shows (on an arbitrary scale) the number of HEALPix pixels contributing to each estimate. In general, the systematic variation in the BGS target density is less than 5 per cent, with the one exception being a ~ 7 per cent decrease in the target density in regions of high stellar density.

Stars could impact the BGS target density in at least five ways: i) Stellar contamination of the BGS selection could lead to increased target density in regions of the sky with high stellar density. ii) While the impact of very bright stars is dealt with by masking (see Section 3.1.1), the halos and diffraction spikes around slightly fainter stars could still affect the photometry of neighbouring galaxies. iii) High stellar density could lead to an overestimate of the local sky brightness which, when subtracted, would lead to fainter galaxy fluxes and hence a lower BGS target density. iv) star/galaxy superposition. v) Binary stars that TRACTORs resolution is not capable of resolving.

Stellar contamination would lead to an increase in target density with increasing stellar density, whereas we see a decrease that sets in above a stellar density of 10^3 deg⁻². Hence, stellar contamination cannot be the dominant systematic influence on the target density.

Galaxy photometry directly compromised by nearby stars that were not subject to masking also seems unlikely to be the cause for the variation in target density. We test this by implementing the medium bright stars mask with a very little impact on target density and clustering. A further masking with 2 and 3 times the masking radius of equation (3) was also tested with no improvement on target density at high stellar densities.

The effect of high stellar density on the estimation of the sky levels deserves further investigation, but is deferred to another study. There is some variation of the target density with galactic extinction which could indicate systematic errors in the estimation of the amount of dust extinction. However, as there are spatial correlations between stellar density and dust extinction, these trends could be driven by the variation in stellar density and can be mitigated with several techniques such as linear and non-linear regressions and

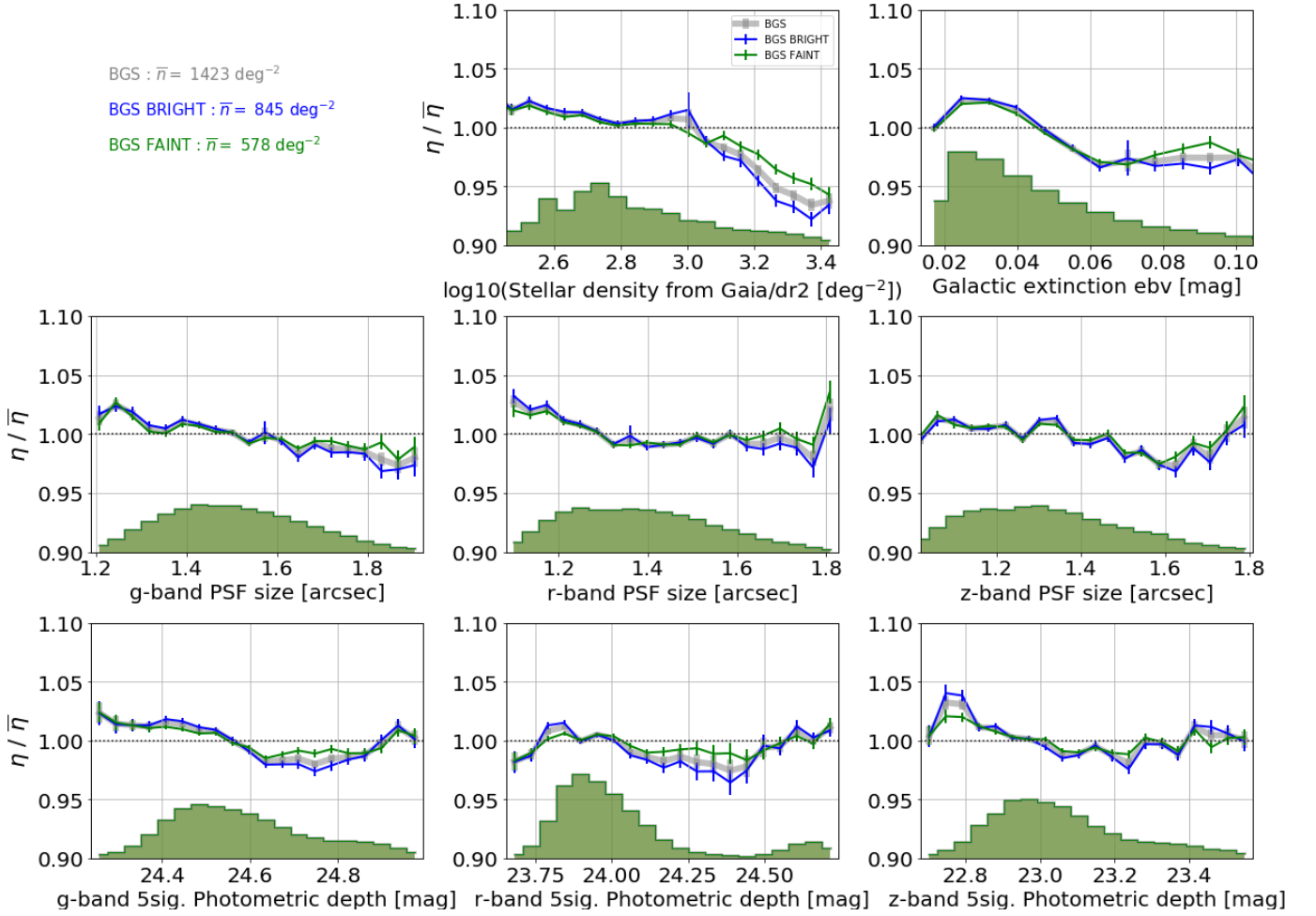


Figure 15. The systematic variation of the BGS BRIGHT (blue) and BGS FAINT (green) and combined (bgs_any, gray) target densities with respect to different properties: the logarithm of the stellar density from *Gaia* DR2, Galactic extinction, PSF size in the three bands (*grz*) and the photometric depth in each of the three bands (*grz*). The target densities and these eight quantities were computed in pixels on the sky using a HEALPix grid with resolution of $N_{\text{side}} = 256$. Histograms show the distribution for each of the x-axis properties. The error bars show the errors on the mean. Each target density, η is expressed in units of its mean across the whole survey $\bar{\eta}$ as given in the legend.

machine learning techniques such as Artificial Neural Networks (Rezaie et al. 2020).

Due to variations in observing conditions, the PSF size varies across the survey. The explicit modelling of the PSF of each image by TRACTOR should make the photometry robust to this variation. Also, our use of *Gaia* to perform star-galaxy separation should also make this classification independent to variations in the seeing. This appears to be borne out by the results shown in Fig. 15 which exhibit only very weak trends with PSF.

In the BGS, while the primary selection is in the *r*-band, TRACTOR simultaneously fits objects in all 3 bands and so the model parameters are affected by data in all three bands. However, any dependence on the depth of the photometry appears very weak in all three bands. This is to be expected as the photometric depth is typically 3 to 4 magnitudes deeper than the $r = 20$ selection limit of the BGS.

5.2.1 Mitigation of systematics using linear weights based on stellar density

One way to mitigate the effect of the systematics in our catalogue is to apply a weight that corrects the target density. If we treat the systematic dependence of the observed target density on a particular quantity, S , as a simple regression problem, we can define the observed target density, η_i^o , averaged over HEALpix pixels with a particular value of $S = S_i$, as

$$\eta_i^o = \eta_i W_i(S_i). \quad (9)$$

Here, η_i is the true target density and $W_i(S_i)$ is the weight for a given systematic attribute, S . As shown in Fig. 15, the most important target density variation is driven by stellar density. Here, we assume that the weight is a simple linear function, $W_i(S_i) = mS_i + c$, where S_i is the stellar density, as we would expect any contamination (or anti-contamination) to be proportional to the stellar density and not to the \log_{10} (stellar density). The best fitting coefficients we find when applying this model to the combined BGS BRIGHT and BGS FAINT sample are $c = 1.03$ and $m = -3.96 \times 10^{-5}$. By construction, this weighting removes the general trend with stellar density for the

combined sample and most of the trend with stellar density for the individual BGS BRIGHT and BGS FAINT samples. At the same time this weighting also reduces the weak systematic trend of target density with galactic extinction.

5.3 Angular correlation function

We measure the angular correlation function, $w(\theta)$, in five apparent magnitude bins from $r_{AB} = 15$ to $r_{AB} = 20$ for the BGS targets in DECaLS South Galactic Cap (SGC) and North Galactic Cap (NGC). Angular correlations were computed using the publicly available code CUTE (Alonso 2012). We compare these with measurements from the mock BGS lightcone catalogue (Smith et al. 2017). This mock catalogue was built by populating the MXXL N-body simulation with galaxies based on a halo occupation distribution model. By construction, the HOD parameters of this mock reproduces both the luminosity function and 2-point clustering measured in the SDSS at low redshift and the GAMA survey at higher redshift.

Fig. 16 shows the comparison of angular clustering measured for the BGS targets with error bars corresponding to the standard deviation of 100 jackknife realisations, the MXXL mock and the SDSS observations by Wang et al. (2013). The angular clustering measurements are consistent between the DECaLS North and South regions, which demonstrates the homogeneity between these two parts of DECaLS. The angular clustering of the BGS targets agrees very well with that displayed in the MXXL lightcone. The HOD parameters of the MXXL mock have been tuned to attempt to match the clustering measured from SDSS MGS, however on large scales HOD models can only alter the amplitude and not the shape of the correlation. Moreover the shape of the large scale correlation function of MXXL is very similar to that of all LCDM models that are consistent with CMB observations. Hence it is interesting that for the two faintest bins BGS is more consistent with MXXL (and hence with LCDM) than is SDSS MGS – possibly indicating reduced systematic errors.

We also look at the angular clustering of the BGS targets after applying the weights that depend on stellar density, as described in the previous section. Overall, applying stellar density weights has a small impact at angular scales larger than 3 – 4 deg. Both the clustering with and without the weights are consistent with each other, within the error bar.

A further test of the fidelity of our BGS catalogue is to check for any spatial correlation of the distribution of BGS targets with stars in the Milky Way. Here we focus our attention on the fainter stars, $12 < G < 17$, which, ideally, should be removed from the BGS targets by our star-galaxy separation scheme. We find a significant anticorrelation on very small scales but no correlation on scales larger than 100 arc seconds.

5.4 Angular cross-correlation with large galaxies

In order to determine whether we are missing faint BGS targets around large galaxies due to the LG mask defined in Section 2.2.4, we measure the angular cross-correlation function between the SGA-2020 and faint BGS targets in $18 < r < 19$ (dash-dotted) as shown in Fig. 17. We also measure the angular cross-correlation function between these faint BGS targets and brighter BGS targets in the magnitude range $15 < r < 16$ (solid) where we assume that most of the large galaxies lie, and we do the same using the MXXL lightcone (dashed). The vertical dotted line shows the mean mask radius around large galaxies, which is about 10 arcsec.

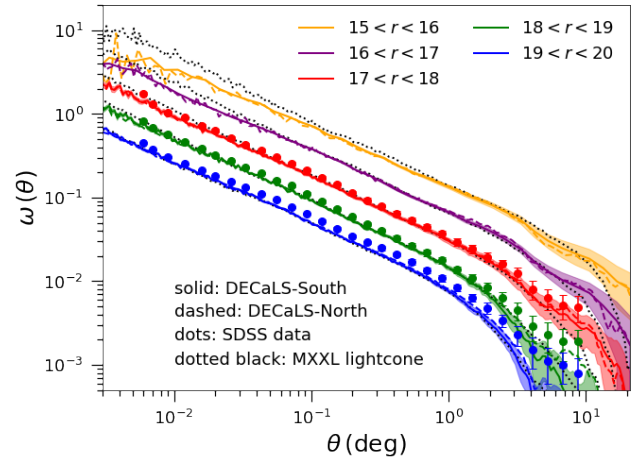


Figure 16. The angular correlation function, $w(\theta)$, measured for the BGS targets in bins of apparent magnitude; different colours indicate different magnitude bins as labelled. The shaded area shows the standard deviation obtained from 100 jackknife regions. The solid curves show the results for DECaLS-South, the dashed curves show DECaLS-North and the dotted curves show the angular clustering in the MXXL lightcone catalogue. The symbols with error bars show measurements from the SDSS by Wang et al. (2013).

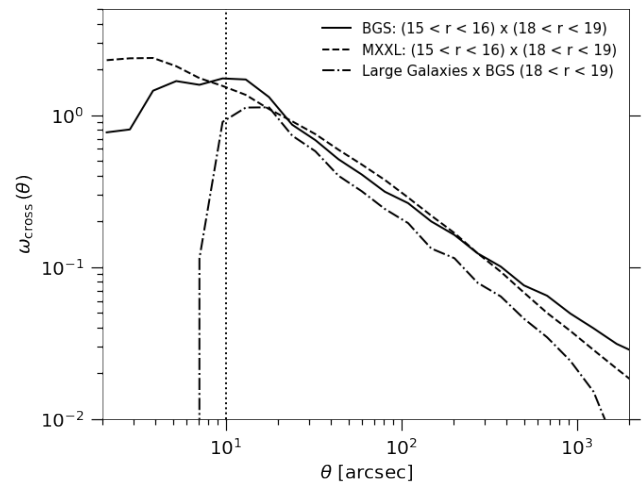


Figure 17. The angular cross-correlation function measured between faint BGS targets in $18 < r < 19$ and large galaxies from the SGA-2020 (dash-dotted) and between the same faint BGS targets and brighter BGS targets in $15 < r < 16$ (solid), the magnitude range in which most of the large galaxies reside. We also compare with the angular cross-correlation between these two bins in apparent magnitude measured in the MXXL lightcone (dashed). The vertical dotted line shows the mean LG mask radius which is about 10 arcsec.

The agreement between the results from the BGS catalogue (solid) and from the MXXL lightcone (dashed) suggests that our treatment of large galaxies is satisfactory and we are only missing BGS targets on scales below 10 arcsec, which is the median large galaxy masking radius (see Section 3.1.2). The difference in amplitude between the solid and dash-dotted curves, with a lower value when cross-correlating with the SGA-2020, suggests that the

catalogue of large galaxies contains either more low- z galaxies or more brighter galaxies, or both, compared to the BGS targets in $15 < r < 16$.

6 CONCLUSIONS

Here we have presented the steps needed to define and select the Bright Galaxy Survey (BGS) targets for the Dark Energy Spectroscopic Instrument (DESI) project. Our galaxy selection uses DECaLS LS imaging data from Data Release 8 (DR8) reduced by the NSF's OIR Lab CP and TRACTOR pipelines. Our BGS target selection has two main components, one which imposes spatial cuts and the other which applies photometric selections. Figs. 2 and 4 show the flowcharts that set out these two selections. At each step these flowcharts report the remaining survey area and surface density of targets.

The main features of our spatial and photometric cuts are the following:

- The BGS spatial target selection removes area near bright stars (BS mask), large galaxies (LG mask), and globular clusters (GC mask), as well as galaxies with less than a specified minimum number of observations (NOBS mask). The BS mask is a circular aperture that scales with the magnitude of the bright star (see Eqn. 3). The exclusion of areas around bright stars removes $\sim 270 \text{ deg}^2$, this is 2.76 per cent of initial footprint. Inspection of stacked images around bright stars (i.e. those with *Gaia* $G < 13$ or Tycho-2 $V < 13$) shows that the BS masking radius used in TRACTOR is well-motivated, with no sign of contamination around the bright stars in the BGS target density. There is a modest ~ 6 per cent increase in BGS target density just beyond the edge of the masked region. We find that there is a negligible angular cross-correlation between stars and galaxies at scales > 100 arcsec. Below 100 arcsec we have an anti-correlation possibly caused by the stars masked within the range $12 < G < 13$.
- The LG and GC masks account for a smaller number of contaminants than the BS mask, removing just $\sim 9 \text{ deg}^2$ of survey area or 0.09 percent of initial footprint.
- DECaLS DR8 is complete to 99.5 per cent with at least one observation in the three bands grz , as described by the value of NOBS. The selection made on NOBS removes $\sim 39 \text{ deg}^2$ of imaging data.
- We use *Gaia* DR2 to separate stars and galaxies as described in Section 4.1. This classification exploits the small PSF of the *Gaia* imaging compared with that typically present in ground-based observations. In our classification scheme we compare the measurement of the flux of an object by *Gaia* with that from TRACTOR through the parameter $G - rr$. Objects with a TRACTOR flux that is greater than that reported by *Gaia* are considered to be galaxies because this difference implies that they are extended sources (see Fig. 5).
- A small fraction (~ 0.35 per cent) of BGS galaxies are of PSF type according to TRACTOR. About half of these are compact sources for which the PSF model is the best fit, but the other half have only PSF photometry as they were designated stars based on the *Gaia* Astrometric Excess Noise (AEN) parameter before TRACTOR was run. For these objects TRACTOR only performs PSF fits. Matching to GAMA reveals that most (96 per cent) of these BGS PSF-type objects are confirmed to be galaxies by the GAMA spectroscopy. In addition, we find that the ~ 7 GAMA galaxies/deg² that are missed in BGS are mostly (~ 98 per cent) of PSF type according to TRACTOR. We conclude that using the

AEN classification is i) causing ~ 0.17 per cent of BGS galaxies to be falsely classified as of PSF type and ii) compromising the photometry of another 7 objects/deg² which then due to having their fluxes underestimated are falsely classified as stars by the BGS $G - rr$ star-galaxy classification.

- Possible systematic effects in DECaLS leave a small imprint on surface density of BGS sources. The variation in the target density of BGS sources as a function of the main possible systematic effects, such as the stellar density, galactic extinction, seeing and imaging depth, is less than 10 per cent in the case of stellar density and under 5 per cent for the remaining systematics. We implement a weighting scheme based on a linear regression model which uses the density of stars to mitigate these effects. Applying the resulting weights, variation in the target density with stellar density is removed by construction, and is greatly reduced when plotted against the other systematic quantities.

- Angular clustering measurements made from our BGS target catalogue are compared with previous measurements from SDSS and the predictions from the MXXL lightcone mock catalogue, which on large scales can be taken as a prediction of LCDM models (see 5.3). On small scales, the three measurements of the angular correlation function agree well, with the exception of the brightest galaxies considered. At large scales, the angular clustering we find for the BGS targets is closer to that recovered from the MXXL mock catalogue than the SDSS measurements. The agreement between the BGS and the MXXL lightcone is even better after applying the linear weights based on stellar density to the BGS.

Galleries with examples of BGS targets divided in BGS BRIGHT and BGS FAINT can be found at http://astro.dur.ac.uk/~qmxp55/bgs_ts_paper_gallery.html along with galleries showing examples of rejected objects by the different spatial and photometric cuts we apply in BGS. We included also examples of discrepancies between our star-galaxy (SG) classification using *Gaia* with TRACTORs divided into 1) TRACTORs extended objects that fail our SG classification, and the TRACTORs point sources objects that pass our SG classification and 2) are *Gaia* and 3) are not *Gaia* sources. Finally, examples of discrepancies between TRACTORs point source classification for *Gaia* objects and our SG classification divided in two samples: 1) are galaxies by our SG classification but stars according to TRACTORs assessment of *Gaia* sources using the Astrometric Excess Noise (AEN) parameter from *Gaia*, and 2) stars by our SG classification but galaxies by their AEN classification.

In a second paper we will focus on applying this framework to select BGS targets using the additional LS, BASS and MzLS imaging data, and set out what is needed to tune our selection to use the upcoming release of the LS, DR9. Among the main changes in DR9 compared to DR8 are i) the implementation of an iterative source detection process in TRACTOR in which the detection algorithm is rerun after sources have been fitted and subtracted, ii) an extended PSF model to subtract the wings of bright stars, iii) the COMPOSITE (COMP) TRACTOR model has been replaced with a SERSIC model, (SER) iv) the criteria used to determine which *Gaia* objects are forced to be fitted by the PSF model are now more restrictive, v) adjustments have been made to the masking procedure around bright stars and to fainter MEDIUM stars where the masking radius around bright stars has been reduced by a factor of two. In addition, TRACTOR implements a local fit to the sky background around these objects. vi) SGA-2020 and Globular Cluster catalogues have been updated and the large galaxy photometry redone in their own custom run of TRACTOR. It is expected that (i) will marginally

increase the completeness of the BGS catalogue, (iv) will reduce the incidence of galaxies being misclassified as stars, and the other changes will improve the photometry. A second paper will quantify these changes and focus predominately on changes in selection relative to this DR8 selection. Hence most of the details of the BGS selection will be in this paper only. Despite these improvements in the quality of the selection, early test releases of the upcoming DR9 data suggest that BGS targets will not vary more than 5 per cent compared to present selection with DECaLS DR8. The second paper will also include a more complete clustering analysis using mock catalogues and colour based clustering measurements, and a more sophisticated technique for the mitigation of systematic effects. A third paper we will cover the work we have undertaken to define and select the BGS targets for the survey validation programme. This series of papers is intended to be complementary work to the overall DESI key project paper on target selection aimed to be released in 2021.

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DATA AVAILABILITY

The data used in this paper is publicly available. The preliminary BGS target selection described in this paper is public at <https://data.desi.lbl.gov/public/ets/target/catalogs/> and detailed at <https://desidatamodel.readthedocs.io>. The DESI Legacy Imaging Surveys used for this work is public at <https://www.legacysurvey.org/dr8/description/>.

REFERENCES

- Abazajian K., et al., 2003, *Astron. J.*, 126, 2081
 Adelman-McCarthy J. K., et al., 2008, *ApJS*, 175, 297
 Alonso D., 2012, arXiv e-prints, p. arXiv:1210.1833
 Angulo R. E., Springel V., White S. D. M., Jenkins A., Baugh C. M., Frenk C. S., 2012, *MNRAS*, 426, 2046
 Bailer-Jones C. A. L., Fouesneau M., Andrae R., 2019, *MNRAS*, 490, 5615
 Baldry I. K., et al., 2010, *Monthly Notices of the Royal Astronomical Society*, 404, 86
 Baldry I. K., et al., 2017, *Monthly Notices of the Royal Astronomical Society*, 474, 3875
 Bertin E., Arnouts S., 1996, *A&AS*, 117, 393
 Blanton M. R., et al., 2001, *AJ*, 121, 2358
 Carrasco J. M., et al., 2016, *A&A*, 595, A7
 Chambers K. C., et al., 2016, arXiv e-prints, p. arXiv:1612.05560
 DESI Collaboration et al., 2016, arXiv e-prints, p. arXiv:1611.00036
 Dey A., et al., 2019, *The Astronomical Journal*, 157, 168
 Driver S. P., et al., 2012, VizieR Online Data Catalog, p. J/MNRAS/413/971
 Fadelly R., Hogg D. W., Willman B., 2012, *ApJ*, 760, 15
 Flaugher B., et al., 2015, *AJ*, 150, 150
 Gaia Collaboration et al., 2016a, *Astronomy and Astrophysics*, 595, A1
 Gaia Collaboration et al., 2016b, *Astronomy and Astrophysics*, 595, A2
 Gaia Collaboration et al., 2018, *Astronomy and Astrophysics*, 616, A1
 Høg E., et al., 2000, *A&A*, 355, L27
 Kitanidis E., et al., 2020, *Monthly Notices of the Royal Astronomical Society*, 496, 2262
 Lang D., Hogg D. W., Mykytyn D., 2016, The Tractor: Probabilistic astronomical source detection and measurement, Astrophysics Source Code Library (ascl:1604.008)
 Liske J., et al., 2015, *MNRAS*, 452, 2087
 Makarov D., Prugniel P., Terekhova N., Courtois H., Vauglin I., 2014, *A&A*, 570, A13
 Odewahn S. C., Stockwell E. B., Pennington R. L., Humphreys R. M., Zumach W. A., 1992, *AJ*, 103, 318
 Rezaie M., Seo H.-J., Ross A. J., Bunesco R. C., 2020, *MNRAS*, 495, 1613
 Ruiz-Macias O., et al., 2020, *Research Notes of the AAS*, 4, 187
 Schlafly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
 Schlegel D. J., Finkbeiner D. P., Davis M., 1998, *ApJ*, 500, 525
 Seibert N. J., Dudik R. P., Dorland B. N., Zacharias N., Makarov V., Fey A., Frouard J., Finch C., 2015, *ApJS*, 221, 12
 Smith A., Cole S., Baugh C., Zheng Z., Angulo R., Norberg P., Zehavi I., 2017, *Mon. Not. Roy. Astron. Soc.*, 470, 4646
 Strauss M. A., et al., 2002, *AJ*, 124, 1810
 The Dark Energy Survey Collaboration 2005, arXiv e-prints, pp astro-ph/0510346
 Wang Y., Brunner R. J., Dolence J. C., 2013, *Mon. Not. Roy. Astron. Soc.*, 432, 1961

Weir N., Fayyad U. M., Djorgovski S., 1995, *AJ*, 109, 2401
Zonca A., Singer L., Lenz D., Reinecke M., Rosset C., Hivon E., Gorski K., 2019, *Journal of Open Source Software*, 4, 1298

APPENDIX A: GALAXY VIEW

In contrast to the approach taken in the main paper, here we present a ‘galaxy’ view of the BGS selection by implementing the star-galaxy separation before the other BGS cuts (with the exception of first applying the nominal BGS magnitude limit $r < 20$). The results of this exercise are shown in Fig. A1. In this view, the geometric masking does not look as aggressive as it did in Fig. 2, with the size of the rejected area and number of objects typically reduced at each step by an order of magnitude compared to what was seen in Fig. 2. The BS mask step is the stage that is the most affected by this change in order. Next is the application of the selection on NOBS which has half the effect that it did in Fig. 2. Note that the area removed by the cuts remains unchanged as this does not depend on the number of targets but is calculated using the randoms.

In addition to the changing the order in which the star-galaxy separation is applied compared to the selection criteria presented in Sections 3 and 4, we swap the FMC and CCs with the QCs. When comparing both schemes, (Fig. 4 and Fig. A1), we see a high overlap between the QCs and the FMC of ~ 15 objects/deg² which represent $\sim 2/3$ the galaxies rejected by FMC in Section 4. CCs is also affected by the to the sequence of cuts and the rejections due to this cut is reduced by a factor of 2 in the galaxy view.

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.

BGS target selection - Galaxy View

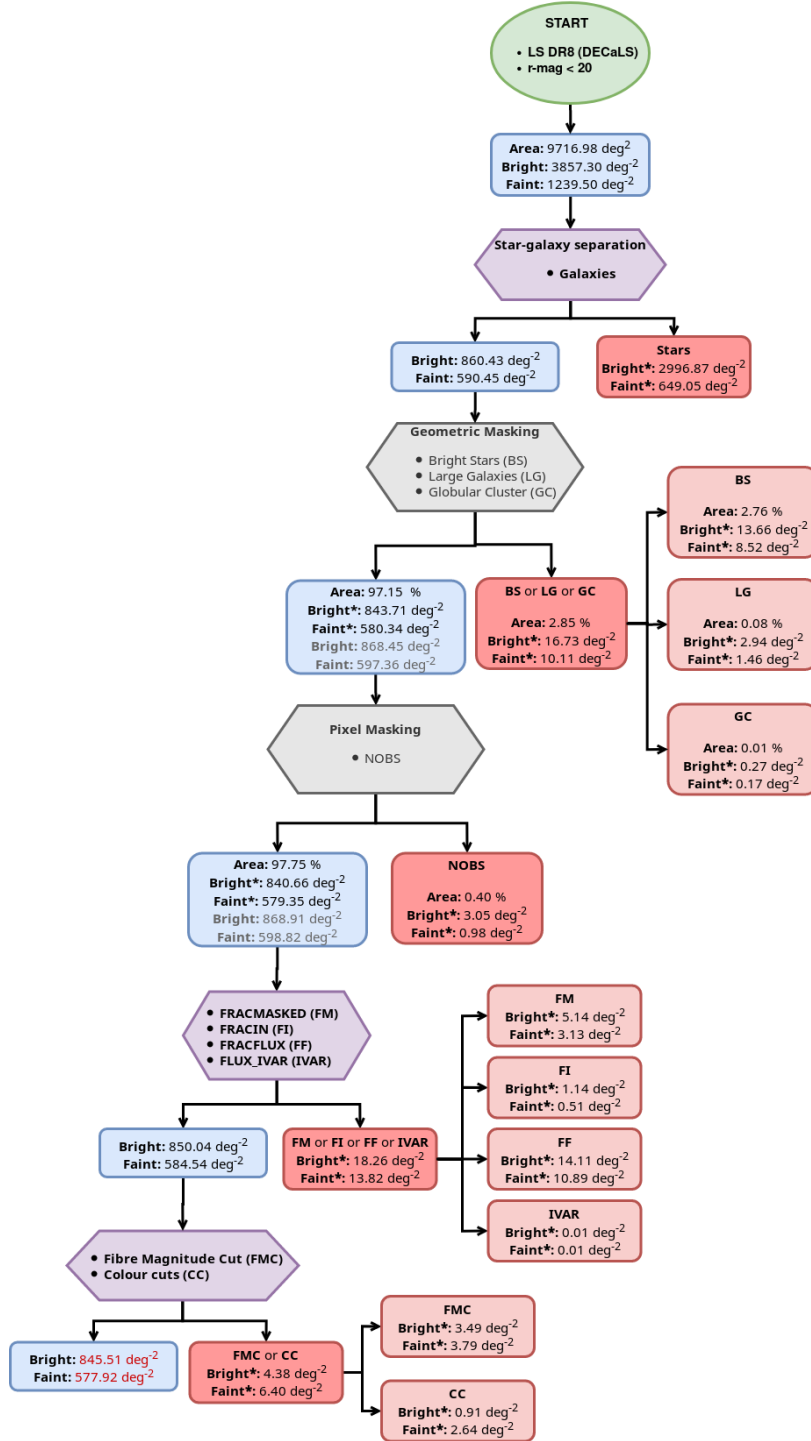


Figure A1. Flow chart showing the spatial and photometric BGS target selections applied to the Legacy Surveys DR8. The spatial selections are shown by gray boxes and are divided into two kinds, one defined by geometric cuts around bright sources i.e. bright stars (BS), large galaxies (LG) and globular clusters (GC), and the other which is at the pixel level, such as the number of observations (NOBS). The photometric selection of BGS targets is divided into four types and is shown by purple boxes; star-galaxy separation, fibre magnitude cuts (FMC), colour cuts (CCs) and quality cuts (QCs) which include FRACMASKED, FRACIN, FRACFLUX and FLUX_IVAR. The blue boxes show the area (in degrees) and the number density (per square degree) of objects retained after each selection, broken down into the numbers for the bright and faint components of the BGS. The red boxes show the equivalent information for the rejected objects. If more than one cut or selection is applied at a given stage, then the darker red boxes show the information about removed objects for the combination of cuts and the lighter red boxes show the corresponding values for each individual cut. The superscript (*) denotes target densities without correcting for the area removed by cuts up to that point, while densities without a superscript (*) do take into account the reduction in area.